

# Alternatives to Dark Matter and Dark Energy

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## Abstract

We review the underpinnings of the standard Newton-Einstein theory of gravity, and identify where it could possibly go wrong. In particular, we discuss the logical independence from each other of the general covariance principle, the equivalence principle and the Einstein equations, and discuss how to constrain the matter energy-momentum tensor which serves as the source of gravity. We identify the a priori assumption of the validity of standard gravity on all distance scales as the root cause of the dark matter and dark energy problems, and discuss how the freedom currently present in gravitational theory can enable us to construct candidate alternatives to the standard theory in which the dark matter and dark energy problems could then be resolved. We identify three generic aspects of these alternate approaches: that it is a universal acceleration scale which determines when a luminous Newtonian expectation is to fail to fit data, that there is a global cosmological effect on local galactic motions which can replace galactic dark matter, and that to solve the cosmological constant problem it is not necessary to quench the cosmological constant itself, but only the amount by which it gravitates.

## 1 Introduction

Following many years of research in cosmology and astrophysics, a picture of the universe has emerged [1, 2, 3, 4, 5] which is as troubling as it is impressive. Specifically, a wide variety of data currently support the view that the matter content of the universe consists of two primary components, viz. dark matter and dark energy, with ordinary luminous matter being relegated to a decidedly minor role. The nature and composition of these dark matter and dark energy components is not at all well-understood, and while both present severe challenges to the standard theory, each presents a different kind of challenge to it. As regards dark matter, there is nothing in principle wrong with the existence of non-luminous material per se (indeed objects such as dead stars, brown dwarfs and massive neutrinos are well-established in nature). Rather, what is disturbing is the ad hoc, after the fact, way in which dark matter is actually introduced, with its presence only being inferred after known luminous astrophysical sources are found to fail to account for any given astrophysical observation. Dark matter thus seems to know where, and in what amount, it is to be needed, and to know when it is not in fact needed (dark matter has to avoid being abundant in the solar system in order to not impair the success of standard gravity in accounting for solar system observations using visible sources alone); and moreover, in the cases where it is needed, what it is actually made of (astrophysical sources (Machos) or new elementary particles (Wimps)) is as yet totally unknown and elusive.

Disturbing as the dark matter problem is, the dark energy problem is even more severe, and not simply because its composition and nature is as mysterious as that of dark matter. Rather, for dark energy there actually is a very good, quite clear-cut candidate, viz. a cosmological constant, and the problem here is that the value for the cosmological constant as anticipated from fundamental theory is orders of magnitude larger than the data can possibly permit. With dark matter then, we see that luminous sources alone underaccount for the data, while for dark energy, a cosmological constant overaccounts for the data. Thus, within the standard picture, arbitrary as their introduction might be, there is nonetheless room for dark matter candidates should they ultimately be found, but for dark energy there is a need not to find something which might only momentarily be missing, but rather to get rid of something which is definitely there. And indeed, if it does not prove possible to quench the cosmological constant by the requisite orders and orders of magnitude, one would have to conclude that the prevailing cosmological theory simply does not work.

In arriving at the predicament that contemporary astrophysical and cosmological theory thus finds itself in, it is important to recognize that the entire case for the existence of dark matter and dark energy is based on just one thing alone, viz. on the validity on all distance scales of the standard Newton-Einstein gravitational theory as expressed through the Einstein equations of motion

$$-\frac{c^3}{8\pi G} \left( R^{\mu\nu} - \frac{1}{2} g^{\mu\nu} R^\alpha{}_\alpha \right) = T^{\mu\nu} \quad (1)$$

for the gravitational field  $g_{\mu\nu}$ . Specifically, the standard approach to cosmology and astrophysics is to take the Einstein equations of motion as given, and whenever the theory is found to encounter observational difficulties on any particular distance scale, modifications are to then be made to  $T^{\mu\nu}$  through the introduction of new, essentially ad hoc, gravitational sources so that agreement with observation can then be restored. While a better understanding of dark energy and explicit observational detection of dark matter sources might eventually be achieved in the future, at the present time the only apparent way to avoid the dark matter and dark energy problems is to modify or generalize not the right-hand side of Eq. (1) but rather its left. In order to see how one might actually do this, it is thus necessary to carefully go over the entire package represented by Eq. (1), to determine whether some of its ingredients might not be as secure as others, and investigate whether the weaker ones could possibly be replaced. To do this we will thus need to make a critical appraisal of both of the two sides of Eq. (1).

## 2 The underpinnings of the standard gravitational picture

### 2.1 Massive test particle motion

Following a year of remarkable achievement, achievement whose centennial is currently being celebrated, at the end of 1905 Einstein found himself in a somewhat paradoxical situation. While he had surmounted an enormous hurdle in developing special relativity, its very establishment created an even bigger hurdle for him. Specifically, with the development of special relativity Einstein resolved the conflict between the Lorentz invariance of the Maxwell equations and the Galilean invariance of Newtonian mechanics by realizing that it was Lorentz invariance which was the more basic of the two principles, and that it was Newtonian mechanics which therefore had to be modified. While special relativity thus ascribed primacy to Lorentz invariance so that observers moving with large or small uniform velocity could then all agree on the same physics, such observers still occupied a highly privileged position since uniform velocity observers form only a very small subset of all possible allowable observers, observers who could move with arbitrarily non-uniform velocity. Additionally, if it was special relativity that was to be the all-embracing principle, all interactions would then need to obey it, and yet what was at the time the accepted theory of gravity, viz. Newtonian gravity, in fact did not. It was the simultaneous resolution

of these two issues (viz. accelerating observers and the compatibility of gravity with relativity) at one and the same time through the spectacular development of general relativity which not only established the Einstein theory of gravity, but which left the impression that there was only one possible resolution to the two issues, viz. that based on Eq. (1). To pinpoint what it is in the standard theory which leads us to the dark matter and dark energy problems, we thus need to unravel the standard gravitational package into what are in fact logically independent components, an exercise which is actually of value in and of itself regardless in fact of the dark matter and dark energy problems.

In order to make such a dissection of the standard picture, we begin with a discussion of a standard, free, spinless, relativistic, classical-mechanical Newtonian particle of non-zero kinematic mass  $m$  moving in flat spacetime according to the special relativistic generalization of Newton's second law of motion

$$m \frac{d^2 \xi^\alpha}{d\tau^2} = 0 \quad , \quad R_{\mu\nu\sigma\tau} = 0 \quad , \quad (2)$$

where  $d\tau = (-\eta_{\alpha\beta} d\xi^\alpha d\xi^\beta)^{1/2}$  is the proper time and  $\eta_{\alpha\beta}$  is the flat spacetime metric, and where we have indicated explicitly that the Riemann tensor is (for the moment) zero. As such, Eq. (2) will be left invariant under linear transformations of the coordinates  $\xi^\mu$ , but on making an arbitrary non-linear transformation to coordinates  $x^\mu$  and using the definitions

$$\Gamma_{\mu\nu}^\lambda = \frac{\partial x^\lambda}{\partial \xi^\alpha} \frac{\partial^2 \xi^\alpha}{\partial x^\mu \partial x^\nu} \quad , \quad g_{\mu\nu} = \frac{\partial \xi^\alpha}{\partial x^\mu} \frac{\partial \xi^\beta}{\partial x^\nu} \eta_{\alpha\beta} \quad , \quad (3)$$

we directly find (see e.g. [6], a reference whose notation we use throughout) that the invariant proper time is brought to the form  $d\tau = (-g_{\mu\nu} dx^\mu dx^\nu)^{1/2}$ , and that the equation of motion of Eq. (2) is rewritten as

$$m \left( \frac{D^2 x^\lambda}{D\tau^2} \right) \equiv m \left( \frac{d^2 x^\lambda}{d\tau^2} + \Gamma_{\mu\nu}^\lambda \frac{dx^\mu}{d\tau} \frac{dx^\nu}{d\tau} \right) = 0 \quad , \quad R_{\mu\nu\sigma\tau} = 0 \quad , \quad (4)$$

with Eq. (4) serving to define  $D^2 x^\lambda / D\tau^2$ . As derived, Eq. (4) so far only holds in a strictly flat spacetime with zero Riemann curvature tensor, and indeed Eq. (4) is only a covariant rewriting of the special relativistic Newtonian second law of motion, i.e. it covariantly describes what an observer with a non-uniform velocity in flat spacetime sees, with the  $\Gamma_{\mu\nu}^\lambda$  term emerging as an inertial, coordinate-dependent force. The emergence of such a  $\Gamma_{\mu\nu}^\lambda$  term originates in the fact that even while the four-velocity  $dx^\lambda/d\tau$  transforms as a general contravariant vector, its ordinary derivative  $d^2 x^\lambda / d\tau^2$  (which samples adjacent points and not merely the point where the four-velocity itself is calculated) does not, and it is only the four-acceleration  $D^2 x^\lambda / D\tau^2$  which transforms as a general contravariant four-vector, and it is thus only this particular four-vector on whose meaning all (accelerating and non-accelerating) observers can agree. The quantity  $\Gamma_{\mu\nu}^\lambda$  is not itself a general coordinate tensor, and in flat spacetime one can eliminate it everywhere by working in Cartesian coordinates. Despite this privileged status for Cartesian coordinate systems, in general it is Eq. (4) rather than Eq. (2) which should be used (even in flat spacetime) since this is the form of Newton's second law of motion which an accelerating flat spacetime observer sees.

Now while all of the above remarks were developed purely for flat spacetime, Eq. (4) has an immediate generalization to curved spacetime where it then takes the form

$$m \left( \frac{D^2 x^\lambda}{D\tau^2} \right) \equiv m \left( \frac{d^2 x^\lambda}{d\tau^2} + \Gamma_{\mu\nu}^\lambda \frac{dx^\mu}{d\tau} \frac{dx^\nu}{d\tau} \right) = 0 \quad , \quad R_{\mu\nu\sigma\tau} \neq 0 \quad . \quad (5)$$

In curved spacetime, it is again only the quantity  $D^2 x^\lambda / D\tau^2$  which transforms as a general contravariant four-vector, and the Christoffel symbol  $\Gamma_{\mu\nu}^\lambda$  is again not a general coordinate tensor. Consequently at any given point  $P$  it can be made to vanish,<sup>1</sup> though no coordinate transformation can bring it to zero

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<sup>1</sup>Under the transformation  $x'^\lambda = x^\lambda + \frac{1}{2} x^\mu x^\nu (\Gamma_{\mu\nu}^\lambda)_P$ , the primed coordinate Christoffel symbols  $(\Gamma'_{\mu\nu}^\lambda)_P$  will vanish at the point  $P$ , regardless in fact of how large the Riemann tensor in the neighborhood of the point  $P$  might actually be.

at every point in a spacetime whose Riemann tensor is non-vanishing. As an equation, Eq. (5) can also be obtained from an action principle, since it is the stationarity condition  $\delta I_T / \delta x_\lambda = 0$  associated with functional variation of the test particle action

$$I_T = -mc \int d\tau \quad (6)$$

with respect to the coordinate  $x^\lambda$ . To appreciate the ubiquity of the appearance of the covariant acceleration  $D^2 x^\lambda / D\tau^2$ , we consider as action the curved space electromagnetic coupling

$$I_T^{(2)} = -mc \int d\tau + e \int d\tau \frac{dx^\lambda}{d\tau} A_\lambda \quad , \quad (7)$$

to find that its variation with respect to  $x^\lambda$  leads to the curved space Lorentz force law

$$mc \left( \frac{d^2 x^\lambda}{d\tau^2} + \Gamma_{\mu\nu}^\lambda \frac{dx^\mu}{d\tau} \frac{dx^\nu}{d\tau} \right) = e F^\lambda{}_\alpha \frac{dx^\alpha}{d\tau} \quad , \quad R_{\mu\nu\sigma\tau} \neq 0 \quad . \quad (8)$$

Similarly, the coupling of the test particle to the Ricci scalar via

$$I_T^{(3)} = -mc \int d\tau - \kappa \int d\tau R^\alpha{}_\alpha \quad , \quad (9)$$

leads to

$$(mc + \kappa R^\alpha{}_\alpha) \left( \frac{d^2 x^\lambda}{d\tau^2} + \Gamma_{\mu\nu}^\lambda \frac{dx^\mu}{d\tau} \frac{dx^\nu}{d\tau} \right) = -\kappa R^\alpha{}_{\alpha;\beta} \left( g^{\lambda\beta} + \frac{dx^\lambda}{d\tau} \frac{dx^\beta}{d\tau} \right) \quad , \quad (10)$$

( $R_{\mu\nu\sigma\tau}$  necessarily non-zero here), while the coupling of the test particle to a scalar field  $S(x)$  via

$$I_T^{(4)} = -\hat{\kappa} \int d\tau S(x) \quad , \quad (11)$$

leads to the curved space

$$\hat{\kappa} S \left( \frac{d^2 x^\lambda}{d\tau^2} + \Gamma_{\mu\nu}^\lambda \frac{dx^\mu}{d\tau} \frac{dx^\nu}{d\tau} \right) = -\hat{\kappa} S_{;\beta} \left( g^{\lambda\beta} + \frac{dx^\lambda}{d\tau} \frac{dx^\beta}{d\tau} \right) \quad , \quad R_{\mu\nu\sigma\tau} \neq 0 \quad , \quad (12)$$

an expression incidentally which reduces to Eq. (5) when  $S(x)$  is a spacetime constant (with the mass parameter then being given by  $mc = \hat{\kappa} S$ ). In all of the above cases then it is the quantity  $D^2 x^\lambda / D\tau^2$  which must appear, since in each such case the action which is varied is a general coordinate scalar.

## 2.2 The equivalence principle

Now as such, the analysis given above is a purely kinematic one which discusses only the propagation of test particles in curved backgrounds. This analysis makes no reference to gravity per se, and in particular makes no reference to Eq. (1) at all, though it does imply that in any curved spacetime in which the metric  $g_{\mu\nu}$  is taken to be the gravitational field, covariant equations of motion involving four-accelerations would have to be based strictly on  $D^2 x^\lambda / D\tau^2$ , with neither  $d^2 x^\lambda / d\tau^2$  nor  $\Gamma_{\mu\nu}^\lambda (dx^\mu / d\tau)(dx^\nu / d\tau)$  having any coordinate independent significance or meaning. As such, we thus recognize the equivalence principle as the statement that the  $d^2 x^\lambda / d\tau^2$  and  $\Gamma_{\mu\nu}^\lambda (dx^\mu / d\tau)(dx^\nu / d\tau)$  terms must appear in all of the above propagation equations in precisely the combination indicated,<sup>2</sup> and that via a coordinate transformation it is possible to remove the Christoffel symbols at any chosen point  $P$ , with it thereby being possible to

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<sup>2</sup>This thereby secures the equality of the inertial and passive gravitational masses of material particles.

simulate the Christoffel symbol contribution to the gravitational field at such a point  $P$  by an accelerating coordinate system in flat spacetime. We introduce this particular formulation of the equivalence principle quite guardedly, since in an equation such as the fully covariant Eq. (10), one cannot remove the dependence on the Ricci scalar by any coordinate transformation whatsoever, and we thus define the equivalence principle not as the statement that all gravitational effects at any point  $P$  can be simulated by an accelerating coordinate system in flat spacetime (viz. that test particles unambiguously move on geodesics and obey Eq. (5) and none other), but rather that no matter what propagation equation is to be used for test particles, the appropriate acceleration for them is  $D^2x^\lambda/D\tau^2$ .<sup>3</sup> With Eq. (4) thus showing the role of coordinate invariance in flat spacetime, with Eqs. (5), (8), and (10) exhibiting the equivalence principle in curved spacetime, and with all of these equations being independent of the Einstein equations of Eq. (1), the logical independence of the general covariance principle, the equivalence principle and the Einstein equations is thus established. Hence, any metric theory of gravity in which the action is a general coordinate scalar and the metric is the gravitational field will thus automatically obey both the relativity principle and the equivalence principle, no matter whether or not the Einstein equations are to be imposed as well.<sup>4</sup>

### 2.3 Massless field motion

Absent from the above discussion is the question of whether real, as opposed to test, particles actually obey curved space propagation equations such as Eq. (5) at all, and whether such a discussion should apply to massless particles as well since for them both  $m$  and the proper time  $d\tau$  vanish identically, with equations such as Eq. (5) becoming meaningless. Now it turns out that these two issues are not actually independent, as both reduce to the question of how waves rather than particles couple to gravity, since light is described by a wave equation, and elementary particles are actually taken to be the quanta associated with quantized fields which are also described by wave equations. And since the discussion will be of relevance for the exploration of the energy-momentum tensor to be given below, we present it in some detail now. For simplicity we look at the standard minimally coupled curved spacetime massless Klein-Gordon scalar field with wave equation

$$S^{;\mu}_{;\mu} = 0 \quad , \quad R_{\mu\nu\sigma\tau} \neq 0 \quad (13)$$

where  $S^{;\mu}$  denotes the contravariant derivative  $\partial S/\partial x_\mu$ . If for the scalar field we introduce a eikonal phase  $T(x)$  via  $S(x) = \exp(iT(x))$ , the scalar  $T(x)$  is then found to obey the equation

$$T^{;\mu}T_{;\mu} - iT^{;\mu}_{;\mu} = 0 \quad , \quad (14)$$

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<sup>3</sup>This particular formulation of the equivalence principle does no violence to observation, since Eotvos experiment type testing of the equivalence principle is made in Ricci-flat Schwarzschild geometries where all Ricci tensor or Ricci scalar dependent terms (such as for instance those exhibited in Eq. (10)) are simply absent, with such tests (and in fact any which involve the Schwarzschild geometry) not being able to distinguish between Eqs. (4) and (10).

<sup>4</sup>To sharpen this point, we note that it could have been the case that the resolution of the conflict between gravity and special relativity could have been through the introduction of the gravitational force not as a geometric entity at all, but rather as an analog of the way the Lorentz force is introduced in Eq. (8). In such a case, in an accelerating coordinate system one would still need to use the acceleration  $D^2x^\lambda/D\tau^2$  and not the ordinary  $d^2x^\lambda/d\tau^2$ . However, if the gravitational field were to be treated the same way as the electromagnetic field, left open would then be the issue of whether physics is to be conducted in flat space or curved space, i.e. left open would be the question of what does then fix the Riemann tensor. There would then have to be some additional equation which would fix it, and curvature would still have to be recognized as having true, non-coordinate artifact effects on particles if the Riemann tensor were then found to be non-zero. Taking such curvature to be associated with the gravitational field (rather than with some further field) is of course the most economical, though doing so would not oblige gravitational effects to only be felt through  $D^2x^\lambda/D\tau^2$ , and would not preclude some gravitational Lorentz force type term (such as the one exhibited in Eq. (10)) from appearing as well.

an equation which reduces to  $T^{;\mu}T_{;\mu} = 0$  in the short wavelength limit. From the relation  $T^{;\mu}T_{;\mu;\nu} = 0$  which then ensues, it follows that in the short wavelength limit the phase  $T(x)$  obeys

$$T^{;\mu}T_{;\mu;\nu} = T^{;\mu}[T_{;\nu;\mu} + \partial_\nu T_{;\mu} - \partial_\mu T_{;\nu}] = T^{;\mu}[T_{;\nu;\mu} + \partial_\nu \partial_\mu T - \partial_\mu \partial_\nu T] = T^{;\mu}T_{;\nu;\mu} = 0 \quad . \quad (15)$$

Since normals to wavefronts obey the eikonal relation

$$T^{;\mu} = \frac{dx^\mu}{dq} = k^\mu \quad (16)$$

where  $q$  is a convenient scalar affine parameter which measures distance along the normals and  $k^\mu$  is the wave vector of the wave, on noting that  $(dx^\mu/dq)(\partial/\partial x^\mu) = d/dq$  we thus obtain

$$k^\mu k^\lambda_{;\mu} = \frac{d^2 x^\lambda}{dq^2} + \Gamma^\lambda_{\mu\nu} \frac{dx^\mu}{dq} \frac{dx^\nu}{dq} = 0 \quad , \quad (17)$$

a condition which we recognize as being the massless particle geodesic equation, with rays then precisely being found to be geodesic in the eikonal limit. Since the discussion given earlier of the coordinate dependence of the Christoffel symbols was purely geometric, we thus see that once rays such as light rays are geodesic, they immediately obey the the equivalence principle,<sup>5</sup> with phenomena such as the gravitational bending of light then immediately following.

Now while we do obtain strict geodesic motion for rays when we eikonalize the minimally coupled Klein-Gordon equation, the situation becomes somewhat different if we consider a non-minimally coupled Klein-Gordon equation instead. Thus, on replacing Eq. (13) by

$$S^{;\mu}_{;\mu} + \frac{\xi}{6} S R^\alpha_\alpha = 0 \quad , \quad R_{\mu\nu\sigma\tau} \neq 0 \quad , \quad (18)$$

the curved space eikonal equation then takes the form

$$T^{;\mu}T_{;\mu} - \frac{\xi}{6} R^\alpha_\alpha = 0 \quad , \quad (19)$$

with Eq. (17) being replaced by the Ricci scalar dependent

$$\frac{d^2 x^\lambda}{dq^2} + \Gamma^\lambda_{\mu\nu} \frac{dx^\mu}{dq} \frac{dx^\nu}{dq} = \frac{\xi}{12} (R^\alpha_\alpha)^{;\lambda} \quad (20)$$

in the short wavelength limit.

A similar dependence on the Ricci scalar or tensor is also obtained for massless spin one-half and spin one fields. Specifically, even though the massless Dirac equation in curved space, viz.  $i\gamma^\mu(x)[\partial_\mu + \Gamma_\mu]\psi(x) = 0$  (where  $\Gamma_\mu(x)$  is the fermion spin connection) contains no explicit direct dependence on the Ricci tensor, nonetheless the second order differential equation which the fermion field then also obeys is found to take the form  $[\partial_\mu + \Gamma_\mu][\partial^\mu + \Gamma^\mu]\psi(x) + (1/4)R^\alpha_\alpha\psi(x) = 0$ . Likewise, even though the curved space Maxwell equations, viz.  $F^{\mu\nu}_{;\nu} = 0$ ,  $F_{\mu\nu;\lambda} + F_{\lambda\mu;\nu} + F_{\nu\lambda;\mu} = 0$  also possess no direct coupling to the Ricci tensor, manipulation of the Maxwell equations leads to the second order  $g^{\alpha\beta}F^{\mu\nu}_{;\alpha;\beta} + F^{\mu\alpha}R^\nu_\alpha - F^{\nu\alpha}R^\mu_\alpha = 0$ , and thus to the second order equation  $g^{\alpha\beta}A_{\mu;\alpha;\beta} - A^\alpha_{;\alpha;\mu} + A^\alpha R_{\mu\alpha} = 0$  for the vector potential  $A^\mu$ . Characteristic of all these massless field equations then is the emergence of an explicit dependence on the Ricci tensor, and thus of some non-geodesic motion analogous to that exhibited in Eq. (20) when we eikonalize.<sup>6</sup> Now equations such as Eq. (20) will still obey the

<sup>5</sup>According to Eq. (17), an observer in Einstein's elevator would not be able to tell if a light ray is falling downwards under gravity or whether the elevator is accelerating upwards.

<sup>6</sup>We note that none of these particular curved space field equations involve the Riemann tensor, but only the Ricci tensor. This is fortunate since Schwarzschild geometry tests of gravity would be sensitive to the Riemann tensor, a tensor which in contrast to the Ricci tensor, does not vanish in a Schwarzschild geometry.

equivalence principle (as we have defined it above), since the four-acceleration which appears is still one which contains the non-tensor Christoffel symbols. Moreover, despite the fact that all of these particular curved space equations are non-geodesic, nonetheless, in the flat space limit all of them degenerate into the flat space geodesics, with the eikonal rays travelling on straight lines. Now in general, what is understood by covariantizing is to replace flat spacetime expressions by their curved space counterparts, with Eq. (4) for instance being replaced by Eq. (5). However, this is a very restrictive procedure, since Eqs. (5) and (10) both have the same flat space limit. The standard covariantizing prescription will thus fail to generate any terms which explicitly depend on curvature (terms which in some cases we see must be there), and thus constructing a curved space energy-momentum tensor purely by covariantizing a flat space one is not at all a general prescription, an issue we shall return to below when we discuss the curved space energy-momentum tensor in detail.

## 2.4 Massive field motion

A situation analogous to the above also obtains for massive fields. Specifically for the quantum-mechanical minimally coupled curved space massive Klein-Gordon equation, viz.

$$S^{;\mu}_{;\mu} - \frac{m^2 c^2}{\hbar^2} S = 0 \quad , \quad R_{\mu\nu\sigma\tau} \neq 0 \quad , \quad (21)$$

the substitution  $S(x) = \exp(iP(x)/\hbar)$  yields

$$P^{;\mu} P_{;\mu} + m^2 c^2 = i\hbar P^{;\mu}_{;\mu} \quad . \quad (22)$$

In the eikonal (or the small  $\hbar$ ) approximation the  $i\hbar P^{;\mu}_{;\mu}$  term can be dropped, so that the phase  $P(x)$  is then seen to obey the purely classical condition

$$g_{\mu\nu} P^{;\mu} P^{;\nu} + m^2 c^2 = 0 \quad . \quad (23)$$

We immediately recognize Eq. (23) as the covariant Hamilton-Jacobi equation of classical mechanics, an equation whose solution is none other than the stationary classical action  $\int p_\mu dx^\mu$  as evaluated between relevant end points. In the eikonal approximation then we can thus identify the wave phase  $P(x)$  as  $\int p_\mu dx^\mu$ , with the phase derivative  $P^{;\mu} = \partial^\mu P$  then being given as the particle momentum  $p^\mu = mcdx^\mu/d\tau$ , a four-vector momentum which accordingly has to obey

$$g_{\mu\nu} p^\mu p^\nu + m^2 c^2 = 0 \quad , \quad (24)$$

the familiar fully covariant particle energy-momentum relation. With covariant differentiation of Eq. (24) immediately leading to the classical massive particle geodesic equation

$$p^\mu p^\lambda_{;\mu} = \frac{d^2 x^\lambda}{d\tau^2} + \Gamma^\lambda_{\mu\nu} \frac{dx^\mu}{d\tau} \frac{dx^\nu}{d\tau} = 0 \quad (25)$$

(as obtained here with the proper time  $d\tau$  appropriate to massive particles and not the affine parameter  $q$ ), we thus recover the well known result that the center of a quantum-mechanical wave packet follows the stationary classical trajectory. Further, since we may also reexpress the stationary  $\int p_\mu dx^\mu$  as  $-mc \int d\tau$ , we see that we can also identify the quantum-mechanical eikonal phase as  $P(x) = -mc \int d\tau$ , to thus enable us to make contact with the  $I_T$  action given in Eq. (6).<sup>7</sup> Though we have thus made contact with  $I_T$ , it is important to realize that we were only able to arrive at Eq. (23) after having started with the equation of motion of Eq. (21), an equation whose own validity requires that stationary

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<sup>7</sup>We make contact with the  $I_T$  of Eq. (6) since we start with the minimally coupled Eq. (21).

variation of the Klein-Gordon action from which it is derived had already been made, with only the stationary classical action actually being a solution to the Hamilton-Jacobi equation. The action  $I_T = -mc \int d\tau$  as evaluated along the stationary classical path is thus a part of the solution to the wave equation, i.e. the output, rather than a part of the input.<sup>8</sup> Thus in the quantum-mechanical case we never need to assume as input the existence of any point particle action such as  $I_T$  at all. Rather, we need only assume the existence of equations such as the standard Klein-Gordon equation, with eikonalization then precisely putting particles onto classical geodesics just as desired. To conclude then, we see that not only does the equivalence principle hold for light (even though it has no inertial mass or gravitational mass at all) and hold for quantum-mechanical particles, we also see that the equivalence principle need not be intimately tied to the classical test particle action  $I_T$  at all. Given this analysis, we turn now to consideration of the curved space energy-momentum tensor.

## 3 Structure of the energy-momentum tensor

### 3.1 Kinematic perfect fluids

At the time of the development of general relativity, the prevailing view of gravitational sources was that they were to be treated like billiard balls, i.e. as purely mechanical kinematic particles which carry energy and momentum; with the advent of general relativity then requiring that such energy and momentum be treated covariantly, so that the way the energy-momentum tensor was to be introduced in gravitational theory was to simply covariantize the appropriate flat spacetime expressions. Despite the subsequent realization that particles are very far from being such kinematic objects (elementary particles are now thought to get their masses dynamically via spontaneous symmetry breaking), and despite the fact that the now standard  $SU(3) \times SU(2) \times U(1)$  theory of strong, electromagnetic and weak interactions ascribes primacy to fields over particles, the kinematically prescribed energy-momentum tensor (rather than an  $SU(3) \times SU(2) \times U(1)$  based one) is nonetheless still used in treatments of gravity today. Apart from this already disturbing shortcoming (one we shall remedy below by constructing the energy-momentum tensor starting from fields rather than particles), an additional deficiency of a purely kinematic prescription is that since flat spacetime energy-momentum tensors had no need to know where the zero of energy was (in flat spacetime only energy differences are observable), their covariantizing left unidentified where the zero of energy might actually be; and since gravity couples to energy itself rather than to energy differences, this kinematic prescription is thus powerless to address the cosmological constant problem, an issue to which we shall return below.

Historically there was of course good reason to treat particles kinematically, since such a treatment did lead to geodesic motion for the particles. Thus for the test particle action  $I_T$  of Eq. (6), its functional variation with respect to the metric allows one to define an energy-momentum tensor according to<sup>9</sup>

$$\frac{2}{(-g)^{1/2}} \frac{\delta I_T}{\delta g_{\mu\nu}} = T^{\mu\nu} = \frac{mc}{(-g)^{1/2}} \int d\tau \delta^4(x - y(\tau)) \frac{dy^\mu}{d\tau} \frac{dy^\nu}{d\tau} \quad , \quad (26)$$

with its covariant conservation ( $T^{\mu\nu}{}_{;\nu} = 0$ ) leading right back to the geodesic equation of Eq. (5).<sup>10</sup> Moreover, an analogous situation is met when the energy-momentum tensor is taken to be a perfect

<sup>8</sup>Thus we cannot appeal to  $I_T$  to put particles on geodesics, since we already had to put them on the geodesics which follow from Eq. (23) in order to get to  $I_T$  in the first place. While the use of an action such as  $I_T$  will suffice to obtain geodesic motion, as we thus see, its use is not at all necessary, with it being eikonalization of the quantum-mechanical wave equation which actually puts massive particles on geodesics.

<sup>9</sup>With our definition here and throughout of  $T^{\mu\nu}$  as  $T^{\mu\nu} = 2(\delta I_T / \delta g_{\mu\nu}) / (-g)^{1/2}$ , it is  $cT_{00}$  which then has the dimension of an energy density rather than  $T_{00}$  itself.

<sup>10</sup>While our ability to impose a conservation condition on  $T^{\mu\nu}$  would follow from Eq. (1) since the Einstein tensor  $G^{\mu\nu} = R^{\mu\nu} - (1/2)g^{\mu\nu}R^\alpha{}_\alpha$  obeys the Bianchi identity, the use of a conservation condition in no way requires the validity of Eq. (1). Specifically, in any covariant theory of gravity in which the pure gravitational piece of the action,  $I_{GRAV}$ ,

fluid of the form

$$T^{\mu\nu} = \frac{1}{c} [(\rho + p)U^\mu U^\nu + pg^{\mu\nu}] \quad (27)$$

with energy density  $\rho$ , pressure  $p$  and a fluid four-vector which is normalized to  $U_\mu U^\mu = -1$ . Specifically, for such a fluid, covariant conservation leads to

$$[(\rho + p)U^\mu U^\nu + pg^{\mu\nu}]_{;\nu} = [(\rho + p)U^\nu]_{;\nu} U^\mu + (\rho + p)U^\mu_{;\nu} U^\nu + p^{;\mu} = 0 \quad , \quad (28)$$

and thus to

$$- [(\rho + p)U^\nu]_{;\nu} + U_\mu p^{;\mu} = -(\rho + p)U^\mu_{;\nu} U^\nu U_\mu = 0 \quad , \quad (29)$$

with the last equality following since  $U_\mu U^\mu_{;\nu} = 0$ . With the insertion of Eq. (29) into (28) then yielding

$$(\rho + p)U^\mu_{;\nu} U^\nu + p_{;\nu} [g^{\mu\nu} + U^\mu U^\nu] = 0 \quad , \quad (30)$$

viz.

$$\frac{D^2 x^\mu}{D\tau^2} = -[g^{\mu\nu} + U^\mu U^\nu] \frac{p_{;\nu}}{(\rho + p)} \quad , \quad (31)$$

we see that the geodesic equation then emerges whenever the right-hand side of Eq. (31) is negligible, a situation which is for instance met when the fluid is composed of non-interacting pressureless dust.

Since Eqs. (26) and (31) do lead to geodesic motion, it is generally thought that gravitational sources should thus be described in this way. However, even within an a priori kinematic perfect fluid framework, Eq. (31) would in no way need to be altered if to the perfect fluid of Eq. (27) we were to add an additional  $T_{EXTRA}^{\mu\nu}$  which was itself independently covariantly conserved. Since a geometric tensor such as the Einstein tensor  $G^{\mu\nu}$  is both covariantly conserved and non-existent in flat spacetime, a curved space generalization of a flat space energy-momentum tensor which would include a term of the form  $T_{EXTRA}^{\mu\nu} = G^{\mu\nu}$  would not affect Eq. (31) at all. Moreover, since the metric tensor  $g^{\mu\nu}$  is also covariantly conserved, the inclusion in  $T_{EXTRA}^{\mu\nu}$  of a  $\Lambda g^{\mu\nu}$  term with  $\Lambda$  constant would also leave Eq. (31) untouched, and while such a term would not vanish in the flat spacetime limit, its presence in flat spacetime would only lead to a non-observable overall shift in the zero of energy. Restricting to the perfect fluid of Eq. (27) is thus only sufficient to recover Eq. (31), and not at all necessary.

Within the kinematic perfect fluid framework, the use of such fluids as gravitational sources is greatly facilitated if some equation of state of the form  $p = w\rho$  can be prescribed where  $w$  would be a constant. To see what choices are suggested for  $w$  from flat space physics, we consider a relativistic flat space ideal  $N$  particle classical gas of particles of mass  $m$  in a volume  $V$  at a temperature  $T$ . For this system the Helmholtz free energy  $A(V, T)$  is given as

$$e^{-A(V,T)/NkT} = V \int d^3p e^{-(p^2+m^2)^{1/2}/kT} \quad , \quad (32)$$

so that the pressure takes the simple form

$$P = - \left( \frac{\partial A}{\partial V} \right)_T = \frac{NkT}{V} \quad , \quad (33)$$

while the internal energy  $U = \rho V$  evaluates in terms of Bessel functions as

$$U = A - T \left( \frac{\partial A}{\partial T} \right)_V = 3NkT + Nm \frac{K_1(m/kT)}{K_2(m/kT)} \quad . \quad (34)$$

---

is a general coordinate scalar function of the metric, the quantity  $A^{\mu\nu} = (2/(-g)^{1/2})(\delta I_{GRAV}/\delta g_{\mu\nu})$  will, because of covariance, automatically obey  $A^{\mu\nu}_{;\nu} = 0$ , and through the gravitational equation of motion  $A^{\mu\nu} = T^{\mu\nu}$  then lead to the covariant conservation of  $T^{\mu\nu}$ . The use of the Einstein equations is only sufficient to yield  $T^{\mu\nu}_{;\nu} = 0$ , but not at all necessary, with the conservation ensuing in any general covariant pure metric theory of gravity.

In the high and low temperature limits (the radiation and matter eras) we then find that the expression for  $U$  simplifies to

$$\begin{aligned} \frac{U}{V} &\rightarrow \frac{Nm}{V} + \frac{3NkT}{2V} = \frac{Nm}{V} + \frac{3P}{2} \approx \frac{Nm}{V} & , & \quad \frac{m}{kT} \rightarrow 0 & , \\ \frac{U}{V} &\rightarrow \frac{Nm}{V} + \frac{3NkT}{2V} = \frac{Nm}{V} + \frac{3P}{2} \approx \frac{Nm}{V} & , & \quad \frac{m}{kT} \rightarrow \infty & . \end{aligned} \quad (35)$$

Consequently, while  $p$  and  $\rho$  are nicely proportional to each other in the high temperature radiation and the low temperature matter eras (where  $w(T \rightarrow \infty) = 1/3$  and  $w(T \rightarrow 0) = 0$ ), we also see that in transition region between the two eras their relationship is altogether more complicated. Use of a  $p = w\rho$  equation of state would at best only be valid at temperatures which are very different from those of order  $m/K$ , though for massless particles it would be of course be valid to use  $p = \rho/3$  at all temperatures, a point to which we shall return below.

### 3.2 Perfect Robertson-Walker fluids

In trying to develop equations of state in curved space, one should replace the partition function in Eq. (32) by its curved space generalization (i.e. one should covariantize it, just as is proposed for  $T^{\mu\nu}$  itself),<sup>11</sup> and then follow the steps above to see what generalization of Eq. (35) might then ensue. However for curved backgrounds of high symmetry, the use of the isometry structure of the background can greatly simplify the discussion. Thus, for instance, for the Robertson-Walker geometry of relevance to cosmology, viz.

$$ds^2 = -c^2 dt^2 + R^2(t) \left( \frac{dr^2}{(1-kr^2)} + r^2 d\Omega_2 \right) , \quad (36)$$

the maximal 3-symmetry of the background entails that any rank two tensor such as the energy-momentum tensor itself must have the generic form

$$T^{\mu\nu} = [C(t) + D(t)]U^\mu U^\nu + D(t)g^{\mu\nu} , \quad (37)$$

to thus automatically be of a perfect fluid form with comoving fluid 4-vector  $U^\mu = (1, 0, 0, 0)$  and general functions  $C$  and  $D$  which can only depend on the comoving time  $t$ . With the energy-momentum tensor being covariantly conserved,  $C$  and  $D$  have to obey

$$\frac{d}{dt} \left( R^3(C + D) \right) = R^3 \frac{dD}{dt} , \quad (38)$$

with  $D$  and  $C$  thus being related according to

$$D = -\frac{d}{dR^3} \left( R^3 C \right) . \quad (39)$$

While  $D$  and  $C$  must be also directly proportional to each other in the Robertson-Walker case since each only depends on the single parameter  $t$ , nonetheless, even though we therefore can set  $D/C = w$  in such cases, in general the quantity  $w$  could still be a function of  $t$  and need not necessarily be a constant.<sup>12</sup> In the above we have purposefully not identified  $C$  with a fluid  $\rho$  and  $D$  with a fluid  $p$ , since should the energy-momentum tensor actually consist of two types of fluid (this being the conventional dark matter plus dark energy picture), even if both of them have their own independent  $w$  according

<sup>11</sup>Typically one replaces  $(p^2 + m^2)^{1/2}/kT$  by  $(dx^\mu/d\tau)U^\nu g_{\mu\nu}/kT$  [ $U^\nu$  is a four-vector] and replaces  $\int d^3p$  by a sum over a complete set of basis modes associated with the propagation of a spinless massive particle in the chosen  $g_{\mu\nu}$  background.

<sup>12</sup>When  $w$  is a constant, we can then set  $C = 1/R^{3(w+1)}$ .

to  $p_1 = w_1\rho_1$ ,  $p_2 = w_2\rho_2$ , the sum of their pressures would then obey  $p_1 + p_2 = w_1\rho_1 + w_2\rho_2$  and would not in general be proportional to  $\rho_1 + \rho_2$ , with the total  $\rho_1 + \rho_2$  and  $p_1 + p_2$  obeying

$$p_1 + p_2 = -\frac{d}{dR^3} \left( R^3(\rho_1 + \rho_2) \right) \quad , \quad (40)$$

and with neither  $\rho_1 + \rho_2$  or  $p_1 + p_2$  then scaling as a power of  $R$ .<sup>13</sup> However, if the full energy-momentum tensor is to describe radiation fluids alone, then no matter how many of them there might actually be in total, in such a situation the full  $T^{\mu\nu}$  must additionally obey the tracelessness condition  $T^\mu{}_\mu = 0$ , to then unambiguously fix  $w = D/C$  to the unique value  $w = 1/3$ .

### 3.3 Schwarzschild fluid sources

While the isometry of the Robertson-Walker geometry does automatically lead us to the perfect fluid form given in Eq. (37) (though not necessarily to any particular equation of state unless  $T^\mu{}_\mu = 0$ ), the situation for lower symmetry backgrounds is not as straightforward. Specifically, for a standard static, spherically symmetric geometry of the form

$$ds^2 = -B(r)c^2dt^2 + A(r)dr^2 + r^2d\Omega_2 \quad , \quad (41)$$

the three Killing vector symmetry of the background entails that the most general energy-momentum tensor must have the generic diagonal form

$$T_{00} = \rho(r)B(r) \quad , \quad T_{rr} = p(r)A(r) \quad , \quad T_{\theta\theta} = q(r)r^2 \quad , \quad T_{\phi\phi} = q(r)r^2\sin^2\theta \quad (42)$$

with conservation condition

$$\frac{dp}{dr} + \frac{(\rho + p)}{2B} \frac{dB}{dr} + \frac{2}{r}(p - q) = 0 \quad , \quad (43)$$

with the function  $q(r)$  not at all being required to be equal to  $p(r)$ .<sup>14</sup> Thus while a flat space static, spherically symmetric perfect fluid would have  $p = q$  and equation of state  $p = w\rho$ , it does not follow that a curved space one would as well, with this being a dynamical and not a kinematic issue whose resolution would require an evaluation of the covariant partition function in the background of Eq. (41).<sup>15</sup>

Despite these considerations, it turns out that a quite a bit of the standard phenomenology associated with static, spherically symmetric sources still holds even if  $q$  and  $p$  are quite different from each other. Specifically, for the geometry of Eq. (41) the components of the Ricci tensor obey

$$\frac{R_{00}}{2B} + \frac{R_{rr}}{2A} + \frac{R_{\theta\theta}}{r^2} = \frac{1}{r^2} \left[ \frac{d}{dr} \left( \frac{r}{A} \right) - 1 \right] \quad , \quad R_{\theta\theta} = -1 + \frac{r}{2A} \left( \frac{1}{B} \frac{dB}{dr} - \frac{1}{A} \frac{dA}{dr} \right) + \frac{1}{A} \quad , \quad (44)$$

while the components of the energy-momentum tensor obey

$$\left( T_{00} - \frac{g_{00}T^\alpha{}_\alpha}{2} \right) = \frac{B}{2}(\rho + p + 2q) \quad ,$$

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<sup>13</sup>As well as needing to require that  $w_1$  and  $w_2$  both be constant, to secure the conventional  $\rho_1 = 1/R^{3(w_1+1)}$ ,  $\rho_2 = 1/R^{3(w_2+1)}$  in the two fluid case additionally requires the separate covariant conservation of the energy-momentum tensor of each fluid, so that the two fluids do not then exchange energy and momentum with each other. A dynamics which is to secure this would have to be identified in any cosmological model (such as a dark matter plus quintessence fluid model) which uses two such fluids unless one of the two fluids just happens to have  $w = -1$ , since this would correspond to a cosmological constant whose energy momentum tensor  $T^{\mu\nu} = -\Lambda g^{\mu\nu}$  is in fact independently conserved.

<sup>14</sup>Unlike the Robertson-Walker case, a static, spherically symmetric geometry is only spherically symmetric about a single point and not about all points in the spacetime.

<sup>15</sup>Since  $q$  would be equal to  $p$  in the flat space limit, one might anticipate that for weak gravity the difference between  $q$  and  $p$  would still be small, though  $q$  could differ radically from  $p$  in the strong gravity black hole limit.

$$\begin{aligned}
\left(T_{rr} - \frac{g_{rr}T_{\alpha}^{\alpha}}{2}\right) &= \frac{A}{2}(\rho + p - 2q) \ , \\
\left(T_{\theta\theta} - \frac{g_{\theta\theta}T_{\alpha}^{\alpha}}{2}\right) &= \frac{r^2}{2}(\rho - p) \ , \\
\frac{1}{2B}\left(T_{00} - \frac{1}{2}g_{00}T_{\alpha}^{\alpha}\right) + \frac{1}{2A}\left(T_{rr} - \frac{1}{2}g_{rr}T_{\alpha}^{\alpha}\right) + \frac{1}{r^2}\left(T_{\theta\theta} - \frac{1}{2}g_{\theta\theta}T_{\alpha}^{\alpha}\right) &= \rho \ ,
\end{aligned} \tag{45}$$

with the last expression conveniently being independent of  $q$ . If we now do impose the Einstein equations of Eq. (1), then in terms of the quantity

$$\tilde{M}(r) = 4\pi \int_0^r dr r^2 \rho(r) \tag{46}$$

we immediately obtain

$$A^{-1}(r) = 1 - \frac{2G}{r}\tilde{M}(r) \ , \tag{47}$$

to then find that  $B(r)$  obeys

$$\frac{1}{B} \frac{dB}{dr} = \frac{2G}{r(r - 2G\tilde{M})} (\tilde{M} + 4\pi r^3 p) \ . \tag{48}$$

We recognize Eqs. (47) and (48) as being of precisely the same form as the standard expressions which are obtained (see e.g. [6]) for  $A$  and  $(1/B)dB/dr$  when  $q(r)$  is equal to  $p(r)$ , though the substitution of these expressions into Eq. (43) leads to

$$\frac{dp}{dr} + \frac{(\rho + p)G}{r(r - 2G\tilde{M})} (\tilde{M} + 4\pi r^3 p) + \frac{2}{r}(p - q) = 0 \ , \tag{49}$$

an equation which differs from the standard expression by the presence of the  $2(p - q)/r$  term. If the matter density terminates at some finite  $r = R$ , then outside of the fluid the geometry is a standard exterior Schwarzschild geometry with metric

$$B(r > R) = A^{-1}(r > R) = 1 - \frac{2MG}{r} \ . \tag{50}$$

Matching this exterior solution to the interior solution at  $r = R$  then yields for  $M$  the standard

$$M = 4\pi \int_0^R dr r^2 \rho(r) \ , \tag{51}$$

with the integration constant required for Eq. (48) then being fixed to yield for  $B(r)$  the standard expression

$$B(r) = \exp\left(-2G \int_r^{\infty} dr \frac{(\tilde{M} + 4\pi r^3 p)}{r(r - 2G\tilde{M})}\right) \ . \tag{52}$$

As we thus see, the functional forms of Eqs. (47), (51) and (52) are completely unaffected by whether or not  $q$  is equal to  $p$ , with any difference between  $q$  and  $p$  only showing up in Eq. (49). As far as the geometry outside of the fluid is concerned, the structure of the exterior Schwarzschild metric is the standard one with the standard form for the total mass  $M$  as given in Eq. (51). However, inside the fluid the dynamics could be different from the conventional treatment because of modifications to the equation of state. However, for weak gravity where  $p$  is already of order  $G$ , should the term of order  $G$  in  $dp/dr$  be given exactly by  $G\rho\tilde{M}/r^2$ , the quantity  $p - q$  would then only begin in order  $G^2$  and the standard lowest order in  $G$  hydrostatic treatment of sources such as stars would not be affected. Nonetheless, even if that is to be the case (something which is not immediately clear), for strong gravity

inside sources (where there are currently no data), any difference between  $p$  and  $q$  could have substantial consequences, and so we should not in general expect a static, spherically symmetric source to possess an energy-momentum tensor of the form  $T^{\mu\nu} = [\rho(r) + p(r)]U^\mu U^\nu + p(r)g^{\mu\nu}$  just because it does so in flat space.<sup>16</sup>

### 3.4 Scalar field fluid sources

As we have thus seen, wisdom gained from experience with kinematic particle sources in flat space serves as a quite limited guide to the structure of gravitational sources in curved spacetime. However, that is not their only shortcoming, with their connection to the structure of the energy-momentum tensor which is suggested by fundamental theory being quite remote. Thus we need to discuss what is to be expected of the energy-momentum tensor in a theory in which the action is built out of fields rather than particles, and in which the fields develop masses by spontaneous symmetry breakdown. However, before going to the issue of dynamical masses, we first need to see how we can connect our above analysis of geodesic motion of eikonalized fields to the structure of the energy-momentum tensor.

To illustrate what is involved, it is convenient to consider a massive complex flat spacetime scalar field, and to get its energy-momentum tensor (and to subsequently enforce a tracelessness condition for it when we restrict to massless fields), we take as action the non-minimally coupled curved space action<sup>17</sup>

$$I_M = - \int d^4x (-g)^{1/2} \left[ \frac{1}{2} S^{;\mu} S^*_{;\mu} + \frac{1}{2} m^2 S S^* - \frac{\xi}{12} S S^* R^\mu{}_\mu \right] . \quad (53)$$

Its variation with respect to the scalar field yields the equation of motion

$$S^{;\mu}{}_{;\mu} + \frac{\xi}{6} S R^\mu{}_\mu - m^2 S = 0 , \quad (54)$$

while its variation with respect the metric yields the energy-momentum tensor

$$\begin{aligned} T_{\mu\nu} = & \left( \frac{1}{2} - \frac{\xi}{6} \right) (S_{;\mu} S^*_{;\nu} + S_{;\nu} S^*_{;\mu}) - \frac{(3-2\xi)}{6} g_{\mu\nu} S^{;\alpha} S^*_{;\alpha} - \frac{\xi}{6} (S S^*_{;\mu;\nu} + S^* S_{;\mu;\nu}) \\ & + \frac{\xi}{6} g_{\mu\nu} (S^* S^{;\alpha}{}_{;\alpha} + S S^{*;\alpha}{}_{;\alpha}) - \frac{1}{2} g_{\mu\nu} m^2 S S^* - \frac{\xi}{6} S S^* \left( R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} R^\alpha{}_\alpha \right) . \end{aligned} \quad (55)$$

With the trace of this energy-momentum tensor evaluating to

$$T^\mu{}_\mu = (\xi - 1) \left( S^{;\mu} S^*_{;\mu} + \frac{1}{2} S^* S^{;\mu}{}_{;\mu} + \frac{1}{2} S S^{*;\mu}{}_{;\mu} \right) - m^2 S S^* \quad (56)$$

in field configurations which obey Eq. (54), we see that the choice  $\xi = 1$ ,  $m = 0$  will enforce the tracelessness of  $T^{\mu\nu}$ . Bearing this in mind we thus set  $\xi = 1$ ,<sup>18</sup> so that the flat space limit of the  $\xi = 1$  theory then takes the form

$$T_{\mu\nu} = \frac{1}{3} (\partial_\mu S \partial_\nu S^* + \partial_\nu S \partial_\mu S^*) - \frac{1}{6} g_{\mu\nu} \partial_\alpha S \partial^\alpha S^* - \frac{1}{6} (S \partial_\mu \partial_\nu S^* + S^* \partial_\mu \partial_\nu S)$$

<sup>16</sup>In passing we additionally note that once  $q$  is not equal  $p$ , for a traceless fluid with  $T^\mu{}_\mu = -\rho + p + 2q = 0$ , Eq. (43) reduces to  $dp/dr + (\rho + p)/2B)(dB/dr) + (3p - \rho)/r = 0$ , and dependent on how  $p$  depends on  $\rho$ , it may be possible for radiation to support a static, stable source. (When  $\rho = 3p$  the only solution is  $p \sim 1/B^2$  which would require  $B(r)$  to be singular at the point  $r = R$  where  $p$  vanishes.)

<sup>17</sup>To construct the correct energy-momentum tensor in flat space, it is necessary to first vary the curved space action with respect to the metric and then take the flat limit.

<sup>18</sup>When  $\xi = 1$ , the coupling of the scalar field to the geometry is conformal, with the massless action  $I_M = - \int d^4x (-g)^{1/2} [S^{;\mu} S^*_{;\mu}/2 - S S^* R^\mu{}_\mu/12]$  being invariant under the local conformal transformation  $S(x) \rightarrow e^{-\alpha(x)} S(x)$ ,  $g_{\mu\nu}(x) \rightarrow e^{2\alpha(x)} g_{\mu\nu}(x)$ .

$$\begin{aligned}
& +\frac{1}{6}g_{\mu\nu}(S^*\partial_\alpha\partial^\alpha S + S\partial_\alpha\partial^\alpha S^*) - \frac{1}{2}g_{\mu\nu}m^2SS^* \\
= & \frac{1}{3}(\partial_\mu S\partial_\nu S^* + \partial_\nu S\partial_\mu S^*) - \frac{1}{6}g_{\mu\nu}\partial_\alpha S\partial^\alpha S^* - \frac{1}{6}(S\partial_\mu\partial_\nu S^* + S^*\partial_\mu\partial_\nu S) - \frac{1}{6}g_{\mu\nu}m^2SS^* . \quad (57)
\end{aligned}$$

In a plane wave solution to the  $\partial_\mu\partial^\mu S = m^2S$  wave equation of the form  $S(x) = e^{ik\cdot x}/V^{1/2}(E_k)^{1/2}$  where  $k^\mu k_\mu = -m^2$ ,  $E_k = (k^2 + m^2)^{1/2}$  and  $V$  is the 3-volume,  $T_{\mu\nu}$  then readily evaluates to

$$T_{\mu\nu} = \frac{k_\mu k_\nu}{VE_k} , \quad (58)$$

and even though the wavefront to the massive plane wave is geodesic, this particular energy-momentum tensor does not look anything like a perfect fluid form, since for a single massive plane wave  $k^\mu = (E_k, 0, 0, k)$  propagating geodetically in the  $z$ -direction  $T_{\mu\nu}$  evaluates to

$$T_{\mu\nu} = \frac{1}{V} \begin{pmatrix} E_k & 0 & 0 & -k \\ 0 & 0 & 0 & 0 \\ 0 & 0 & 0 & 0 \\ -k & 0 & 0 & k^2/E_k \end{pmatrix} . \quad (59)$$

To get a perfect fluid we need to incoherently add an entire family of these plane waves.<sup>19</sup> Thus if we take a set of six plane waves moving in the  $\pm x$ ,  $\pm y$ ,  $\pm z$  directions, all with the same  $k = |\vec{k}|$  and  $E_k$ , viz.  $k^\mu = (E_k, k, 0, 0)$ ,  $k^\mu = (E_k, -k, 0, 0)$ ,  $k^\mu = (E_k, 0, k, 0)$ ,  $k^\mu = (E_k, 0, -k, 0)$ ,  $k^\mu = (E_k, 0, 0, k)$ , and  $k^\mu = (E_k, 0, 0, -k)$ , and then incoherently add up their individual contributions to  $T_{\mu\nu}$ , we obtain

$$T_{00} = \frac{6E_k}{V} , \quad T_{xx} = T_{yy} = T_{zz} = \frac{2k^2}{E_k V} , \quad T^\mu{}_\mu = -\frac{6m^2}{E_k V} . \quad (60)$$

We recognize Eq. (60) as being of precisely the perfect fluid form with  $\rho = 6E_k/V$ ,  $p = 2k^2/E_k V$ . As such, we not only get a perfect fluid form, we even see that in the event that the mass is introduced kinematically as in the action of Eq. (53), then for small  $k \ll m$  we obtain  $p \ll \rho$ , with the effective  $p/\rho = w$  being zero, while for very large  $k$  (or equivalently for zero  $m$ ) we obtain  $p = \rho/3$ . As such the above procedure shows how to obtain a perfect fluid form starting from field theory, with its generalization to curved space requiring an equivalent incoherent averaging over the basis modes associated with the curved space wave equations appropriate to the curved space backgrounds of interest, a procedure which we indicated earlier is not guaranteed to automatically yield the straightforward  $T_{xx} = T_{yy} = T_{zz}$  condition found in flat space.

### 3.5 Dynamical fluid sources

While the above analysis shows how the perfect fluid form for massive particles can be obtained from field theory, and shows why the test particle action and energy-momentum tensor of Eqs. (6) and (26) are not at all germane to the issue, this analysis still fails to take into consideration the fact that elementary particle masses are not kinematic, but rather that they are acquired dynamically by spontaneous breakdown. To investigate what is to happen in the dynamical mass case, it is convenient to consider a spin one-half matter field fermion  $\psi(x)$  which is to get its mass through a real spin-zero Higgs scalar boson field  $S(x)$ . In order to illustrate the difference between dynamic and kinematic masses in the sharpest way possible, we take the action to possess no intrinsic mass scales and the energy-momentum tensor to be traceless. We thus give neither the fermion nor the scalar field any

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<sup>19</sup>This is equivalent to using (the zero temperature limit of) the partition function discussed earlier.

kinematic mass at all, and in order to secure the tracelessness of the energy-momentum tensor couple the scalar field conformally to gravity, to thus yield as curved space matter action<sup>20</sup>

$$I_M = - \int d^4x (-g)^{1/2} \left[ \frac{1}{2} S^{;\mu} S_{;\mu} - \frac{1}{12} S^2 R^\mu{}_\mu + \lambda S^4 + i\bar{\psi}\gamma^\mu(x)[\partial_\mu + \Gamma_\mu(x)]\psi - hS\bar{\psi}\psi \right] \quad (61)$$

where  $h$  and  $\lambda$  are dimensionless coupling constants.<sup>21</sup> Variation of this action with respect to  $\psi(x)$  and  $S(x)$  yields the equations of motion

$$i\gamma^\mu(x)[\partial_\mu + \Gamma_\mu(x)]\psi - hS\psi = 0 \quad , \quad (62)$$

and

$$S^{;\mu}{}_{;\mu} + \frac{1}{6} S R^\mu{}_\mu - 4\lambda S^3 + h\bar{\psi}\psi = 0 \quad , \quad (63)$$

while variation with respect to the metric yields (without use of any equation of motion) an energy-momentum tensor of the form

$$\begin{aligned} T_{\mu\nu} = & i\bar{\psi}\gamma_\mu(x)[\partial_\nu + \Gamma_\nu(x)]\psi + \frac{2}{3} S_{;\mu} S_{;\nu} - \frac{1}{6} g_{\mu\nu} S^{;\alpha} S_{;\alpha} - \frac{1}{3} S S_{;\mu;\nu} + \frac{1}{3} g_{\mu\nu} S S^{;\alpha}{}_{;\alpha} \\ & - \frac{1}{6} S^2 \left( R_{\mu\nu} - \frac{1}{2} g_{\mu\nu} R^\alpha{}_\alpha \right) - g_{\mu\nu} \left[ \lambda S^4 + i\bar{\psi}\gamma^\alpha(x)[\partial_\alpha + \Gamma_\alpha(x)]\psi - hS\bar{\psi}\psi \right] \quad , \end{aligned} \quad (64)$$

with use of the matter field equations of motion then permitting us to rewrite  $T_{\mu\nu}$  as

$$\begin{aligned} T_{\mu\nu} = & i\bar{\psi}\gamma_\mu(x)[\partial_\nu + \Gamma_\nu(x)]\psi + \frac{2}{3} S_{;\mu} S_{;\nu} - \frac{1}{6} g_{\mu\nu} S^{;\alpha} S_{;\alpha} - \frac{1}{3} S S_{;\mu;\nu} \\ & + \frac{1}{12} g_{\mu\nu} S S^{;\alpha}{}_{;\alpha} - \frac{1}{6} S^2 \left( R_{\mu\nu} - \frac{1}{4} g_{\mu\nu} R^\alpha{}_\alpha \right) - \frac{1}{4} g_{\mu\nu} h S \bar{\psi}\psi \quad . \end{aligned} \quad (65)$$

Additional use of the matter field equations of motion then confirms that this energy-momentum tensor is indeed traceless.

In the presence of a spontaneously broken non-zero constant expectation value  $S_0$  for the scalar field, the energy-momentum tensor is then found to simplify to

$$T_{\mu\nu} = i\bar{\psi}\gamma_\mu(x)[\partial_\nu + \Gamma_\nu(x)]\psi - \frac{1}{4} g_{\mu\nu} h S_0 \bar{\psi}\psi - \frac{1}{6} S_0^2 \left( R_{\mu\nu} - \frac{1}{4} g_{\mu\nu} R^\alpha{}_\alpha \right) \quad , \quad (66)$$

with flat space limit

$$T_{\mu\nu} = i\bar{\psi}\gamma_\mu \partial_\nu \psi - \frac{1}{4} \eta_{\mu\nu} h S_0 \bar{\psi}\psi \quad . \quad (67)$$

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<sup>20</sup>In Eq. (61) the general-relativistic Dirac matrices  $\gamma^\mu(x)$  and the fermion spin connection  $\Gamma_\mu(x)$  are defined as  $\gamma^\mu(x) = V_a^\mu(x)\hat{\gamma}^a$  and  $\Gamma_\mu = [\gamma^\nu(x), \partial_\mu \gamma_\nu(x)]/8 - [\gamma^\nu(x), \gamma_\sigma(x)]\Gamma_{\mu\nu}^\sigma/8$  where  $V_a^\mu(x)$  is a vierbein and  $\hat{\gamma}^a$  is a special-relativistic Dirac gamma matrix which, with the  $ds^2 = -(dx^0)^2 + (dx^1)^2 + (dx^2)^2 + (dx^3)^2 = \eta_{ab} dx^a dx^b$  metric, obeys  $\hat{\gamma}^a \hat{\gamma}^b + \hat{\gamma}^b \hat{\gamma}^a = -2\eta^{ab}$ , while  $\bar{\psi}$  is given by  $\bar{\psi} = \psi^\dagger \hat{D}$  where  $\hat{D}$  is a Hermitian flat spacetime Dirac matrix which effects  $\hat{D}\hat{\gamma}^a\hat{D}^{-1} = \hat{\gamma}^{a\dagger}$ . In our notation, here and throughout Hermiticity in the fermion kinetic energy sector is implicit, with  $i\bar{\psi}\gamma^\mu(x)[\partial_\mu + \Gamma_\mu(x)]\psi$  denoting  $(i/2)\bar{\psi}\gamma^\mu(x)[\partial_\mu + \Gamma_\mu(x)]\psi - (i/2)\overleftarrow{\bar{\psi}}[\overleftarrow{\partial}_\mu + \Gamma_\mu(x)]\gamma^\mu(x)\psi = (i/2)\bar{\psi}\gamma^\mu(x)[\partial_\mu + \Gamma_\mu(x)]\psi + \text{Hermitian conjugate}$ , and  $i\bar{\psi}\gamma_\mu(x)[\partial_\nu + \Gamma_\nu(x)]\psi$  denoting  $(i/4)\bar{\psi}\gamma_\mu(x)[\partial_\nu + \Gamma_\nu(x)]\psi + (i/4)\overleftarrow{\bar{\psi}}\gamma_\mu(x)[\partial_\nu + \Gamma_\nu(x)]\psi + \text{Hermitian conjugate}$  in the associated energy-momentum tensor constructed below as the symmetric  $T_{\mu\nu} = 2g_{\mu\alpha}g_{\nu\beta}(\delta I_M/\delta g_{\alpha\beta})/(-g)^{1/2} = -(1/2)[(V_{\mu a}(x)/|V(x)|)(\delta I_M/\delta V_a^\nu(x)) + (V_{\nu a}(x)/|V(x)|)(\delta I_M/\delta V_a^\mu(x))]$  where  $|V(x)| = (-g)^{1/2}$ . In this construction we note that  $\delta g_{\mu\nu} = V_{\mu a}\delta V_\nu^a + \delta V_{\mu a}V_\nu^a$ ,  $\delta|V(x)| = -|V(x)|V_a^\mu(x)\delta V_a^\mu(x)$ , and  $\delta V_{\sigma c} = -V_{\mu c}V_\sigma^a\delta V_a^\mu$ , to thereby yield  $(V_{\nu a}(x)/|V(x)|)(\delta/\delta V_a^\mu(x))[f d^4x (-g)^{1/2}\bar{\psi}\gamma^\sigma(x)[\partial_\sigma + \Gamma_\sigma(x)]\psi] = \bar{\psi}\gamma_\nu(x)[\partial_\mu + \Gamma_\mu(x)]\psi + (1/8)\bar{\psi}\gamma^\sigma(x)([\gamma_\nu, \gamma_\mu])_{;\sigma}\psi - g_{\nu\mu}\bar{\psi}\gamma^\sigma(x)[\partial_\sigma + \Gamma_\sigma(x)]\psi$ , from which Eq. (64) then follows. (The author wishes to thank Dr. R. K. Nesbet for helpful comments on the material presented here.)

<sup>21</sup>As such, the action of Eq. (61) is the most general curved space matter action for the  $\psi(x)$  and  $S(x)$  fields which is invariant under the local conformal transformation  $S(x) \rightarrow e^{-\alpha(x)}S(x)$ ,  $\psi(x) \rightarrow e^{-3\alpha(x)/2}\psi(x)$ ,  $\bar{\psi}(x) \rightarrow e^{-3\alpha(x)/2}\bar{\psi}(x)$ ,  $g_{\mu\nu}(x) \rightarrow e^{2\alpha(x)}g_{\mu\nu}(x)$ .

With the fermion now obeying

$$i\gamma^\mu \partial_\mu \psi - hS_0 \psi = 0 \quad , \quad (68)$$

the tracelessness of the flat space energy-momentum tensor of Eq. (67) is manifest. The quantization of the flat space limit theory is straightforward since Eq. (68) just describes a free fermion with mass  $m = hS_0$ , and yields one particle plane wave eigenstates  $|k\rangle$  of four momentum  $k^\mu = (E_k, \vec{k})$  where  $E_k = (k^2 + m^2)^{1/2}$ . For a positive energy, spin-up, Dirac spinor propagating in the  $z$ -direction, in analog to Eq. (59) we obtain the matrix elements

$$T_{\mu\nu} = \frac{1}{V} \begin{pmatrix} E_k & 0 & 0 & -k \\ 0 & 0 & 0 & 0 \\ 0 & 0 & 0 & 0 \\ -k & 0 & 0 & k^2/E_k \end{pmatrix} + \frac{1}{V} \begin{pmatrix} -m^2/4E_k & 0 & 0 & 0 \\ 0 & m^2/4E_k & 0 & 0 \\ 0 & 0 & m^2/4E_k & 0 \\ 0 & 0 & 0 & m^2/4E_k \end{pmatrix} . \quad (69)$$

with trace

$$T^\mu{}_\mu = \frac{1}{V} \left[ -E_k + \frac{k^2}{E_k} + \frac{m^2}{E_k} \right] = 0 \quad . \quad (70)$$

In Eq. (69) we recognize a two-component structure to the fermion energy-momentum tensor, a standard kinematic piece in which  $\langle k|T_{00}|k\rangle = E_k/V$ ,  $\langle k|T_{33}|k\rangle = k^2/E_k V$ , and a dynamic part coming from the symmetry breaking in which  $\langle k|T_{00}|k\rangle = -m^2/4E_k$ ,  $\langle k|T_{11}|k\rangle = \langle k|T_{22}|k\rangle = \langle k|T_{33}|k\rangle = m^2/4E_k$ . On incoherently averaging over the directions of  $\vec{k}$ , the energy momentum tensor is then found to take the form

$$T_{\mu\nu} = (\rho + p)U_\mu U_\nu + p\eta_{\mu\nu} + \Lambda\eta_{\mu\nu} \quad , \quad (71)$$

where

$$\rho = \frac{6E_k}{V} \quad , \quad p = \frac{2k^2}{VE_k} \quad , \quad \Lambda = \frac{3m^2}{2VE_k} \quad , \quad (72)$$

and where the tracelessness of  $T^{\mu\nu}$  is enforced by the relation

$$3p - \rho + 4\Lambda = 0 \quad . \quad (73)$$

As an energy-momentum tensor, Eq. (71) provides an explicit example of a phenomenon we referred to earlier, namely that it is possible to add on to a kinematic energy-momentum tensor an additional tensor which is itself covariantly conserved without affecting the covariant conservation of the kinematic energy-momentum tensor itself. As well as showing how a cosmological constant type term  $\Lambda$  can naturally arise in dynamical mass theories, the great virtue of using Eq. (71) is that the tracelessness condition of Eq. (73) constrains the value of  $\Lambda$  to be neither smaller nor larger than  $\rho - 3p$ , an issue which we shall revisit below when we discuss the dark energy problem, with the tracelessness condition thus being seen to give us control of the cosmological constant.

It is instructive to compare this dynamical theory of fermion masses with a strictly kinematic fermion mass theory in which the action is given by

$$I_M = - \int d^4x (-g)^{1/2} \left[ i\bar{\psi}\gamma^\mu(x) [\partial_\mu + \Gamma_\mu(x)]\psi - m\bar{\psi}\psi \right] \quad , \quad (74)$$

with the flat space fermion wave equation being given by

$$i\gamma^\mu \partial_\mu \psi - m\psi = 0 \quad , \quad (75)$$

the flat space energy-momentum tensor being given by

$$T_{\mu\nu} = i\bar{\psi}\gamma_\mu \partial_\nu \psi \quad , \quad (76)$$

and the trace being given by the non-zero

$$T^\mu{}_\mu = m\bar{\psi}\psi . \quad (77)$$

This time, an incoherent averaging over plane wave states yields

$$T_{\mu\nu} = (\rho + p)U_\mu U_\nu + p\eta_{\mu\nu} , \quad \rho = \frac{6E_k}{V} , \quad p = \frac{2k^2}{VE_k} , \quad (78)$$

with trace  $T^\mu{}_\mu = 3p - \rho \neq 0$ . With both Eq. (71) and Eq. (78) leading to the covariant conservation of the kinematic  $(\rho + p)U_\mu U_\nu + p\eta_{\mu\nu}$ , we see that both energy-momentum tensors lead to eikonalized geodesic motion, with the validity of the equivalence principle in no way requiring the use of Eq. (78).<sup>22</sup> The key distinction between Eqs. (71) and (78) is that Eq. (71) contains not just the energy on the fermion field, but also that of the Higgs field that gave it its mass, an energy which couples to gravity. In dynamical theories of mass generation, gravity is thus sensitive to the Higgs field associated with the fermion, while in kinematic case it of course is not.

### 3.6 Implications of elementary particle physics

While it is conventional to use Eq. (78) and its analogs in current standard model cosmological studies, it is actually Eq. (71) and its analogs which is suggested by elementary particle physics. Specifically, in the standard  $SU(3) \times SU(2) \times U(1)$  theory of strong, electromagnetic and weak interactions, the only mass scale present in the fundamental Lagrangian is the wrong sign, tachyonic mass term in the Higgs potential  $V(S) = \lambda S^4 - \mu^2 S^2/2$  with all fermions and gauge bosons getting their masses dynamically through a non-vanishing Higgs field expectation value, and all couplings between fields being dimensionless. As regards the Higgs field, it could be a fundamental field with a bona fide fundamental tachyonic mass, or it could be a dynamical manifestation of an underlying symmetry breaking through bilinear fermion condensates, with  $S(x)$  then only being a Ginzburg-Landau type long range order parameter (and  $V(S)$  its effective Ginzburg-Landau action) which arises when the fermion bilinear takes a non-zero expectation value in a spontaneously broken vacuum. In such a case the underlying theory of fermions and gauge bosons would possess no intrinsic mass scales at all and would then have a traceless energy-momentum tensor. Beyond these two possibilities for the Higgs field, inspection of the action of Eq. (61) reveals yet another origin for the wrong sign mass term, namely that it could even arise from curvature in a theory with no fundamental mass scale at all, with  $R^\mu{}_\mu$  serving as  $\mu^2$  in a non-flat background. Thus either the energy-momentum tensor of  $SU(3) \times SU(2) \times U(1)$  is traceless (dynamical Higgs) or its trace is given (fundamental Higgs) by the non-zero  $T^\mu{}_\mu = \mu^2 S_0^2$  (c.f. Eq. (56)) when  $S$  is constant. At best then the trace can only be related to the Higgs field itself, and would thus not be given by the  $T^\mu{}_\mu = 3p - \rho$  expected of the kinematic fluid of Eq. (78). The use of the perfect fluids commonly employed in standard gravity and cosmology would thus appear to be at variance with the standard  $SU(3) \times SU(2) \times U(1)$  model of particle physics.<sup>23</sup> Hence, even before we

<sup>22</sup>Analogously, if we take the scalar  $S(x)$  to be a constant in Eq. (12), Eq. (12) then reduces to Eq. (5), with test particle motion being geodesic whether mass is kinematic or dynamic.

<sup>23</sup>In the standard particle physics model radiative corrections will lead to the generation of a trace anomaly in theories whose starting Lagrangian has no fundamental mass scales at all. However, such anomalies while then making the trace of the renormalized energy-momentum tensor non-zero, do not themselves introduce any new mass scale (the dimension four trace is proportional to terms which are quadratic in the curvature tensor) and would not in any way make the trace equal to the kinematic  $3p - \rho$  required of Eq. (78). Moreover, it is possible to actually cancel the trace anomaly altogether, by a judicious choice of fields, a judicious choice of geometric background, or by a renormalization group fixed point at which the coefficient of the trace anomaly is then zero. That the trace anomaly is to be cancelled would appear to be part of the standard cosmological wisdom anyway, since no such anomaly term contribution is ever included alongside the kinematic ordinary matter, kinematic dark matter and dark energy components which make up the energy content of the standard cosmological model, and so we shall not consider the trace anomaly any further here.

enter into the issue of the fact that the use of standard gravity in astrophysics and cosmology leads to the dark matter and dark energy problems, we see that the very use of such dark matter and dark energy sources at all in the kinematic way in which they are commonly used is already not in accord with the standard model of particle physics. With this word of caution in mind we shall now discuss the status of standard gravity with its conventional kinematic perfect fluid sources, and in looking for departures from standard gravity then allow for theories in which the exact energy-momentum tensor is traceless. Since the use of a traceless energy-momentum tensor in the Einstein equations of Eq. (1) would lead to a Ricci scalar which would have to vanish in every conceivable situation (to rapidly then bring standard gravity into conflict with data), when we do consider the source of gravity on the right-hand side of the gravitational equations of motion to be traceless, we will have to modify the left-hand side as well in a way which would make it automatically traceless as well. As we shall see, such a departure from the standard theory will readily allow us to resolve both the dark matter and the dark energy problems. Having now explored constraints on the right-hand side of Eq. (1) in some detail, we turn next to an exploration of its left-hand side, and begin first with the Newtonian gravity which was its antecedent.

## 4 Newtonian gravity

### 4.1 The Newtonian potential

The Newtonian prescription for determining the non-relativistic potential  $\phi(\vec{r})$  at any point  $\vec{r}$  due to a set of static, mass sources,  $m_i$ , at points  $\vec{r}_i$  is to sum over them according to

$$\phi(\vec{r}) = - \sum_i^N \frac{m_i G}{|\vec{r} - \vec{r}_i|} \quad (79)$$

where  $G$  is Newton's constant, with the motions of material test particles then being determined via

$$\frac{d^2 \vec{r}}{dt^2} = -\vec{\nabla} \phi \quad . \quad (80)$$

As such, Eq. (79) contains the full content of Newton's law of gravity, and for any given set of gravitational sources, any candidate theory of gravity needs to recover the associated  $\phi(\vec{r})$  in any kinematic region in which Eq. (79) has been confirmed, to the level of precision required by available data. To determine what are the appropriate kinematic regions requires first testing Newtonian predictions in a candidate region using sources which are already known and prescribed in advance of data, and after success has been achieved, and only after, is one then free to use the law again in that same region to infer the existence of other, previously unknown, sources. At the present time one can say that in this way Newton's law of gravity has been well-established on distance scales from the order of millimeters (the smallest distance scale on which there has been testing) out to distances of the order of  $10^{15}$  cm or so (viz. solar system distance scales). At much larger distances all tests which use only the sources which were known in advance of observation have been found to fail, with additional (dark matter) sources always having to be invoked after the fact, with the validity of Newton's law yet to be confirmed on those distances.<sup>24</sup> The establishing of a given law in one kinematic region cannot be regarded as evidence for its validity in others.

As regards solar system tests of Newton's law of gravity, these tests are spectacular, with the planets being found to move according to Eq. (80) to very high accuracy when the sun and the planets

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<sup>24</sup>Testing the validity of Newton's law on a distance scale such as galactic would require knowing the dark matter distribution in a galaxy in advance of measurement of the orbital velocities of the material in it.

themselves are taken to be the sources. To a very good approximation the planets move around the sun with a centripetal acceleration of the form

$$\frac{v^2}{r} = \frac{GM_\odot}{r^2} , \quad (81)$$

and have Keplerian fall-off of velocity with distance of the form

$$v = \frac{G^{1/2}M_\odot^{1/2}}{r^{1/2}} \quad (82)$$

where  $M_\odot$  is the mass of the sun. In fact so reliable was this law found to be in the solar system when only known visible sources were used, that when the orbit of Uranus was found to depart slightly from the Newtonian expectation, a perturbation due to a then undetected nearby planet was proposed, with the subsequent detection of Neptune giving dramatic confirmation of Newton's law of gravity at that distance scale.

While Newton's law had always been thought to hold on all solar and sub-solar distance scales, recently it has come in to question on very small laboratory scales of order millimeters or so. While the meter or so distance region had actually been quite thoroughly searched (though to no avail) when it was thought that there might be a so-called fifth force operative on those distance scales (fifth since it would be in addition to the strong, electromagnetic, weak and gravitational forces), recent studies of large extra dimension physics had led to a reopening of the issue with a focus on the even smaller millimeter region distance scale. While there has been a longstanding theoretical interest in the possible existence of additional spacetime dimensions beyond the four established ones, it had generally been presupposed that such dimensions would be truly microscopic, possibly being as small as the  $10^{-33}$  cm Planck length. In an attempt to resolve the longstanding hierarchy problem of understanding why there was such a huge disparity between the  $M_{EW} = 10^3$  GeV electroweak and the  $M_{PL} = 10^{19}$  GeV gravitational mass scales, Arkani-Hamed, Dimopoulos and Dvali [7] found a candidate extra dimension based solution in which rather than be microscopic, the additional dimensions beyond four would instead need to be of the order of millimeters. In such a situation gravitational flux lines could then spread out to such distances in the extra dimensions, and therefore lead to modifications to standard gravity at the millimeter level. Following on the work of Arkani-Hamed, Dimopoulos and Dvali, Randall and Sundrum [8, 9] found an alternate way to address the hierarchy problem which required the geometry in the extra dimensions not to be the Minkowski one that had always been assumed in higher dimensional theories, but rather to be anti-de Sitter. In such a situation the extra dimensions could then not only be large but even infinite in size, with the curvature of the anti-de Sitter space acting as a sort of refractive medium which would sharply inhibit the penetration of gravitational flux lines into the higher dimensions. In such theories there could again be departures from Newton's law at the millimeter level, which thus prompted a renewed search for possible millimeter region departures. Since for the moment, none has yet actually been detected [10], the issue should be regarded as open, and so we shall assume here the validity of Eq. (79) on all solar and sub-solar distance scales including millimeter ones and below, with the deriving of the Newtonian phenomenology on such scales thus being set as a requisite for any candidate gravitational theory.

## 4.2 The second order and fourth order Poisson equations

While Eq. (79) contains the full content of Newton's law of gravity, in order to actually perform the needed sum over sources, it very convenient to recast Newton's law in the form of a second order Poisson equation. Thus if we set

$$\nabla^2\phi(\vec{r}) = g(\vec{r}) , \quad (83)$$

we can write the potential as

$$\phi(\vec{r}) = -\frac{1}{4\pi} \int d^3\vec{r}' \frac{g(r')}{|\vec{r} - \vec{r}'|} , \quad (84)$$

with the potential exterior and interior to a spherically symmetric static source of radius  $R$  then being given by

$$\phi(r > R) = -\frac{1}{r} \int_0^R dr' r'^2 g(r') , \quad \phi(r < R) = -\frac{1}{r} \int_0^r dr' r'^2 g(r') - \int_r^R dr' r' g(r') . \quad (85)$$

However, Eq. (83) is not the only Poisson type equation which will lead to a  $1/r$  potential. Thus if we for instance consider the fourth order

$$\nabla^4 \phi(\vec{r}) = h(\vec{r}) , \quad (86)$$

we can write the potential as

$$\phi(\vec{r}) = -\frac{1}{8\pi} \int d^3\vec{r}' h(r') |\vec{r} - \vec{r}'| , \quad (87)$$

with the potential exterior and interior to a spherically symmetric static source of radius  $R$  then being given by

$$\begin{aligned} \phi(r > R) &= -\frac{r}{2} \int_0^R dr' r'^2 h(r') - \frac{1}{6r} \int_0^R dr' r'^4 h(r') , \\ \phi(r < R) &= -\frac{r}{2} \int_0^r dr' r'^2 h(r') - \frac{1}{6r} \int_0^r dr' r'^4 h(r') - \frac{1}{2} \int_r^R dr' r'^3 h(r') - \frac{r^2}{6} \int_r^R dr' r' h(r') . \end{aligned} \quad (88)$$

Moreover, recovering a  $1/r$  potential is not even restricted to the above choices of Poisson equation, since the sixth order  $\nabla^6 \phi(\vec{r}) = k(\vec{r})$  would yield an exterior potential of the generic form  $\phi \sim 1/r + r + r^3$ , with the pattern repeating for all higher even number of derivative Poisson equations. Characteristic of all of these possible Poisson equations is the fact that we recover a  $1/r$  term for each and every one of them, and in all of them the terms which depart from a pure  $1/r$  form are all only important at larger distances (viz. just the distances on which the dark matter problem is encountered), with there being no modifications to the  $1/r$  law at very small distances. Providing that the  $r$ ,  $r^3$  and so on type terms are all negligible on solar system distance scales, all of these Poisson equations will thus reproduce the standard Newtonian solar system phenomenology.

Now at first glance the expressions for the coefficients of the  $1/r$  terms in Eqs. (85) and (88) appear to differ since they are given as different moments of the source. However, this is not of concern since measurements of  $\phi(r)$  in the  $r > R$  region where there are data can never uncover any information regarding the behavior of the integrand in Eq. (85) in the  $r < R$  region. What is required of the  $1/r$  term in the  $r > R$  region is only that it conform with the large  $r$  behavior required by the original Newton law of Eq. (79) (and thus the all  $r > R$  behavior since  $1/r$  is the unique behavior of the Newton term in the entire exterior region), viz. that for a set of sources all with the same mass  $m$  the potential behave as

$$\phi(r > R) = -\frac{NmG}{r} , \quad (89)$$

i.e. that the coefficient of the  $1/r$  term be linear in the number of sources  $N$ . To achieve such a linearity in the second order Poisson equation case one can choose to set

$$g(r < R) = mG \sum_{i=1}^N \frac{\delta(r - r_i)}{r^2} \quad (90)$$

in the  $r < R$  region, though as we had just noted, this is not mandated by Eq. (79). (While use of this particular  $g(r < R)$  is sufficient to give Eq. (89), it is not necessary.) What is necessary is only that

there be  $N$  discrete sources, with Eq. (79) only counting the number of them in the  $r > R$  region. To achieve exactly the same result through the use of the fourth order Eq. (88), it is convenient, though not mandatory, to consider as source<sup>25</sup>

$$h(r < R) = -\gamma c^2 \sum_{i=1}^N \frac{\delta(r - r_i)}{r^2} - \frac{3\beta c^2}{2} \sum_{i=1}^N \left[ \nabla^2 - \frac{r^2}{12} \nabla^4 \right] \left[ \frac{\delta(r - r_i)}{r^2} \right] , \quad (91)$$

with its insertion in Eq. (88) yielding

$$\phi(r > R) = -\frac{N\beta c^2}{r} + \frac{N\gamma c^2 r}{2} . \quad (92)$$

As we see, the coefficient of the  $1/r$  term is again linear in the number of sources, with Newton's law of gravity only counting the number of discrete sources that are present in the  $r < R$  region. Comparing Eqs. (89) and (92), we see that we can once and for all define  $\beta = mG/c^2$  at the individual discrete source level (a level which could be microscopic), without ever needing to decompose  $\beta c^2$  into separate  $m$  and  $G$  pieces, with only the Schwarzschild radius of the source (viz.  $2mG/c^2$  in the second order case and  $2\beta$  in the fourth order) ever being measurable gravitationally.<sup>26</sup>

It is important to stress that it is the presence of two independent singularities in the source of Eq. (91) which leads to the logical independence of the  $\beta$  and  $\gamma$  coefficients, since if we could approximate the source  $h(r)$  by a constant source  $h(r) = h$ , we would instead have to conclude that the coefficients of the two potential terms would be related according to

$$\phi(r > R) = -\frac{NhR^5}{30r} - \frac{NhR^3 r}{6} , \quad (93)$$

and thus have a ratio of order the radius squared of the source. This is to be contrasted with the choice of a constant source  $g(r) = g$  in the second order case where the potential is then given by

$$\phi(r > R) = -\frac{NgR^3}{3r} . \quad (94)$$

Since a comparison of Eqs. (89) and (94) would entail that one can make the identification  $mG = gR^3/3$ , a view of gravitational sources has developed that macroscopic gravitational sources can be treated as being continuous rather than as being a collection of independent discrete microscopic sources. However, apart from not at all being mandated by the successful use of Eq. (89), the issue is in fact not even addressable if one only makes measurements in the exterior  $r > R$  region. However, in the fourth order Poisson case the issue does become relevant, because here one does measure more than just one moment of the source, with the higher derivative theory thus being able to probe deeper into the source. Thus within the fourth order theory, an experimental determination of the ratio of the coefficients of the  $1/r$  and  $r$  terms which would indicate that the ratio of the two coefficients is very far from being of order the square of the radius of a given macroscopic source, would indicate that the macroscopic source would

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<sup>25</sup>We choose the particular  $[\nabla^2 - (r^2/12)\nabla^4](\delta(r - r_i)/r^2)$  source rather than the more straightforward  $\nabla^2(\delta(r - r_i)/r^2)$  one, since unlike the latter, the former one happens to be positive definite. An additional virtue of this particular choice of source is that it only couples to the fourth moment integral and not to the second moment one; and with the  $\delta(r - r_i)/r^2$  source only coupling to the second moment integral and not to the fourth moment one, the logical independence of the two moments is thus established.

<sup>26</sup>Even if the microscopic  $\beta c^2$  were not to be equal to the product of the conventionally defined  $m$  and  $G$  (not that we know what the gravitational coupling of a microscopic source is), the use of the coefficient  $N\beta c^2$  would then lead to an estimate for the number of atoms in a macroscopic source such as the sun which would be different from the conventional one, a point which is not of concern since without appealing to the  $mG$  coefficient, we do not know how many atoms there anyway are in the sun, with the standard estimate of the number being based on the a priori use of the  $mG$  coefficient.

have to be composed of microscopic components which are discrete. Quite remarkably then, higher order derivative theories open up the possibility of establishing discreteness at the microscopic level from macroscopic measurements alone.<sup>27</sup> Moreover, below we shall actually apply the potential of Eq. (92) to galactic rotation curves, to find that the  $\beta^*$  and  $\gamma^*$  coefficients of the fourth order theory stellar potential

$$\phi^*(r > R) = -\frac{\beta^* c^2}{r} + \frac{\gamma^* c^2 r}{2} \quad (95)$$

actually are very far from possessing a ratio  $\beta^*/\gamma^*$  which is anything like the square of a typical stellar radius (numerically we find  $(\beta^*/\gamma^*)^{1/2} \sim 10^{23}$  cm, with the  $\gamma^* c^2 r/2$  term first becoming competitive with the  $\beta^* c^2/r$  term on galactic distance scales). Within the framework of a gravity which is based on fourth order equations then we can conclude that matter must be discrete at the microscopic level.

While the second order Poisson equation of Eq. (83) does yield the Newtonian potential of Eq. (89) as solution, there is actually a critical difference between the use of the second order Newton potential and Poisson equation on the one hand and the use of their fourth order counterparts given as the potential of Eq. (92) and the Poisson equation of Eq. (86) on the other. Specifically, while the fourth order potential of Eq. (92) does nicely reduce to the second order potential of Eq. (89) in the small  $r$  limit, the fourth order Poisson equation itself never reduces to the second order one in any limit at all. It is thus possible for solutions to two different equations to approximate each other very closely in some kinematic region even while the equations to which they are solutions never approximate each other at all. The equations which describe any candidate alternate theory of gravity thus do not at all need to reduce to those which describe the standard theory. Rather, it is only the solutions to the alternate theory which have to recover the solutions to the standard theory in any of the kinematic regions in which the standard theory solutions have themselves been tested.<sup>28</sup> Outside of such regions the alternate theory can then provide for departures from the standard theory expectations, something we will capitalize on below to explain galactic rotation curve systematics without the need to invoke dark matter. The essential difference between the second order Poisson equation and the fourth order one, is that once the second order Poisson equation is specified as being the fundamental equation of motion, there is no possibility to ever get any departure from Newton's law on any distance scale at all, something which is allowed for in a higher order theory. Since the Einstein equations of Eq. (1) were explicitly constructed so that their non-relativistic weak gravity limit would be the second order Poisson equation, the Einstein equations thus lock non-relativistic weak gravity into being Newtonian on all distance scales, and to thus be so even on distances well beyond the solar system ones where Newton's law of gravity was originally established. In this sense then the dark matter problem could originate in the lack of reliability of the extrapolation of standard solar system wisdom to altogether larger distance scales.

### 4.3 Universal acceleration scale and the MOND theory

While the above discussion has focused on the standard Newtonian picture and those allowed departures from it which are associated with a critical distance scale, an alternate form of allowed departure has been suggested by Milgrom [11, 12, 13], one which has proven to be remarkably instructive and fruitful. Specifically, on noting that the centripetal accelerations of particles in orbits around galaxies were

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<sup>27</sup>As a historical aside, we note that since the use of higher order gravity can thus provide for a macroscopic manifestation of the existence of atoms, it thus meets an objective which was first sought by Boltzmann, one which was part of his motivation for developing kinetic theory (and which was subsequently achieved by a determination of a finite value for Avogadro's number by macroscopic means).

<sup>28</sup>If the alternate equations of motion did reduce to the standard ones in some limit, then of course so would their solutions. However, the converse is not true, since the alternate theory solutions are still able to reduce to the standard theory solutions even if there is no such parallel for the equations of motion themselves.

typically much smaller than those associated with the orbits of planets around the sun or satellites around the earth, Milgrom suggested that the determining factor was not in fact a distance scale at all but rather an acceleration one, with departures from the standard wisdom occurring whenever accelerations of particles dropped below some critical universal acceleration scale  $a_0$ . Quite remarkably, Milgrom found that it is actually possible to choose such an  $a_0$  so that its effects would be of significance for galaxies and yet leave solar system phenomenology untouched.<sup>29</sup> In contrast to proposals which seek to modify the right-hand side (viz. the gravitational side) of Eq. (80), Milgrom instead proposed to modify the left-hand side (viz. the inertial side) with the proposal thus being known as MOND (modified Newtonian dynamics). Specifically Milgrom replaced Eq. (80) by

$$\mu\left(\frac{a}{a_0}\right)\vec{a} = \vec{f} \quad (96)$$

where  $\vec{a} = d^2\vec{r}/dt^2$  is the ordinary acceleration,  $\vec{f} = -\vec{\nabla}\phi$  is the gravitational force and  $\mu(a/a_0)$  is the modification. In order for this modification to reduce to the standard Newton law when  $a \gg a_0$ , the function  $\mu(x)$  has to behave as  $\mu(x \gg 1) \rightarrow 1$ . Guided by the fact that many prominent spiral galaxies possess rotational velocity curves which are flat (i.e. the rotational velocity  $v$  is independent of the distance  $R$  from the center of the galaxy), and that the velocities of these particular galaxies obey the empirical Tully-Fisher law which relates  $v$  to the total luminosity  $L$  of the galaxy as  $v^4 \sim L$ , Milgrom further suggested that at very small  $x$  the function  $\mu(x)$  should behave as  $\mu(x \ll 1) \rightarrow x$ , since the centripetal accelerations  $a = v^2/R \ll a_0$  of orbits in a galaxy of mass  $M$  would then obey

$$\frac{a^2}{a_0} = \frac{v^4}{a_0 R^2} = \frac{MG}{R^2} \quad , \quad (97)$$

and thus yield the  $R$  independent  $v^4 = a_0 MG$ . Then, if the mass involved was just the luminous mass of the galaxy, it would be related to the luminosity of the galaxy by  $M = (M/L)L$  where  $M/L$  is the mass to light ratio of the galaxy, with both the flatness of galactic rotation curves and the Tully-Fisher relation thereby being obtained from luminous matter alone.

As a proposal, the relation of Eq. (96) is perhaps a little bit too strong a departure from the standard picture, since as such it would modify the inertial side of Newton's second law of motion not just for gravitational processes but for non-gravitational ones as well, processes for which there are no known problems. Additionally, since the proposal is a modification of the inertial side of Newton's second law of motion, in an accelerating coordinate system Eq. (96) would not readily recover Eqs. (4) and (5). To get a sense of what might instead happen, one possible way to effect a covariant generalization of Eq. (96) is to replace the acceleration  $\vec{a}$  by the contravariant  $D^2x^\lambda/D\tau^2$  of Eqs. (4) and (5), to introduce a four-vector  $V^\lambda$  which reduces to  $V^\lambda = (1, 0, 0, 0)$  in a non-accelerating coordinate system, and to take  $\mu(x)$  to be a general coordinate scalar function of

$$x = \frac{V_\lambda D^2x^\lambda}{a_0 D\tau^2} \quad (98)$$

where the parameter  $a_0$  is then a general coordinate scalar with the dimensions of acceleration. While the presence of such a  $V_\lambda$  would lead to frame-dependent effects, such departures from standard wisdom need not necessarily be in conflict with observation if  $\mu(x)$  is sensibly close to one in the kinematic regime where tests of the equivalence principle and preferred frame effects have so far been made.

Nonetheless, since it would be more economical not to have to make so radical a conceptual departure from standard wisdom, an alternative procedure for introducing a universal acceleration  $a_0$  would be

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<sup>29</sup>Phenomenologically  $a_0$  was found to be of order  $10^{-8}$  cm/sec<sup>2</sup>, to be compared with an acceleration at the outer limits of the solar system of order  $10^{-3}$  cm/sec<sup>2</sup> due to the pull of the sun.

to have it instead appear on the gravitational side of Eq. (80), with Eq. (80) then being replaced by

$$\vec{a} = \nu \left( \frac{f}{a_0} \right) \vec{f} \quad (99)$$

instead. To get a sense of what the function  $\nu(y)$  might look like, we consider a particularly simple choice for  $\mu(x)$  which meets its required large and small  $x$  limits, viz.

$$\mu(x) = \frac{x}{(1+x^2)^{1/2}} \quad , \quad (100)$$

to then find that  $\nu(y)$  and  $\vec{a}$  are given by

$$\nu(y) = \left( \frac{1}{2} + \frac{(y^2+4)^{1/2}}{2y} \right)^{1/2} \quad (101)$$

$$\vec{a} = \left( \frac{1}{2} + \frac{(f^2+4a_0^2)^{1/2}}{2f} \right)^{1/2} \vec{f} \quad . \quad (102)$$

With the two relevant limits of the function  $\nu(y)$  being  $\nu(y \gg 1) \rightarrow 1$ ,  $\nu(y \ll 1) \rightarrow 1/y^{1/2}$ , we thus recover Eq. (80) when  $f \gg a_0$ , and come right back to Eq. (97) when  $f \ll a_0$ , with Eqs. (96) and (99) leading to the same gravitational phenomenology. Having now presented an analysis of Newtonian gravity and some possible allowed departures from it, we turn next to a discussion of Einstein gravity.

## 5 Einstein gravity

### 5.1 Einstein gravity for static sources and the Newtonian limit

In his construction of a covariant theory of gravity, even though Einstein used a fundamental principle, viz. the equivalence principle, to identify the metric  $g_{\mu\nu}$  as the gravitational field, his choice for the dynamical equation of motion which was to then fix the metric, viz. Eq. (1), appears to have been based on a phenomenological consideration, namely that the theory recover the standard second order Poisson equation in the non-relativistic, weak gravity limit. As such, the second order Einstein equations can be viewed as being a covariant generalization of the second order Poisson equation of Eq. (83), replacing it by ten equations (reducible to six by gauge invariance) of which only the (00) component equation would be of order one in the standard non-relativistic, weak gravity limit in which the components of  $T_{\mu\nu}$  other than  $T_{00}$  are taken to be negligible compared to  $T_{00}$  itself. The great virtue of the Einstein equations is that in prescribing a covariant generalization of the second order Poisson equation, it not only incorporated Newtonian gravity, but also provided it with general relativistic corrections which could then be tested, leading to the three classic tests of general relativity, the gravitational red shift, gravitational bending of light, and the precession of planetary orbits. With the Einstein equations not only satisfying these three tests, but doing so in truly spectacular fashion, by and large the consensus in the community is that the issue is therefore considered settled, with the Einstein equations being the true equations of gravity which no longer need to be open to question.

However, it is important to note that while Einstein's original construction was motivated by the second order Poisson equation, the application of Eq. (1) to the three classic solar system tests is not actually sensitive to this particular aspect of the Einstein equations at all. Specifically, what is needed for the three tests is only knowledge of the metric exterior to a static, spherically symmetric source, viz. the Schwarzschild metric as given earlier as Eqs. (41) and (50), a metric which obeys  $R_{\mu\nu} = 0$  in the source-free region. As such, the exterior Schwarzschild solution is parameterized by a parameter  $MG$

associated with a given source, and while this parameter can be determined from a matching of the exterior solution of Eq. (50) to the interior solution of Eq. (47) to yield Eq. (51), no use of Eq. (51) is actually made in the analysis of the classic tests. While the  $R_{\mu\nu} = 0$  condition does indeed follow from Eq. (1) in the source-free  $T_{\mu\nu} = 0$  region, Eq. (1) does not follow from this fact, i.e. knowing that  $R_{\mu\nu}$  vanishes in the source-free region is not sufficient to determine what  $R_{\mu\nu}$  is to be equal to in regions where it does not vanish. To emphasize the point, we note that while functional variation of the second order Einstein-Hilbert action

$$I_{EH} = -\frac{c^3}{16\pi G} \int d^4x (-g)^{1/2} R^\alpha{}_\alpha \quad (103)$$

with respect to the metric yields the Einstein tensor according to

$$\frac{16\pi G}{c^3 (-g)^{1/2}} \frac{\delta I_{EH}}{\delta g_{\mu\nu}} = G^{\mu\nu} = R^{\mu\nu} - \frac{1}{2} g^{\mu\nu} R^\alpha{}_\alpha \quad , \quad (104)$$

variation of the fourth order general coordinate scalar actions

$$I_{W_1} = \int d^4x (-g)^{1/2} [R^\alpha{}_\alpha]^2 \quad (105)$$

and

$$I_{W_2} = \int d^4x (-g)^{1/2} R_{\alpha\beta} R^{\alpha\beta} \quad (106)$$

respectively yield [14]

$$\frac{1}{(-g)^{1/2}} \frac{\delta I_{W_1}}{\delta g_{\mu\nu}} = W_{(1)}^{\mu\nu} = 2g^{\mu\nu} (R^\alpha{}_\alpha)_{;\beta}{}^{;\beta} - 2(R^\alpha{}_\alpha)^{;\mu;\nu} - 2R^\alpha{}_\alpha R^{\mu\nu} + \frac{1}{2} g^{\mu\nu} (R^\alpha{}_\alpha)^2 \quad , \quad (107)$$

and

$$\frac{1}{(-g)^{1/2}} \frac{\delta I_{W_2}}{\delta g_{\mu\nu}} = W_{(2)}^{\mu\nu} = \frac{1}{2} g^{\mu\nu} (R^\alpha{}_\alpha)_{;\beta}{}^{;\beta} + R^{\mu\nu;\beta}{}_{;\beta} - R^{\mu\beta;\nu}{}_{;\beta} - R^{\nu\beta;\mu}{}_{;\beta} - 2R^{\mu\beta} R^\nu{}_\beta + \frac{1}{2} g^{\mu\nu} R_{\alpha\beta} R^{\alpha\beta} \quad . \quad (108)$$

Then, since the Schwarzschild solution is one in which the Ricci tensor vanishes, in the Schwarzschild solution it follows that covariant derivatives of the Ricci tensor vanish as well, with both  $W_{(1)}^{\mu\nu}$  and  $W_{(2)}^{\mu\nu}$  thus vanishing when  $R_{\mu\nu} = 0$ . Moreover, this same argument immediately generalizes to general coordinate scalar actions based on even higher powers of the Ricci scalar or Ricci tensor. The exterior Schwarzschild solution is thus an exterior solution to any pure metric theory of gravity which uses any Ricci tensor based higher order action in place of the Einstein-Hilbert action,<sup>30</sup> with all of them thus satisfying the three classic tests.<sup>31</sup> Thus, in complete parallel to our earlier discussion of the difference between the second and higher order Poisson equations, we see that even though there is no limit in which any linear combination of  $W_{(1)}^{\mu\nu}$  and  $W_{(2)}^{\mu\nu}$  reduces to the Einstein tensor of Eq. (104), nonetheless the use of any such combination as a gravitational tensor which could replace  $G_{\mu\nu}$  in Eq. (1) would still lead to the Schwarzschild solution. The Schwarzschild solution thus bears the same relation to the Einstein-Hilbert action and its higher order generalizations as the Newton potential does to the Poisson equation and its higher order generalizations, with use of the Einstein equations only being sufficient to secure the three classic tests but not at all necessary. Hence, as had first been noted by Eddington as long ago as only shortly after he himself led an expedition in 1919 to confirm that light did indeed bend

<sup>30</sup>This would not necessarily be true for any higher order Riemann tensor based action since the Schwarzschild metric is not Riemann flat, only Ricci flat.

<sup>31</sup>While the vanishing of the Ricci tensor implies the vanishing of its derivatives as well, the vanishing of the derivatives can be achieved without the Ricci tensor itself needing to vanish, with the vanishing of tensors such as  $W_{(1)}^{\mu\nu}$  or  $W_{(2)}^{\mu\nu}$  thus being achievable by additional, non-Schwarzschild type solutions as well, a point we shall return to below.

as it passed by the sun, the success of the three classic tests does not secure the validity of the Einstein equations, with other options being possible. Indeed, at that time Eddington issued a challenge to the community to tell him which gravitational action then was the correct one to use, a challenge that the community has never answered even though it goes directly to the heart of gravitational theory.

Beyond the three classic tests of general relativity, the only other test on solar system sized distance scales is the decay of the orbit of a binary pulsar, an effect which again only involves an exterior metric, this one due to the presence of two stellar sources rather than just the one involved in the solar system tests. The decay of the orbit is due to gravitational radiation reaction, with each star in the binary responding to retarded gravitational signals emitted by the other and not to instantaneous ones. Such retardation effects will exist in any covariant metric theory of gravity since in all of them gravitational information cannot be communicated faster than the speed of light, and in all of them there will thus be some decay of the binary pulsar orbit. While there thus will be orbit decay not only in Einstein-Hilbert based gravity but also in its higher order alternatives, nonetheless all such alternatives face the (calculationally daunting) challenge of actually obtaining the specific amount of decay which has actually been observed, with calculations having so far only been carried through (and with stunning success) is the second order Einstein theory itself.

## 5.2 The Einstein static universe and the cosmological constant

The one situation in which the full content of the Einstein equations is felt is in cosmology, since in that case both sides of Eq. (1) are non-zero, with cosmological observations being made inside the universe and thus in the region where  $T_{\mu\nu}$  is explicitly non-zero. In order to prepare for the discussion of the dark energy problem and the accelerating universe to be given below, it is useful to recall Einstein's effort to produce a static universe solution to the Einstein equations through the introduction of a cosmological constant. In order to see the structure of the Einstein model it is instructive to start not with a static model but rather with the general time dependent Robertson-Walker geometry of Eq. (36) and the generic source of Eq. (37) with its two separate time-dependent functions  $C(t)$  and  $D(t)$ . For such a set-up the Einstein equations of Eq. (1) yield

$$3c\ddot{R} = -4\pi G(C + 3D)R \quad (109)$$

and

$$\dot{R}^2 + kc^2 = \frac{8\pi G}{3c}CR^2 \quad (110)$$

A solution in which  $C$ ,  $D$  and  $R$  are all constant is then obtained when

$$\frac{D}{C} = -\frac{1}{3} \quad , \quad k = \frac{8\pi G}{3c^3}CR^2 \quad (111)$$

On taking  $C$  to be positive, we see that  $k$  would need to be positive too (viz. topologically closed 3-space), while  $D$  would have to be negative, with the Einstein static universe thus being supportable by a source with  $D/C = -1/3$ . Fluids in which  $D/C$  is negative are known as quintessence fluids and have recently come into prominence [15] through the discovery of the accelerating universe, with a quintessence fluid with  $C + 3D < 0$  then leading to a net cosmic acceleration ( $\ddot{R} > 0$ ) in Eq. (109), though its role in the Einstein model is to put the acceleration precisely at the  $\dot{R} = 0$  borderline where it is neither positive nor negative. In his actual construction, Einstein himself did not use the language of quintessence fluids. Rather, Einstein used an energy-momentum tensor containing just an ordinary non-relativistic matter fluid with a density  $\rho_m$  and zero pressure  $p_m$ , and actually modified Eq. (1) by adding a cosmological constant term to its left-hand side according to

$$-\frac{c^3}{8\pi G} \left( R^{\mu\nu} - \frac{1}{2}g^{\mu\nu}R^\alpha{}_\alpha \right) + \Lambda g^{\mu\nu} = T^{\mu\nu} \quad (112)$$

With the cosmological constant term itself having the form of a perfect Robertson-Walker geometry fluid with energy density  $\rho_\Lambda = c\Lambda > 0$  and pressure  $p_\Lambda = -c\Lambda < 0$  (so that the cosmological constant term itself can be thought of as being a quintessence fluid with  $p_\Lambda = -\rho_\Lambda$  and  $\rho_\Lambda + 3p_\Lambda < 0$ ), we can relate  $\rho_m$  and  $\Lambda$  to  $C$  and  $D$  according to  $C = \rho_m/c + \Lambda$ ,  $D = -\Lambda$ . The requirement that  $C$  and  $D$  obey  $C + 3D = 0$  then requires that  $\rho_m$  and  $\Lambda$  be related according to the very specific fine-tuned relation

$$2c\Lambda = \rho_m \quad . \quad (113)$$

Since on its own the ordinary  $\rho_m$  fluid would lead to  $\ddot{R} < 0$  in Eq. (109) while on its own a positive cosmological constant would lead to  $\ddot{R} > 0$ , as we thus see, to get the Einstein model to be static we need to fine-tune the attraction generated by the energy density  $\rho_m$  and the repulsion generated by the cosmological constant  $\Lambda$  so that they precisely cancel each other. Interestingly, we shall encounter a reflection of such a fine-tuning below in our discussion of the dark energy problem.

With Slipher and Hubble's discovery of the expansion of the universe Einstein realized that there was no further need for a static cosmology,<sup>32</sup> and so he abandoned the cosmological constant. However, with the recent discovery of the acceleration of the universe (i.e. not just an expanding  $\dot{R} > 0$ , but also an accelerating  $\ddot{R} > 0$ ), the cosmological constant has come back into prominence, not just as a fundamental term of the form inserted on the left-hand side of Eq. (112) but also as a dynamically generated spontaneous symmetry breakdown contribution to the energy-momentum tensor on the right-hand side of Eq. (112), with current cosmology requiring these two types of cosmological constant to cancel each other to unbelievable precision.

From a theoretical perspective, the very fact that one is able to introduce the  $\Lambda g_{\mu\nu}$  term on the left-hand side of Eq. (112) at all is a reflection of the fact that Einstein's very choice of the original Eq. (1) did not derive from a fundamental principle but only from the phenomenological need to recover Newton's law.<sup>33</sup> Indeed, it was phenomenology which lead him first to introduce and then subsequently to remove the cosmological constant term. It is this lack of a principle which renders the choice of gravitational action non-unique, with there then being no principle with which to control the contribution of the cosmological constant. In addition, the use of Eq. (1) dictates that the low energy limit of the theory is the second order Poisson equation not just on solar system distances but on all distances. And so we turn now to a study of the motions of particles on larger distance scales to see just how well the second order Poisson equation and its Newtonian potential solution then fare.

## 6 The dark matter problem

### 6.1 The galactic rotation curve problem

It is really quite remarkable how early in the study of galaxies and then clusters of galaxies that mass discrepancies started to appear. Specifically, not all that long after Shapley's work on the shape and size of the Milky Way galaxy and Hubble's demonstration that the spiral nebulae were distant galaxies in their own right, Oort's measurements [16] of the velocities of stars in our local solar neighborhood and Zwicky's [17] and then Smith's [18] measurements of the velocity dispersions of galaxies in clusters of galaxies all pointed to discrepancies. Since the discrepancy that Zwicky detected was based on the use of the Newtonian gravity virial theorem, Zwicky was able to conclude that unless the virial theorem itself did not hold for clusters (either because Newton's law did not hold on those distances or because the cluster had not yet virialized), there then had to be more mass in the cluster than he could detect. As such, it was thought at the time that any such missing mass would be in the form of astrophysical

<sup>32</sup>Even in his model, any departure from the  $2c\Lambda = \rho_m$  fine-tuning condition would lead to a non-static cosmology.

<sup>33</sup>Phenomenological that is unless there is some as yet unknown principle which would actually require the validity of Newton's law on all distance scales, since for the moment Newton's law itself is only phenomenological.

objects which were too faint to be detectable, with the notion that the missing matter might be of an entirely different form that could never be detected optically at all coming only altogether later when it was realized that modern elementary particle physics could actually provide such so called Wimp (weakly interacting massive particle) objects.

While there continued to be further indications of mass discrepancies over time (for a history see e.g. [19]), the situation started to become critical following the HI 21 cm radio studies of the spiral galaxies NGC 300 and M33 by Freeman [20] and of M31 by Roberts and Whitehurst [21] which found no sign of the expected Keplerian fall-off of rotational velocity  $v$  with distance  $R$  from the center of the galaxy. With hydrogen gas being distributed in spiral galaxies out to much further distances than the stars themselves, the HI studies were thus the ideal way to probe the outer reaches of spiral galaxies, and thus the ideal probe to test for the Keplerian fall-off expected of the rotation velocities of particles which were located far beyond the visible stars (viz. the visible mass) of the galaxies. Moreover, the measurement of such rotation velocities would be unhindered by projection effects if studies were made of spirals that were close to being edge on along our line of sight.

The continuing persistence of the cluster problem and concern in general about mass discrepancies prompted detailed exploration of outer region galactic rotation curves, with early systematic surveys being made by Bosma [22] and Begeman [23] as part of the University of Groningen Westerbork Synthesis Radio Telescope survey. What was found in these and in many related studies was that the outer region rotation velocities were systematically larger than the Keplerian expectation, with none giving any indication of a fall-off at all. In Fig. 1 we display the rotational velocity curves of a sample of eleven spiral galaxies which have been identified by Begeman, Broeils and Sanders [24] as being particularly reliable and characteristic of the general pattern of behavior of galactic rotation curves which has been found, with their paper giving complete data references. With the stellar surface matter density of each spiral galaxy in the sample behaving as

$$\begin{aligned}\Sigma(R) &= \Sigma_0 e^{-R/R_0} \ , \\ N^* &= 2\pi \int_0^\infty dR R \Sigma(R) = 2\pi \Sigma_0 R_0^2 \ ,\end{aligned}\tag{114}$$

where  $R_0$  is the scale length and  $N^*$  is the total number of stars,<sup>34</sup> in Fig. 1 we have plotted the rotation velocities (in units of km/sec) in each given galaxy as a function of  $R$  as measured in units of that particular galaxy's own scale length, with the rotation curves being displayed in the figure in order of increasing luminosity from the lowest (DDO 154) to the highest (NGC 2841) in the sample. In Table 1 we provide some characteristics of the galaxies, and note that the luminosities of the galaxies in the selected sample range over a factor of order a thousand.

To get a quick sense of what the luminous Newtonian expectation for galaxies might look like, we note that for an infinitesimally thin disk of stars with distribution  $\Sigma(R) = \Sigma_0 e^{-R/R_0}$  and potential  $V^*(r) = -\beta^* c^2 / r = -M_\odot G / r$ , the total centripetal acceleration is then given by [20]

$$\frac{v_{\text{lum}}^2}{R} = g_\beta^{\text{lum}} = \frac{N^* \beta^* c^2 R}{2R_0^3} \left[ I_0 \left( \frac{R}{2R_0} \right) K_0 \left( \frac{R}{2R_0} \right) - I_1 \left( \frac{R}{2R_0} \right) K_1 \left( \frac{R}{2R_0} \right) \right] \ .\tag{115}$$

(In Appendix A we provide a straightforward derivation of this relation.) From the behavior of the modified Bessel functions, the  $v(R)$  velocity profile associated with Eq. (115) is found to consist of an initial rise from  $R = 0$  up to a peak at  $R = 2.2R_0$  which is then followed by a steady fall-off. With  $\Sigma(R)$

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<sup>34</sup>We assume that light traces mass, so that the surface matter density is proportional to the measured luminous surface brightness with a mass to light ratio which is independent of  $R$ .

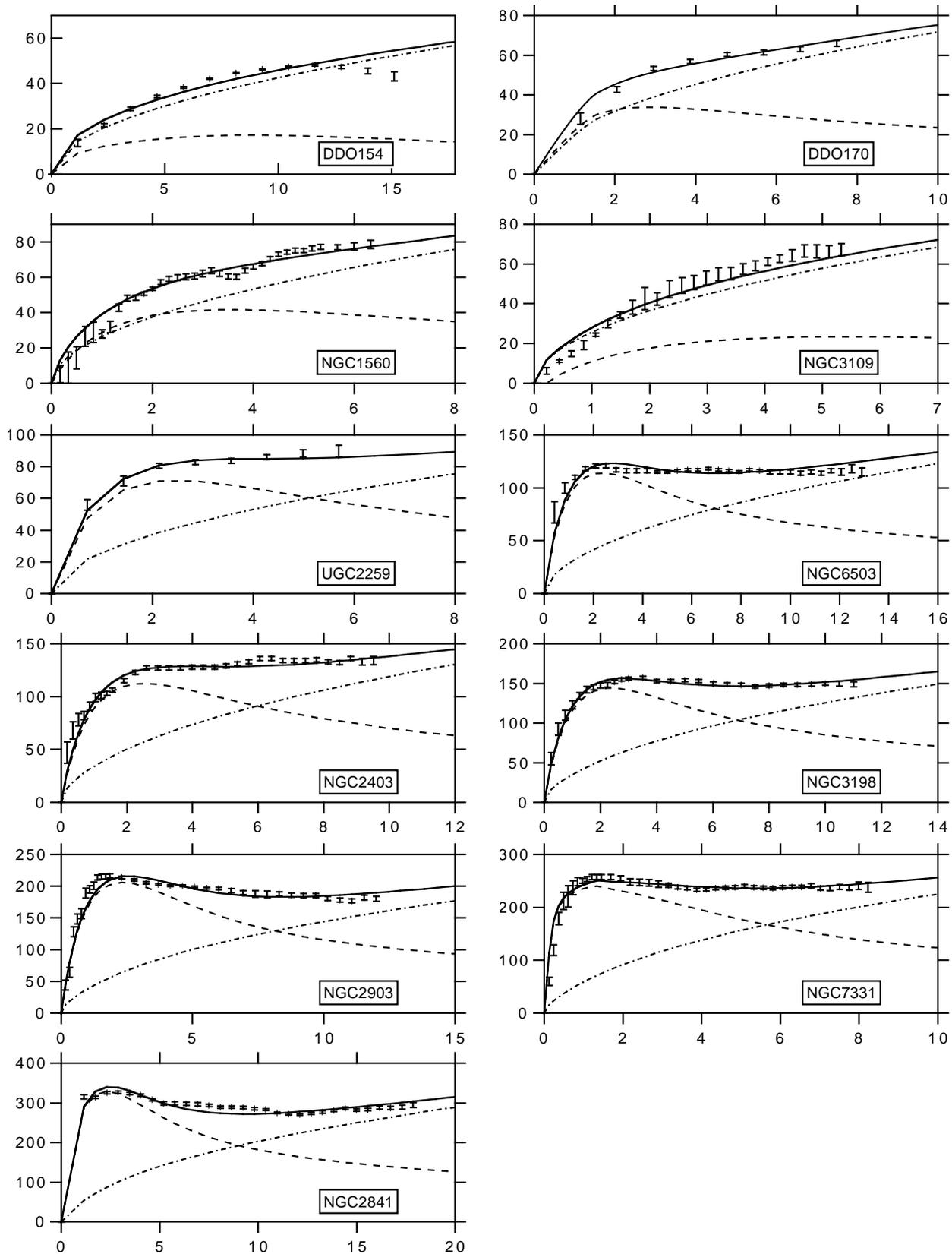


Figure 1: Some typical measured galactic rotation velocities (in km/sec) with their quoted errors as plotted as a function of  $R/R_0$  where  $R_0$  is each galaxy's own optical disk scale length. The dashed (falling) curve shows the velocities that would be produced by the luminous Newtonian contribution alone. The dash-dotted (rising) curve shows the velocities that would be produced by the linear potentials of the alternate conformal gravity theory which is discussed below, and the full curve shows the overall predictions of that theory, with no dark matter being assumed.

Table 1: Characteristics of the eleven galaxy sample

Galaxy	Distance (Mpc)	Luminosity ( $10^9 L_{B_\odot}$ )	$R_0$ (kpc)	$(v^2/c^2 R)_{\text{last}}$ ( $10^{-30} \text{cm}^{-1}$ )	$(M/L)$ ( $M_\odot L_{B_\odot}^{-1}$ )
DDO 154	3.80	0.05	0.48	1.51	0.71
DDO 170	12.01	0.16	1.28	1.63	5.36
NGC 1560	3.00	0.35	1.30	2.70	2.01
NGC 3109	1.70	0.81	1.55	1.98	0.01
UGC 2259	9.80	1.02	1.33	3.85	3.62
NGC 6503	5.94	4.80	1.73	2.14	3.00
NGC 2403	3.25	7.90	2.05	3.31	1.76
NGC 3198	9.36	9.00	2.72	2.67	4.78
NGC 2903	6.40	15.30	2.02	4.86	3.15
NGC 7331	14.90	54.00	4.48	5.51	3.03
NGC 2841	9.50	20.50	2.39	7.25	8.26

typically reaching to about four scale lengths before becoming negligible, the visible stellar material in the galaxy is thus contained almost entirely within this region (the optical disk), with Eq. (115) asymptoting to the Keplerian

$$\frac{v_{\text{lum}}^2}{R} \rightarrow \frac{N^* \beta^* c^2}{R^2} \quad (116)$$

far outside the optical disk. To treat the luminous Newtonian contribution more accurately we need to allow for the fact that the optical disk actually has a thickness (with its own scale height  $z_0$ ), that the HI gas comes with its own surface brightness distribution and can have a total mass as high as 15 per cent or so of the luminous mass, and that the two largest galaxies in the sample also have a central bulge. Using the calculational tools presented in Appendices A, B and C the contributions of all of these various components can readily be determined for each galaxy in the sample, with the exact luminous Newtonian expectation for  $v(R)$  reported in the typical calculation of [25] being exhibited as the dashed curves in Fig. 1. In generating this total luminous Newtonian contribution, the total number of stars  $N^*$  in each galaxy was adjusted so that the luminous Newtonian contribution would account for the entire rotational velocity at the  $R = 2.2R_0$  peak (the maximum disk prescription), with the obtained values for  $N^*$  being presented in Table 1 as galactic mass to light ratios  $M/L$  which thus express  $N^*$  in terms of the measured total luminosity of each galaxy. As we see, normalizing to this maximum disk prescription then completely accounts for the entire initial rise in the rotation curve, to therefore indicate that the inner region is well described by the luminous Newtonian contribution alone. However, as we also see from Fig. 1, in the outer regions of the galaxies the luminous Newtonian contributions totally underaccount for the measured velocities.<sup>35</sup> This then is the mass discrepancy problem in galaxies, and as such it is actually far more severe than the cluster problem, since it is based not on any statistical averaging over the velocities of many galaxies in a cluster or on any assumption regarding their virialization, but only on looking at individual orbits in galaxies one by one just as was done for planetary orbits in the solar system. Not only do we find such discrepancies in galaxies, we wryly note if the galactic data were the only data that were available to us and we had no solar system

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<sup>35</sup>The maximal disk prescription takes the luminous Newtonian contribution to be as large as it possibly can be without overshooting the inner region data. Since this simultaneously fixes the normalization of the Newtonian contribution to the outer region, it thus represents the best that the luminous Newtonian contribution can possibly do in the outer region. And even if we were to change the overall normalization of  $N^*$ , this would anyway not change the shape of the luminous Newtonian contribution, with its behavior at large  $R$  still falling off even then.

data at all, then from the measured galactic rotational velocities alone it would not be possible to infer an inverse square gravitational law at all.<sup>36</sup>

## 6.2 The dark matter solution for galaxies

Given the prior experience with first the prediction and then the discovery of the planet Neptune, it was most natural to try to explain the galactic mass discrepancies by positing the presence of non-luminous dark matter in galaxies. In fact, there was already some theoretical basis for doing so since Ostriker and Peebles [26] had suggested that there should be non-luminous spherical halos (non-luminous since there was no sign of any luminous such halos of the size they envisaged) surrounding unbarred spiral galaxies, due to the fact that in Newtonian gravity such disk shaped structures would not be gravitationally stable, though spherical ones would. However, any such stabilizing halo would only need to envelop the optical disk in order to stabilize it and would thus have no immediate need to extend beyond four disk scale lengths or so. Moreover, since both elliptical galaxies and clusters of galaxies are also found to have mass discrepancies, it would appear that whatever is causing such spherically shaped systems to possess these discrepancies, it could not be stability since such spherical distributions of luminous Newtonian matter are perfectly capable of being stable on their own. Nonetheless, the difficulty with not taking galactic dark matter to be spherically distributed is that if it were also to lie in a disk then it would not be stable either. With the measured rotational velocities far exceeding the luminous Newtonian contributions in the outer galactic regions, within Newtonian gravity any proposed dark matter which would be capable of affecting the shape of the rotation curves in those outer regions would explicitly have to be located in those selfsame regions itself. Moreover, with the measured velocities in the outer galactic regions being at least twice as big as the luminous Newtonian contributions, the dark matter contribution to the  $v^2/R$  centripetal accelerations would thus have to be three to four times as big as the luminous one. Thus not only would dark matter have to be located in precisely the region where there are very few stars, the total amount of dark matter in galaxies would have to greatly exceed the total amount of visible matter, matter which itself is predominantly located in the inner region optical disk. In this sense then the galactic dark matter solution is not entirely equivalent to the Neptune solution to the solar system Uranus problem, since the Neptune modification to solar system orbits was a very minor modification to the Newtonian gravity produced by the sun and not an effect which was of the same order as it. From the point of view of gravity theory, a modification to the galactic luminous Newtonian expectation which is comparable with it in magnitude is a quite drastic modification, while from the point of view of observational astronomy, the notion that such a huge amount of matter is not only not optically observable but is not even predominantly located in the region occupied by the visible stars is equally drastic, to thus underscore how very serious the dark matter problem is.

Beyond the actual assumption of the existence of galactic dark matter at all, its explicit use in galaxies still leaves much to be desired. Specifically, in order to actually fit the galactic rotation curve data explicitly using dark matter, it is conventional to take the spherical dark matter density  $\sigma(r)$  to be the distribution associated with an isothermal Newtonian sphere in hydrostatic equilibrium, viz.

$$\sigma(r) = \frac{\sigma_0}{(r^2 + r_0^2)} \quad , \quad (117)$$

as modified by the introduction of a core radius  $r_0$  to prevent the distribution from diverging at  $r = 0$ . Given such a distribution, the associated centripetal accelerations in the galactic plane are given by (see e.g. Eq. (A29))

$$\frac{v_{\text{dark}}^2}{R} = g_{\beta}^{\text{dark}} = \frac{4\pi\beta^*c^2\sigma_0}{R} \left[ 1 - \frac{r_0}{R} \arctan\left(\frac{R}{r_0}\right) \right] \quad . \quad (118)$$

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<sup>36</sup>It is amusing to ponder how gravity theory might have evolved if the earth had been the only object in orbit around the sun.

While Eq. (118) leads to rotation curves which are asymptotically flat, an actual fitting to an observed flat rotation curve (such as the high luminosity ones displayed in Fig. 1) is not automatically achievable simply by adding  $g_{\beta}^{\text{dark}}$  to  $g_{\beta}^{\text{lum}}$  of Eq. (115). Rather, the two free parameters in  $g_{\beta}^{\text{dark}}$  have to be adjusted galaxy by galaxy. As we see from a typical galaxy such as NGC 3198, the observed rotation curve is pretty flat from the inner region peak near  $R = 2.2R_0$  where (c.f. Eq. (115))

$$v_{\text{lum}}^2 \sim \frac{0.4N^*\beta^*c^2}{R_0} \quad , \quad v_{\text{lum}}^4 \sim \left[0.32\pi\Sigma_0\beta^{*2}c^4\right] N^* \quad , \quad (119)$$

all the way out to the last detected point at  $R \sim 10R_0$ . While the asymptotic limit of Eq. (118) leads to

$$v_{\text{dark}}^2 \rightarrow 4\pi\beta^*c^2\sigma_0 \quad , \quad (120)$$

there is no apparent reason why the dark matter  $\sigma_0$  should be related to the luminous parameters  $N^*$  and  $R_0$  via

$$\sigma_0 = \frac{N^*}{10\pi R_0} \quad , \quad (121)$$

and yet without such a fine-tuning between the dark halo and optical disk parameters, the asymptotic halo contribution would not match the inner region peak velocity at all, being either larger or smaller than it. Moreover, even if this particular fine-tuning is invoked, at best that would only match the velocity at  $R = 2.2R_0$  to the velocity at  $r = 10R_0$  and not require the rotation curve to be flat at all points in between. To achieve such intermediate region flatness, in addition it is necessary to also fine-tune the halo core radius parameter  $r_0$  as well (so as to match at  $R = 6R_0$  or so), and only after this is also done is the flatness of the rotation curve over the entire  $2.2R_0 \leq R \leq 10R_0$  region then secured. Unsatisfactory as this use of two fine-tuned parameters is, the situation is even worse than that since such a fine-tuning has to be made individually for each and every galaxy; and while a two parameter per galaxy prescription is basically found to work well phenomenologically for the rotation curves that have so far been observed, at two halo parameters per galaxy this would already lead to 22 free halo parameters for just the eleven galaxy sample of Fig. 1 alone.

Beyond the isothermal sphere shape for  $\sigma_0$ , numerical  $N$ -body cold dark matter (CDM) simulations of hierarchical clustering have also yielded halos, of which the typical one found by Navarro, Frenk and White [27, 28] is quite close in form to the profile

$$\sigma(r) = \frac{\sigma_0}{r(r+r_0)^2} \quad . \quad (122)$$

Unlike the halos of Eq. (117), with the halos of Eq. (122) being cuspy ones which diverge at  $r = 0$ , they make a quite substantial (though still finite) contribution to rotational velocities in the inner region, and have been criticized on this score [29] since they are found to have some difficulty in fitting galaxies such as dwarfs and low surface brightness galaxies<sup>37</sup> where the inner region luminous Newtonian contribution is small, a region where the true nature of the halo should then be manifest. Moreover, even though these cuspy dark matter halos do provide good fitting to high surface brightness galaxies, since these particular halos would contribute in the inner region, getting good fitting in the high surface brightness case entails giving up the nice accounting of the inner region which the luminous Newtonian contribution provides, and instead requiring an interplay between dark and luminous matter even in that region. Very recently attempts have been made to remedy the cuspieness problem [30, 31] leading to halos close in form to

$$\sigma(r) = \sigma_0 \exp \left[ -\frac{2}{\alpha} \left( \frac{r^n}{r_s^n} - 1 \right) \right] \quad , \quad (123)$$

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<sup>37</sup>Viz. galaxies with a small  $N^*$  or a small central surface brightness  $\Sigma_0$ .

and the issue should be regarded as being open at the present time. However, apart from issues such as this, there is actually a far more serious concern for essentially all dark halo models, namely even if one can generate dark matter halos dynamically, it is not clear what fixes which halo is to go with which given luminous matter distribution, i.e. if the only information available is the measured surface brightness of a given galaxy, from that how do we then determine which values of  $\sigma_0$  and  $r_0$  to use with each such given galaxy so as to be able to predict the structure of the galactic rotation curve in advance of its measurement, and thereby render dark matter theory falsifiable. Forms such as those of Eqs. (122) and (123) only give the generic shapes of halos, and do not provide any a priori determination of values for the  $\sigma_0$  and  $r_0$  parameters or their relation to the luminous  $\Sigma(R) = \Sigma_0 e^{-R/R_0}$  profile, leaving  $\sigma_0$  and  $r_0$  to be free parameters to be fit to galactic data galaxy by galaxy, to then quickly generate large numbers of free parameters as more and more galactic rotation curves are considered.

### 6.3 Clues from the data

In fact perhaps the most serious challenge to the whole dark matter idea is something which is intrinsic to the very motivation which gave rise to it in the first place. Specifically, since galactic rotation curves could not be fitted by luminous Newtonian matter alone, any additional matter which was to be invoked would then have to be non-luminous, and thus be essentially decoupled from the luminous matter. However, as we see from relations such as the  $v^4 \sim L$  Tully-Fisher relation which holds for the prominent high luminosity spiral galaxies, the velocity which is fixed by the total gravitational potential is in fact correlated with the luminosity, and thus whatever is producing this total gravitational potential has to know about the luminous matter even as it is to totally dominate over it in the potential. Moreover, the Tully-Fisher relation is not the only regularity in the data, and in order to get some guidance from the data as to what is needed of dark matter theory (or of any alternate theory for that matter), we now look at some particularly instructive regularities in the galactic data themselves.

As regards the surface brightness, as first noted by Freeman, as well as obey the Tully-Fisher relation, high surface brightness spiral galaxies are found to all have the same central surface brightness  $\Sigma_0^F$ , with all other types of spiral galaxy typically being found to have central surface brightnesses which are lower than this  $\Sigma_0^F$ , with the Freeman value for  $\Sigma_0^F$  serving as an upper limit on the general  $\Sigma_0$ . Moreover, from Eq. (119) we see that Freeman limit galaxies have an inner region peak velocity which obeys

$$v_{\text{lum}}^4 = 0.32\pi\Sigma_0^F\beta^{*2}c^4N^* \quad , \quad (124)$$

with all galaxies with this common  $\Sigma_0^F$  then immediately obeying the Tully-Fisher relation, and doing so entirely because of the behavior of the luminous matter distribution alone, to thus not immediately point in the direction of dark matter.

While most attention in galactic rotation curves has focussed on the flat ones since they are so very striking, nonetheless, as we see from Fig. 1 not all rotation curves are in fact flat. Rather, as first recognized by Casertano and van Gorkom [32], the data fall into three broad classes. Specifically, the rotation curves of the low luminosity dwarf galaxies are typically found to be rising at the last detected data points, the rotation curves of the intermediate to high luminosity galaxies are found to be flat, while the rotation curves of the very highest luminosity galaxies are found to be mildly falling. While there are many flat rotation curves, there are also many which are seen to be rising, and if these curves are to eventually flatten off, this has yet to be observed. Of these various classes, it is only the high luminosity, flat rotation curve galaxies which are Freeman limit, Tully-Fisher galaxies, with the lower luminosity, non-flat galaxies not enjoying this particular form of universality at all.

To explore the rotation curves of these various classes of galaxies in more detail, and to try find a universality which both Freeman limit and sub-Freeman limit galaxies can enjoy, it is very instructive to focus not on the total rotational velocities, but rather on the difference between the total velocity and the luminous Newtonian expectation. With the luminous Newtonian contributions all falling far

outside the optical disk, for galaxies for which the total velocities are flat, the difference between the total and the luminous Newtonian expectation can not itself flat, but must instead actually be rising in the detected region. Inspection of the flat rotation curves in the sample of Fig. 1 shows no galaxy for which the difference between the total and the luminous Newtonian expectation is itself flat, so if the effect of galactic halos is to produce contributions to rotational velocities which are to be asymptotically flat (as per Eq. (120)), no sign of any such asymptotic behavior is yet manifest in the data. A similar shortfall pattern is also found to obtain for the dwarf galaxies, galaxies whose rotation curves are not flat, with the shortfall between the measured rotational velocities and the luminous Newtonian expectation again being seen to be rising. In fact the situation for dwarfs is actually more severe than for bright spirals, since while the ratio of the total velocity to the luminous Newtonian expectation at the last detected data point is a factor of two for the bright spirals, for dwarfs the ratio is a factor of three (and thus a factor of nine in the centripetal acceleration). Mass discrepancies in dwarfs are thus even more pronounced than in the bright spirals.

As we see, the characteristic feature of all the eleven galaxies in the sample is that the shortfall between the measured rotational velocities and the luminous Newtonian expectation is a shortfall which explicitly rises with distance from the center of each galaxy. It is thus of interest to see whether this rise might have any universal structure which is common to all the galaxies in the sample. To this end it is very instructive to evaluate the centripetal acceleration at the last detected data point in each galaxy, with the obtained values being exhibited in Table 1. As we see from the table, the quantity  $(v^2/c^2R)_{\text{last}}$  is found to basically increase with luminosity. With the asymptotic luminous Newtonian contribution to the centripetal acceleration being given by Eq. (116) where  $N^*$  has been fixed once and for all by the inner region maximum disk procedure, its subtraction from  $(v^2/c^2R)_{\text{last}}$  is found to yield a function which satisfies a two parameter formula containing one term which is linear in  $N^*$  and a second term which, quite extraordinarily, is independent of it altogether, with the full  $(v^2/c^2R)_{\text{last}}$  and  $v_{\text{last}}^2$  being given by [25]

$$\begin{aligned} \left(\frac{v^2}{c^2R}\right)_{\text{last}} &= \frac{\beta^*N^*}{R^2} + \frac{\gamma^*N^*}{2} + \frac{\gamma_0}{2} \ , \\ v_{\text{last}}^2 &= \frac{\beta^*c^2N^*}{R} + \frac{\gamma^*c^2N^*R}{2} + \frac{\gamma_0c^2R}{2} \end{aligned} \quad (125)$$

where  $\gamma^*$  and  $\gamma_0$  are two new universal parameters with values

$$\gamma^* = 5.42 \times 10^{-41} \text{cm}^{-1} \ , \quad \gamma_0 = 3.06 \times 10^{-30} \text{cm}^{-1} \ . \quad (126)$$

To get a sense of the relative importance of these various terms, we note that with the solar  $\beta^*$  being given by  $\beta^* = 1.48 \times 10^5 \text{cm}$ , the quantity  $(\beta^*/\gamma^*)^{1/2}$  evaluates to  $(\beta^*/\gamma^*)^{1/2} = .52 \times 10^{23} \text{cm}$ , with the  $\beta^*N^*/R^2$  and  $\gamma^*N^*/2$  terms thus being competitive with each other on none other than galactic distance scales. With the large galaxies typically having of order  $10^{11}$  stars, for them the  $\gamma^*N^*/2$  and  $\gamma_0/2$  are of the same order of magnitude, with all three contributions in Eq. (125) being of importance, and with the falling luminous Newtonian contribution having its most significant effect in the most luminous (largest  $N^*$ ) of the galaxies. However, for small  $N^*$  the  $\gamma_0c^2R/2$  term completely dominates in  $v_{\text{last}}^2$  with the last data points then rising with  $R$ , just as exhibited in Fig. 1 for the dwarf galaxies, and with Eq. (125) precisely capturing the essence of the pattern identified in [32].

In order to assess the significance of Eq. (125), we note that there is nothing at all special about the last data point in each galaxy. Such data points are fixed purely by the observational limit of the radio detectors used to survey the galaxies, a limit which is not fixed by the galaxies themselves but by their distances from us. And yet we are obtaining a universal formula. It appears to us that the only way to achieve such an outcome is to have shortfalls which are rising not just at the last currently detected  $R$  but at all possible  $R$  beyond the optical disk, so that no matter where an individual detector just

happens to cut off, it will be doing so on a curve that is universally rising, to thus always yield the same Eq. (125). Since the eleven galaxy sample is large enough to be characteristic of galactic rotation curve systematics and since it was by no means chosen by the authors of [24] in order to obtain a formula such as that of Eq. (125), we thus believe we can identify Eq. (125) as a universal characteristic of galactic rotation curve systematics, one which is of interest in its own right and which dark matter theory is therefore required to reproduce. With Eq. (125) reducing to the  $N^*$ -independent relation  $v_{\text{last}}^2 = \gamma_0 c^2 R/2$  in the small  $N^*$  limit, the small  $N^*$  limit exposes the departure from the luminous Newtonian expectation in the starkest terms. And it could be that  $\Lambda$ CDM (CDM with a cosmological constant  $\Lambda$ ) generated dark matter halo models are having difficulty fitting dwarfs because relations such as  $v_{\text{last}}^2 = \gamma_0 c^2 R/2$  might not be natural to them.

In regard to Eq. (125), we additionally note that the quantity which measures the departure of the centripetal acceleration  $(v^2/R)_{\text{last}}$  from the luminous Newtonian expectation is a universal acceleration  $\gamma_0 c^2/2$  ( $\gamma_0 + \gamma^* N^*$  is never smaller than  $\gamma_0$ ) whose magnitude is given by  $\gamma_0 c^2/2 = 1.4 \times 10^{-9}$  cm/sec/sec, a value which is close to that found for the universal acceleration parameter  $a_0$  of the MOND theory discussed above. Further support for a role for a universal acceleration was noted by McGaugh [33] who studied the behavior of the quantity  $M_{\text{dyn}}(R)/M_{\text{lum}}(R)$  as a function of the measured centripetal acceleration  $v^2(R)/R$  at points  $R$  within galaxies. Here  $M_{\text{dyn}}(R) = Rv^2(R)/G$  is the total amount of matter interior to  $R$  as would be required by Newtonian gravity to produce the measured  $v^2(R)/R$ , while  $M_{\text{lum}}(R)$  is the total amount of luminous matter which is detected within the same region. In his study of a quite extensive sample of galaxies McGaugh found that mass discrepancies [viz.  $M_{\text{lum}}(R)$  less than the needed  $M_{\text{dyn}}(R)$ ] systematically occurred in galaxies whenever the measured  $v^2(R)/R$  fell below a universal value of order  $10^{-8}$  cm/sec/sec or so, a value which is in accord with that of both the acceleration  $a_0$  associated with MOND and the  $\gamma_0 c^2/2$  acceleration associated with Eq. (125). Thus from both McGaugh's phenomenological study and from the phenomenological Eq. (125), we see that it is indeed an acceleration scale (or equivalently an inverse distance scale times the square of the speed of light) which determines exactly when luminous Newtonian matter fails to account for observed data. As such, the acceleration criterion which Milgrom originally identified has to be regarded as being a bona fide property of nature. As formulated, the MOND theory actually possesses not one but two logically independent components. It comes not just with this acceleration criterion which indicates when the luminous Newtonian prediction is to fail, but also it provides a model (Eq. (99)) for how Newtonian gravity is to then be modified in the region (the MOND region) where the luminous Newtonian expectation does in fact fail. However, regardless of how the theory is to behave in the MOND region, i.e. regardless of what particular form the function  $\nu(f/a_0)$  might take,<sup>38</sup> and regardless of whether one likes any particular form (such as the one of Eq. (101) with its square root structure) that might be used to fit data, it remains true that Milgrom's acceleration criterion is the criterion which determines when departures from the luminous Newtonian expectation are to first occur, a criterion which fundamental theory is therefore obliged to explain. As we thus see, from Milgrom's work, from McGaugh's work and from Eq. (125), there is a lot of universality in the rotation curve data, universality which does not immediately appear to be natural to dark matter and which therefore challenges it, a universality which we will use below to guide us when we consider alternatives to dark matter.

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<sup>38</sup>If in Eq. (125) we ignore the  $\gamma^* N^*/2$  term and set  $a_0 = \gamma_0 c^2/2$ , we can then write Eq. (125) in the generic form of Eq. (99) with the function  $\nu(f/a_0)$  being given by  $\nu(f/a_0) = 1 + a_0/f$ , an issue we will elaborate on further below.

## 7 The dark energy problem

### 7.1 The cosmological dark matter and flatness problems

As we go to scales even larger than galaxies, mass discrepancies are not only found to persist, two new types of mass discrepancy begin to appear, one for clusters of galaxies and the other for the entire universe itself. While non-relativistic luminous Newtonian expectation shortfalls are found not just for galaxies but for clusters of galaxies as well, for such clusters general-relativistic shortfalls are also found, with the use of the Schwarzschild metric for a cluster leading to far less gravitational lensing by the cluster than is observed if only the luminous inventory of the cluster is used as the Schwarzschild source. Thus for clusters one can say that either the dark matter idea is internally consistent since both non-relativistic and general-relativistic shortfalls are found in one and the same system, or one can say that absent dark matter, with clusters it is not just Newtonian gravity which is failing but Einstein gravity too, with the cluster geometry probed by light from a distant quasar not being the Schwarzschild metric associated with the visible sources. In making a statement such as this it is important to note that the detection of gravitational lensing is generally regarded as being a great triumph for Einstein gravity. However the fact of lensing as opposed to the amount of lensing only requires that gravity be a metric theory to which light couples. It is the amount of lensing which is sensitive to the specific make-up of the lens and the specific geometry which it produces, and unless there is dark matter in clusters, then the standard theory would simply not be predicting enough. As with our discussion of  $\Lambda$ CDM generated galactic halos, what is needed to make cluster dark matter into a falsifiable theory is a prediction of the amount of lensing given only a knowledge of the luminous content of the cluster. Currently the amount of dark matter in clusters is inferred only after the lensing measurements have been made. While making lensing predictions based only on the known luminous content of a cluster thus serves as a goal for dark matter, at the present time it also remains an objective for the alternate gravitational theories we shall discuss below, including those for which galactic rotation curve predictions can be made from luminous information alone. Progress on cluster gravitational lensing should thus be definitive for both dark matter and its potential alternatives.

The other large distance scale on which mass discrepancies are found is the largest distance scale there is, viz. that of the universe itself, with the standard theory not only again encountering the issue of dark matter, this time the problem is compounded by a need for dark energy as well. To heuristically see why such problems might be anticipated, it is instructive to consider the implications for a simple Newtonian cosmology of the existence of a universal acceleration. Specifically, we recall that in a critical density Newtonian cosmology of size the Hubble radius  $L = c/H$  (viz. a cosmology where  $\rho_C = 3Mc^2/4\pi L^3 = 3H^2c^2/8\pi G$ ), the Hubble radius happens to be equal to the Schwarzschild radius of the matter within it (viz.  $2MG/c^2 = L$ ). For such a cosmology the centripetal acceleration of a test particle located at the Hubble radius is given by  $v^2/L = MG/L^2 = cH/2$ . Using a value of  $72 \pm 8$  km/sec/Mpc for  $H$  [34], we find that  $cH/2$  evaluates to  $cH/2 = 3.5 \times 10^{-8}$  cm/sec/sec, and with this value being very close to the values found for both the MOND  $a_0$  and the acceleration  $\gamma_0 c^2/2$  associated with Eq. (125), we see that the centripetal accelerations of Newtonian cosmology occur precisely in the acceleration regime where luminous Newtonian matter falls short for galaxies. We this anticipate, and shall in fact shortly see, that a pure luminous matter based standard Einstein gravity cosmology will not adequately fit available cosmic data either. Moreover, in addition this time we will find that the situation can not even be rectified by invoking dark matter, with it being this very failure which has led to the introduction of yet another ad hoc dark object, viz. dark energy.

In standard Einstein gravity a full general-relativistic treatment of cosmology is based on the Friedmann equations of Eqs. (109), (110) and (39)

$$\dot{R}^2 + kc^2 = \frac{8\pi G}{3c} CR^2 \quad , \quad 3\ddot{R} = -\frac{4\pi G}{c}(C + 3D)R \quad , \quad D = -\frac{d}{dR^3} (R^3 C) \quad , \quad (127)$$

as written for the moment in terms of the generic source  $T^{\mu\nu} = [C(t) + D(t)]U^\mu U^\nu + D(t)g^{\mu\nu}$  given in Eq. (37). Given the Robertson-Walker metric of Eq. (36), the objective of cosmology is to then determine the parameters  $k$  and  $R(t)$  which appear in the metric and determine the matter content of the universe as expressed through  $C(t)$  and  $D(t)$ . As a result of many years of observational effort, values for the Hubble parameter  $\dot{R}/R = H$  have now converged on an agreed value for  $H$  as given above. With  $H$  and  $G$  being known, it is convenient to introduce the dimensionless quantities

$$\Omega_C(t) = \frac{8\pi G C(t)}{3cH^2} \quad , \quad \Omega_k(t) = -\frac{kc^2}{\dot{R}^2(t)} \quad , \quad (128)$$

and the deceleration parameter

$$q(t) = -\frac{\ddot{R}R}{\dot{R}^2} \quad , \quad (129)$$

since in terms of them the Friedmann equations of Eq. (127) can be written very compactly as

$$\dot{R}^2 + kc^2 = \dot{R}^2 \Omega_C(t) \quad , \quad \Omega_C(t) + \Omega_k(t) = 1 \quad , \quad q(t) = \frac{1}{2} \left( 1 + \frac{3D}{C} \right) \Omega_C(t) \quad . \quad (130)$$

Historically, the cosmological source was taken to be ordinary matter with  $C = \rho_m/c$ ,  $D = p_m/c$ , with equation of state  $p_m = \rho_m/3$  for radiation and high energy ( $kT \gg mc^2$ ) matter and equation of state  $p_m = 0$  for low energy ( $kT \ll mc^2$ ) matter, with an inventory of the visible universe using galaxy counts leading to a current era value for

$$\Omega_M(t) = \frac{8\pi G \rho_m(t)}{3c^2 H^2} \quad (131)$$

of order 0.01 at the current time  $t_0$ .<sup>39</sup> Such a value for  $\Omega_M(t_0)$  is quite close to the critical value  $\Omega_M = 1$ , the value at which the Friedmann equation would require the parameter  $k$  to be zero (flat universe).<sup>40</sup> As first noted by Dicke, this closeness of  $\Omega_M(t_0)$  to one creates a severe problem for the Friedmann equation, the so-called flatness problem. Specifically, regardless of the value of  $k$ , the use of the above normal matter equations of state in the Friedmann equation leads to the occurrence of a big bang singularity in the early universe, with  $\dot{R}(t=0)$  diverging at  $t=0$ . Near to such a singularity the same Friedmann equation then requires that  $\Omega_M(t=0)$  be incredibly close to one, again regardless of the choice of  $k$ . Then, with  $\Omega_M(t=0)$  being close to one in the early universe, and with the universe then expanding for an entire Hubble time ( $1/H(t_0) \sim 10^{10}$  years), it is very hard to understand why the current era  $\Omega_M(t_0)$  would be anywhere near unity today instead of having redshifted to close to zero by now. In fact for the Friedmann equation to lead to a cosmic evolution in which the current era value of  $\Omega_M(t_0)$  would be close to one today requires fine-tuning the initial value of  $\Omega_M(t=0)$  at a typical early universe time to be one to within one part in  $10^{60}$  or so. As an evolution equation, the Friedmann equation is highly unstable, with it not being able to easily reconcile its initial singularity with the current era value of  $\Omega_M(t_0)$ . More naturally, it would have  $\Omega_M(t)$  fall to zero in the early universe itself.

## 7.2 Fine-tuning and the inflationary universe

A beautiful solution to this conflict was proposed by Guth [35], the so-called inflationary universe, in which a de Sitter phase was to precede the current Robertson-Walker phase, with the rapid exponential expansion of the de Sitter phase precisely providing a dynamics which would lead to the needed initial condition for the subsequent Robertson-Walker phase, no matter what the value of  $k$ . An additional

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<sup>39</sup>Since such inventorying does not undercount, 0.01 thus serves as a lower bound on  $\Omega_M(t_0)$ .

<sup>40</sup>It is called the critical value since the sign of the 3-curvature  $k$  of the universe has the same sign as  $\Omega_M(t_0) - 1$ .

virtue of having such an early universe de Sitter phase was that it also provided a solution to the horizon problem associated with the observed uniformity of the cosmic microwave background across the sky, with the de Sitter phase being able to causally connect regions which a luminous ( $\Omega_M(t_0) = 0.01$ ) or even a luminous plus dark ( $\Omega_M(t_0) > 0.01$ ) matter Friedmann universe could not.<sup>41</sup>

Geometrically it is possible to describe a de Sitter phase using the comoving Robertson-Walker coordinate system, and for it we can take the Friedmann equation to be of the form associated with a cosmological constant source, viz.

$$\dot{R}^2(t) + kc^2 = \alpha c^2 R^2(t) \quad , \quad (132)$$

where  $\alpha = 8\pi G\Lambda/3c^3$ . For the allowable choices of sign for  $\alpha$  and  $k$  solutions to Eq. (132) can readily be found, and they are given as

$$\begin{aligned} R(t, \alpha < 0, k < 0) &= \left(\frac{k}{\alpha}\right)^{1/2} \sin[(-\alpha)^{1/2}ct] \quad , \\ R(t, \alpha = 0, k < 0) &= (-k)^{1/2}ct \quad , \\ R(t, \alpha > 0, k < 0) &= \left(-\frac{k}{\alpha}\right)^{1/2} \sinh(\alpha^{1/2}ct) \quad , \\ R(t, \alpha > 0, k = 0) &= R(t=0)\exp(\alpha^{1/2}ct) \quad , \\ R(t, \alpha > 0, k > 0) &= \left(\frac{k}{\alpha}\right)^{1/2} \cosh(\alpha^{1/2}ct) \quad . \end{aligned} \quad (133)$$

On defining

$$\Omega_\Lambda(t) = \frac{\alpha c^2}{H^2} = \frac{8\pi G\Lambda}{3cH^2} \quad , \quad (134)$$

we find that in the de Sitter phase the quantities  $q(t)$  and  $\Omega_\Lambda(t)$  are related by  $q(t) = -\Omega_\Lambda(t)$ , with  $q(t)$  and  $\Omega_\Lambda(t)$  then being found to be given by

$$\begin{aligned} \Omega_\Lambda(t, \alpha < 0, k < 0) &= -q(t, \alpha < 0, k < 0) = -\tan^2[(-\alpha)^{1/2}ct] \quad , \\ \Omega_\Lambda(t, \alpha = 0, k < 0) &= -q(t, \alpha = 0, k < 0) = 0 \quad , \\ \Omega_\Lambda(t, \alpha > 0, k < 0) &= -q(t, \alpha > 0, k < 0) = \tanh^2(\alpha^{1/2}ct) \quad , \\ \Omega_\Lambda(t, \alpha > 0, k = 0) &= -q(t, \alpha > 0, k = 0) = 1 \quad , \\ \Omega_\Lambda(t, \alpha > 0, k > 0) &= -q(t, \alpha > 0, k > 0) = \coth^2(\alpha^{1/2}ct) \end{aligned} \quad (135)$$

in the various cases. As we see from the solutions, when  $\alpha$  is taken to be positive, then no matter what the value of  $k$ , at late times the expansion radius will grow as  $\exp(\alpha^{1/2}ct)$  with  $\Omega_\Lambda(t)$  asymptoting to one and  $\Omega_k(t) = 1 - \Omega_\Lambda(t)$  to zero. What a period of de Sitter inflation thus achieves is that it quenches the contribution of curvature to cosmic expansion, with the universe inflating so much that curvature is no longer able to contribute to the expansion. If at the end of the inflationary era all of the energy density contained in  $\Omega_\Lambda(t)$  can be converted into a matter  $\Omega_M(t)$ ,  $\Omega_\Lambda(t)$  will then play no further role and  $\Omega_M(t)$  will then enter the Robertson-Walker era with a value that will be extremely close to one in a universe in which curvature is no longer capable of influencing cosmic expansion. With  $\Omega_k(t)$  thus remaining negligibly small throughout the Robertson-Walker era, and with  $\Omega_\Lambda(t)$  have been disposed of,  $\Omega_M(t)$  will then continue to remain close to one, and will thus still be extremely close to one today. The inflationary solution to the flatness problem is thus to make an  $\Omega_M(t)$  which is of order one actually be equal to one (in each and every Robertson-Walker epoch), with the shortfall between a luminous contribution of 0.01 and one itself then requiring a 0.99 contribution through matter that

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<sup>41</sup>That a de Sitter phase would solve the horizon problem had also been noted in [36] and [37].

will contribute to  $\Omega_M(t)$  with exactly the same equation of state as ordinary matter but whose visual absence would require it to be non-luminous. Inflation thus leads to the need for cosmological dark matter in an amount which would make the current era  $\Omega_M(t_0)$  be equal to one, viz. to precisely the critical density value. With such a value for  $\Omega_M(t_0)$ , and with  $D/C$  being zero for the non-relativistic fluids (dark or luminous) of relevance at the three degree current era temperature, inflation would thus entail that the current era  $q(t_0)$  would be equal to  $1/2$  (c.f. Eq. (130)), so that even while inflation itself leads to a negative  $q(t)$  (c.f. Eq. (135)) in the inflationary era, following the transfer of energy density from  $\Omega_\Lambda(t)$  to  $\Omega_M(t)$  at its end, we transit into a Robertson-Walker phase in which  $q(t)$  is positive.<sup>42</sup> While inflation thus provides a nice solution to both the flatness and the horizon problems, the very fact that it produces an era of exponential expansion no matter what the value of  $k$  means that inflation inflates away any dependence on  $k$ , with inflation thus not being sensitive to the sign or the magnitude of  $k$ . The fact that  $\Omega_k(t_0)$  is to be negligible today does not mean that  $k$  itself is zero. Rather inflation enables the current era  $\Omega_k(t_0)$  to be negligible even if  $k$  is non-zero, though this very insensitivity to  $k$  means that the physics of what is to actually fix  $k$  is thus not addressed by inflation. Since the universe does have a global topology (flat, open or closed), something does have to fix it, with physics either beyond inflation (or perhaps instead of it) still being required, a point we return to below.

### 7.3 The accelerating universe and the cosmological dark energy solution

While the development of the inflationary universe model immediately prompted a vigorous search for dark matter candidates, a search which so far has only yielded massive neutrinos (though with masses which are too small to be of significance for the needed  $\Omega_M(t_0)$ ), curiously, at no time did any astrophysical observations themselves ever indicate that  $\Omega_M(t_0)$  was in fact equal to one or that  $q(t_0)$  was equal to  $1/2$ . In fact the best estimates for  $\Omega_M(t_0)$  as obtained from using dark matter to fit the luminous shortfalls found for clusters of galaxies, lead to an  $\Omega_M(t_0)$  of no more than 0.3 [3]. As regards determining  $q(t_0)$ , we note that since the deceleration parameter involves a second derivative with respect to time, its measurement requires a measurement of the change in the first derivative (the Hubble parameter) over cosmologically separated distances, and thus requires an extension of the Hubble plot out to very high redshift. A sufficient such extension out to  $z = 1$  has recently been made using type Ia supernovae as standard candles, and has led [1, 2] to an epochal discovery, namely that far from being equal to  $1/2$ , the current era deceleration parameter is not even positive at all, with the current universe actually being in an accelerating rather than a decelerating phase. However, before addressing this remarkable finding of the supernovae data, we note first that a fit to the 54 most reliable of the original 60 type Ia supernovae sample of [2] using the luminosity function

$$d_L = \frac{2c}{H(t_0)}(1+z) \left( 1 - \frac{1}{(1+z)^{1/2}} \right) \quad (136)$$

associated with a pressureless  $\Omega_M(t_0) = 1$ ,  $\Omega_k(t_0) = 0$  universe is found to yield a  $\chi^2$  value of 92.9,<sup>43</sup> a value which thereby actually renders an  $\Omega_M(t_0) = 1$ ,  $\Omega_k(t_0) = 0$  universe unviable. Consequently, recent supernovae data rule out the  $\Omega_M(t_0) = 1$  inflationary universe solution to the flatness problem.

To get a sense of what is required to fit the supernovae data, we note that as the matter content is reduced in this same pressureless  $\Omega_M(t_0)$ ,  $\Omega_k(t_0)$  universe, a reduction which requires a concomitant increase in  $\Omega_k(t_0)$ , the associated  $\chi^2$  is found to decrease, reaching a value of  $\chi^2 = 61.5$  when  $\Omega_M(t_0)$  reaches its minimum value of zero, a situation in which  $\Omega_k(t_0)$  has then risen to one (viz.  $k$  now negative), the deceleration parameter has dropped to zero, and the luminosity function can be written

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<sup>42</sup>That a positive cosmological constant  $\alpha$  leads to a negative  $q(t)$  is something we shall return to below.

<sup>43</sup>The various  $\chi^2$  values reported here and below are taken from [38, 39].

in the closed form

$$d_L = \frac{c}{H(t_0)} \left( z + \frac{z^2}{2} \right) . \quad (137)$$

Since this particular model is very close to a universe with just regular luminous matter ( $\Omega_M(t_0) = 0.01$ ) and no other components, we see that while this model still gives a  $\chi^2$  which is somewhat larger than the number of independent degrees of freedom, the model does not fare all that badly, suggesting that it may have captured some of the essence of the Hubble plot data. In fact paradoxically, a pure luminous model with  $\Omega_M(t_0) = 0.01$ ,  $\Omega_k(t_0) = 0.99$  actually does better with the supernovae data than a  $k < 0$  luminous plus dark matter model with  $\Omega_M(t_0) = 0.3$ ,  $\Omega_k(t_0) = 0.7$ .<sup>44</sup> Within a pressureless  $\Omega_M(t_0)$ ,  $\Omega_k(t_0)$  model, the only way to reduce the  $\chi^2$  even lower than its  $\Omega_M(t_0) = 0$ ,  $\Omega_k(t_0) = 1$  value is to make either  $G$  or  $\rho_m$  negative – not that one of course can do so in the model, though it does show the needed trend, with the data favoring an even more reduced deceleration parameter, viz. one which would actually be negative, a trend which having large amounts of dark matter actually works against. As it stands then, we see that the supernovae data exclude the only value of  $\Omega_M(t_0)$  for which the Friedmann evolution equation has so far been shown to not have a fine-tuning problem (viz.  $\Omega_M(t_0) = 1$ ), and leaves us with a universe which has even more cosmic repulsion than that achievable by the smallest permissible  $\Omega_M(t_0)$ . We stress this point now to show that even before invoking the cosmological constant (to get some additional cosmic repulsion akin to that which it provided in the original Einstein static universe discussed earlier), once  $\Omega_M(t_0)$  is not equal to one we are right back in the flatness problem, with the Friedmann equation again requiring fine-tuning. Thus even while the introduction of the cosmological constant will also entail fine-tuning, a fine-tuning of its value down from that anticipated from fundamental physics, this particular fine-tuning is quite distinct from that associated with the Friedmann cosmic evolution equation, with the standard cosmological evolution equation already having fine-tuning problems even without a cosmological constant, fine-tuning problems which do not disappear following its introduction.

As regards the possible addition of a cosmological constant to the model, an addition in which Eq. (130) is then replaced by

$$\begin{aligned} \dot{R}^2 + kc^2 &= \dot{R}^2[\Omega_M(t) + \Omega_\Lambda(t)] \quad , \quad \Omega_M(t) + \Omega_\Lambda(t) + \Omega_k(t) = 1 \quad , \\ q(t) &= \frac{1}{2} \left( 1 + \frac{3p_m}{\rho_m} \right) \Omega_M(t) - \Omega_\Lambda(t) \quad , \end{aligned} \quad (138)$$

we now see that a positive  $\Omega_\Lambda(t)$  will nicely serve to reduce  $q(t)$  just as we would want. To get a sense of how big we would need the  $\Omega_\Lambda(t)$  contribution to be, we note that for a pure  $\Omega_\Lambda(t) = 1$ ,  $\Omega_M(t) = 0$ ,  $\Omega_k(t) = 0$  universe the luminosity function can be written in the closed form

$$d_L = \frac{c}{H(t_0)} (z + z^2) \quad , \quad (139)$$

with it leading to  $\chi^2 = 75.8$  for the same 54 supernovae data points. While this value is still too large to be acceptable, it is closer to the number of degrees of freedom than the  $\chi^2$  associated with a pure  $(\Omega_M(t), \Omega_\Lambda(t), \Omega_k(t)) = (1, 0, 0)$  universe. Hence within the family of universes with  $\Omega_k(t_0) = 0$ , we can anticipate a best fit which is weighted more to  $\Omega_\Lambda(t)$  than to  $\Omega_M(t)$ , with the values  $\Omega_M(t_0) = 0.3$ ,  $\Omega_\Lambda(t_0) = 0.7$  being found to yield  $\chi^2 = 57.7$  and the explicit data fit displayed as the lower curve in Fig. 2. With these values of  $\Omega_M(t_0)$  and  $\Omega_\Lambda(t_0)$  the deceleration parameter given by Eq. (138) evaluates to  $q(t_0) = -0.55$ , and is thus seen to be quite negative.

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<sup>44</sup>Of course both of these models will still have to be excluded since neither of them can account for the uniformity of the cosmic microwave background, and both of them suffer from the flatness problem.

## 7.4 The freedom inherent in fitting the current supernovae data

To get a sense of the fitting freedom inherent in the supernovae data, it is instructive to look not just at the curvature-free  $\Omega_k(t) = 0$  family of solutions, but also at a particular class of matter-free  $\Omega_M(t) = 0$  solutions, viz. the  $\alpha > 0, k < 0$  ones as given in Eqs. (133) and (135) with  $R(t, \alpha > 0, k < 0) = (-k/\alpha)^{1/2} \sinh(\alpha^{1/2} ct)$  and  $\Omega_\Lambda(t, \alpha > 0, k < 0) = -q(t, \alpha > 0, k < 0) = \tanh^2(\alpha^{1/2} ct)$ . This particular family of solutions is one which is accelerating in all epochs [ $q(t, \alpha > 0, k < 0)$  is negative at all  $t$  for any positive  $\alpha$ ], and on parameterizing  $\alpha$  in terms of the current era deceleration parameter  $q(t_0) = q_0$ , its luminosity function is given [38, 39] as the one parameter

$$d_L = -\frac{c}{H(t_0)} \frac{(1+z)^2}{q_0} \left( 1 - \left[ 1 + q_0 - \frac{q_0}{(1+z)^2} \right]^{1/2} \right) . \quad (140)$$

Using this luminosity function a best fit of  $\chi^2 = 58.62$  is obtained with  $q_0 = -0.37$ , a deceleration parameter which again is negative, and is exhibited as the upper curve in Fig. 2.<sup>45</sup> With the  $(\Omega_M(t_0), \Omega_\Lambda(t_0), \Omega_k(t_0)) = (0.3, 0.7, 0.0)$  and the  $(\Omega_M(t_0), \Omega_\Lambda(t_0), \Omega_k(t_0)) = (0.0, 0.37, 0.63)$  fits being of equal and indistinguishable quality, the best fits in the  $(\Omega_M(t_0), \Omega_\Lambda(t_0))$  plane will to good approximation thus lie along the line

$$\Omega_\Lambda(t_0) = 1.1\Omega_M(t_0) + 0.37 \quad , \quad \Omega_k(t_0) = 0.63 - 2.1\Omega_M(t_0) \quad (141)$$

which runs through these particular two fits. Along this line we obtain

$$q(t_0) = -0.6\Omega_M(t_0) - 0.37 \quad , \quad (142)$$

with the current era deceleration parameter thus being negative everywhere along the best fit line. With  $\Omega_\Lambda(t_0)$  being zero on the best fit line when  $\Omega_M(t_0)$  is given by  $\Omega_M(t_0) = -0.34$ , we recover our earlier observation that in pure matter universes with no cosmological constant, the best fits are actually obtained with  $\Omega_M(t_0)$  negative.

## 7.5 The cosmic coincidence problem

From the point of view of the standard model, there will be a problem no matter where on the best fit line it finds itself, since along the best fit line  $\Omega_M(t_0)$  and  $\Omega_\Lambda(t_0)$  are of the same order of magnitude. In the presence of a cosmological constant, the occurrence of an early universe initial singularity in the Friedmann equation of Eq. (138) this time translates into the requirement that at the initial time the quantities  $\Omega_M(t_0)$  and  $\Omega_\Lambda(t_0)$  must add up to one, viz.

$$\Omega_M(t=0) + \Omega_\Lambda(t=0) = 1 \quad . \quad (143)$$

However, since  $\rho_m$  redshifts while  $\Lambda$  does not, in order for these two quantities to be of the same order today, in the initial universe  $\Omega_M(t=0)$  must have been incredibly close to one (typically to within one part in  $10^{60}$  or so) and  $\Omega_\Lambda(t=0)$  must have been of order  $10^{-60}$ . There is thus again a fine-tuning problem, with the cosmic evolution Friedmann equation again not being able to naturally reconcile its initial singularity with the current era values of  $\Omega_M(t_0)$  and  $\Omega_\Lambda(t_0)$ . With the most natural expectation for a Friedmann universe with both matter and a cosmological constant being that entirely in the early universe itself  $\Omega_M(t)$  should go from one to zero and  $\Omega_\Lambda(t)$  should go from zero to one, the closeness of their values some ten billion years later is often referred to as the cosmic coincidence problem.

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<sup>45</sup>In Fig. 2 we have labelled this fit as the conformal gravity fit, as it will be associated with the conformal gravity alternative to standard gravity which will be discussed below. However, for the moment it need only be regarded as just one of the possible phenomenological fits obtainable by treating  $\Omega_M(t_0)$  and  $\Omega_\Lambda(t_0)$  as free parameters in Eq. (138).

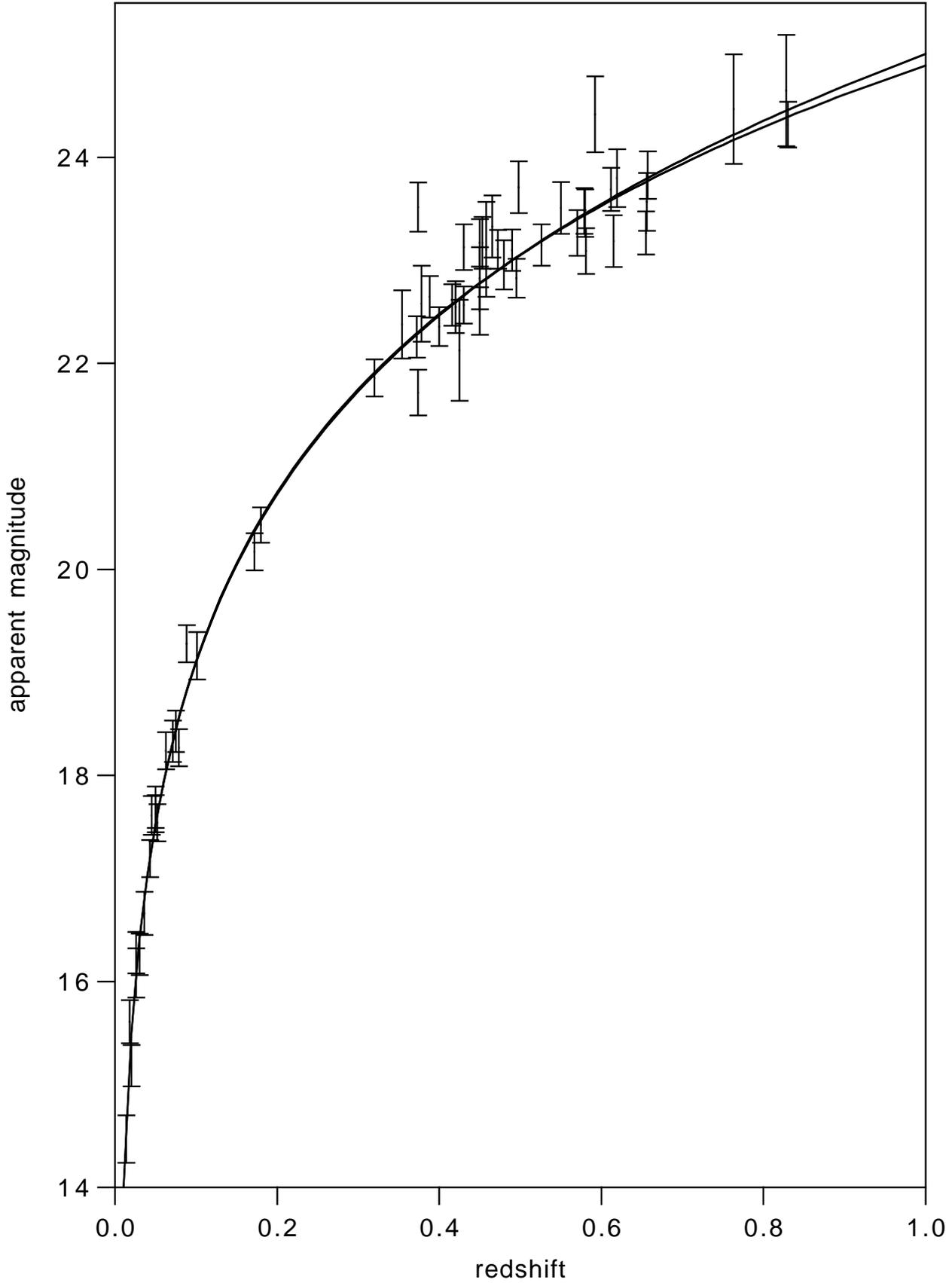


Figure 2: The  $q_0 = -0.37$  conformal gravity fit (upper curve) and the  $\Omega_M(t_0) = 0.3$ ,  $\Omega_\Lambda(t_0) = 0.7$  standard model fit (lower curve) to the  $z < 1$  supernovae Hubble plot data.

While inflation does not solve this cosmic coincidence problem, an early universe inflationary era can nonetheless still occur (and indeed its occurrence would still solve the horizon problem), with  $\Omega_k(t)$  then still being negligible throughout the Robertson-Walker era. Within the class of best fit solutions to the supernovae data this would then single out the  $\Omega_M(t_0) = 0.3$ ,  $\Omega_\Lambda(t_0) = 0.7$  one, and thus bring inflation into agreement with the  $\Omega_M(t_0) = 0.3$  dark matter cluster estimate. Further support for this  $\Omega_M(t_0) = 0.3$ ,  $\Omega_\Lambda(t_0) = 0.7$  realization of inflation has recently been obtained from study of the temperature anisotropy of the cosmic microwave background, since such anisotropies can actually be generated in a period of rapid early universe inflation in which initially microscopic fluctuations could be amplified to macroscopic size. The size to which such inflationary fluctuations will grow at the time of recombination will thus act as a theoretical standard yardstick which will be imprinted on the cosmic microwave background, with the apparent size with which it will then appear to us then reflecting the geometry through which photons have had to travel to us since recombination. Through use of this technique the cosmic microwave background anisotropy measurements are then found to have a best fit in the  $(\Omega_M(t_0), \Omega_\Lambda(t_0))$  plane which scatters around the line [4, 5]

$$\Omega_\Lambda(t_0) = 1 - \Omega_M(t_0) \quad , \quad \Omega_k(t_0) = 0 \quad (144)$$

with the compatibility of Eq. (141) with Eq. (144) remarkably bringing us right back to  $\Omega_M(t_0) = 0.3$ ,  $\Omega_\Lambda(t_0) = 0.7$ . Consequently, these values are now regarded as being the ones both required of and established by the standard cosmological theory, with the standard theory thus having achieved its primary objective of determining the matter content of the universe.

However, even while the standard theory has now ostensibly done its job, and while one should not understate how remarkable it is that there actually is a specific single set of  $\Omega_M(t_0)$ ,  $\Omega_\Lambda(t_0)$  values which fits currently available cosmological data at all,<sup>46</sup> the universe which emerges is still a highly perplexing one which not only possesses huge amounts of still poorly understood dark matter, it needs to augment this dark matter with an even more poorly understood  $\Omega_\Lambda(t_0)$ . However, before looking at the implications of such a value for  $\Omega_\Lambda(t_0)$ , it is important to note that there could be an alternative to it within standard gravity, namely the so-called quintessence model of [15]. The quintessence model was actually developed prior to the discovery of the accelerating universe, with its objective at the time being to try to reconcile an  $\Omega_M(t_0)$  which would be less than one with inflation in a way that would not require any early universe fine-tuning of the Friedmann equation. To this end a quintessence source was added to the dark matter source, one which would then have an  $\Omega_Q(t_0)$  which would be equal to  $1 - \Omega_M(t_0)$  today, and thus be of the same order as  $\Omega_M(t_0)$  today, to initially appear to replace the  $\Omega_M(t_0)$ ,  $\Omega_\Lambda(t_0)$  cosmic coincidence by an  $\Omega_M(t_0)$ ,  $\Omega_Q(t_0)$  cosmic coincidence instead. However, it was found possible to construct a tracking potential for the quintessence fluid which would have the effect of driving  $\Omega_Q(t)$  and  $\Omega_M(t)$  to comparable current era values today regardless of what early universe initial conditions were presupposed, to thus render the cosmic evolution equation free of early universe fine-tuning problems. In order for the quintessence fluid to be able to achieve this objective, it would have to have an equation of state in which  $p/\rho = w$  is negative, and in order for it to have escaped visual detection it would need to be non-luminous. Non-luminous negative pressure fluids such as these are known as dark energy,<sup>47</sup> with a cosmological constant being a special case of such fluids, one with  $w = -1$ . With quintessence fluids having negative  $w$ , they can thus lead to cosmic repulsion, and if  $w$  is negative enough can provide for fitting to the supernovae data of the same quality as the  $\Omega_M(t_0) = 0.3$ ,  $\Omega_\Lambda(t_0) = 0.7$  cosmology fit shown in Fig. 1. While quintessence is thus offered as a candidate solution to the cosmic coincidence problem, it leaves open the question of what is to happen to the cosmological

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<sup>46</sup>While it might be thought that the lines of Eqs. (141) and (144) have to cross somewhere, this is not quite the case since the lines are not actually of infinite extent, with both of them only describing their respective data in limited regions, regions which remarkably do in fact overlap.

<sup>47</sup>Since it is their negative pressure which is abnormal and not their energy density (which remains positive), a possibly better designation for quintessence fluids would be dark pressure.

constant, with quintessence only being able to succeed if some mechanism can be found which would naturally set the cosmological constant to zero. For quintessence models to be valid then still requires a solution to the cosmological constant problem.

## 7.6 The cosmological constant problem

Indeed the real problem with the cosmological constant is that rather than provide an  $\Omega_\Lambda(t_0)$  as small as 0.7 or even one equal to zero, the a priori expectation for  $\Omega_\Lambda(t_0)$  is that it will be huge, being as large as  $10^{120}$  if fixed by the quantum gravity Planck scale, and being as large as  $10^{60}$  or so if fixed by the electroweak phase transition scale, with use of either of these values leading to absolute phenomenological disaster for the standard theory. Moreover, even if  $\Omega_\Lambda(t)$  is quenched to zero at the end of the inflationary era, as the universe then cools down it will go through phase transitions such as the electroweak one at around  $T_V = 10^{15}$  degrees, at which point a brand new cosmological constant will be generated by the change in free energy due to the phase transition, with the difference between the free energy just above the phase transition and its value at current temperatures way below it being of order  $\sigma T_V^4$  where  $\sigma$  is Stefan's constant. Not only is such a  $\sigma T_V^4$  huge, since the free energy is lowered in a phase transition, one would even expect it to produce a negative contribution to  $\Lambda$ , with  $\Lambda$  then differing both in sign and magnitude from the required scale of the order of few degrees which is associated with the positive  $\Omega_\Lambda(t_0) = 0.7$ . Apart from the electroweak scale, there might even be a supersymmetry breaking scale as well, and while supersymmetry could provide some possible dark matter wimp candidates, the very same mechanism which would put their masses into at least the TeV region (the current lower mass bound associated with their lack of detection to date) would at the same time generate an even bigger contribution to the cosmological constant than that associated with electroweak symmetry breaking itself. Thus mechanisms which will generate the cold dark matter particles needed for  $\Omega_M(t_0) = 0.3$  will also generate huge contributions to  $\Omega_\Lambda(t_0)$  in a dynamics which would leave us nowhere near the cosmic coincidence values at all. This then is the cosmological constant problem.

At the present time within the standard model no convincing resolution of the cosmological constant problem has been identified, and while there are some suggestions, none of them are anywhere near being at the stage where they could be falsified. Enormous efforts have been made to find a dynamics which would naturally quench  $\Omega_\Lambda(t_0)$  down from either  $10^{60}$  or  $10^{120}$  to a current value of 0.7, so far to no avail. Additionally, one can consider the possibility that in addition to the overall dynamical cosmological constant which is generated on the right-hand side of the Einstein equations as the matter density goes through its various phase transitions, there could also exist an a priori fundamental cosmological constant on the left-hand side, one which will then almost, but not quite cancel the dynamical one so that the current era observer then sees a net  $\Omega_\Lambda(t_0) = 0.7$ . Such a proposal requires that a fundamental cosmological constant exist as part and parcel of the fundamental gravitational equations, that it be included in gravitational theory in all epochs, and that it be preset with just the value needed to lead to a net  $\Omega_\Lambda(t_0) = 0.7$  today. It is not clear how one might test such an idea, with such a cancellation having to be regarded as being a cosmic coincidence itself.

Another proposed mechanism is the so-called anthropic principle which takes the view that there exists an infinite number of universes in which  $\Lambda$  takes on all possible values, with observers such as ourselves only being able to exist and observe the universe in those particular universes in which the value of  $\Lambda$  permits the emergence of intelligent beings. In a sense such a viewpoint is actually a denial of physics, saying that when we can explain something we do, and when we cannot, we appeal to the anthropic principle.<sup>48</sup> From a more technical viewpoint it may be the case that even with an infinite number of universes, the cosmological constant could be Planck density in all of them, and that

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<sup>48</sup>A possibly more accurate name might be the anthropic lack of principle, with it not being clear which phenomena are to be attributed to the anthropic principle and which not.

while the Planck length itself might vary from universe to universe, the relation of  $\Lambda$  to it might not. Additionally, it would be an extraordinary coincidence if the only possible allowed range of  $\Lambda$  needed for intelligent beings would precisely be within the range needed to solve the early universe Friedmann equation fine-tuning problem. Rather, one might expect a somewhat broader range for  $\Lambda$ , and with the Friedmann fine-tuning required range (viz. to one part in  $10^{60}$ ) then being a very small fraction of it, it is far more likely that the anthropic principle would put us outside the Friedmann required range rather than in it. Finally, if the anthropic principle is to be a part of physics, then there has to be some falsifiable prediction that it makes which would enable us to test it.

## 7.7 Testing dark energy beyond a redshift of one

However, regardless of any of these theoretical concerns, there is actually a direct, model-independent test that one can make to see whether it really is true that the real world  $c\Lambda$  is indeed equal to  $0.7\rho_c \sim 0.6 \times 10^{-8}$  ergs/c.c. (viz.  $\alpha = \Omega_\Lambda H^2/c^2 = 8\pi G\Lambda/3c^3 \sim 0.4 \times 10^{-56}/\text{cm}^2$ ), this being the value which would make  $\Omega_\Lambda(t_0)$  be equal to 0.7 today. Specifically, since  $\rho_m$  redshifts while  $\Lambda$  does not,  $\rho_m$  would then dominate over  $\Lambda$  at earlier times, with the universe then having to be in a decelerating phase rather than an accelerating one. For the given values of  $\Omega_M(t_0)$  and  $\Omega_\Lambda(t_0)$  such a switch between acceleration and deceleration would occur before  $z = 1$ , and so a monitoring of the Hubble plot above  $z = 1$  would then test for such a switch over. As such, the very existence of a switch over near  $z = 1$  is a reflection of the cosmic coincidence, with initial conditions having to be fine-tuned in the early universe so that the switch from  $\rho_m$ -dominated deceleration to  $\Lambda$ -dominated acceleration does not take place in the early universe itself but is delayed for no less than ten billion years. In Fig. 3 we have therefore plotted the higher  $z$  Hubble plot required of an  $\Omega_M(t_0) = 0.3$ ,  $\Omega_\Lambda(t_0) = 0.7$  universe, and if the standard model numbers are right, the luminosity function must be found to follow the indicated curve.

To get a sense of the level of precision which is required to make this test, in Fig. 3 we have also plotted the expectations of the  $(\Omega_M(t_0), \Omega_\Lambda(t_0)) = (0, 0)$  universe with  $q_0 = 0$  and the pure  $\Lambda$   $(\Omega_M(t_0), \Omega_\Lambda(t_0)) = (0, 0.37)$  universe with  $q_0 = -0.37$  which we discussed earlier, since the first of these two universes represents the coasting borderline case between acceleration and deceleration, while the latter represents the best fit permanently accelerating pure  $\Lambda$  universe.<sup>49</sup> Discriminating between the three cases (accelerating/decelerating, coasting, and permanently accelerating) should eventually be achievable since at  $z = 2$  these three typical universes respectively yield apparent magnitudes  $m = 26.75$ , 27.04 and 27.17, while at  $z = 5$  they yield  $m = 29.14$ , 30.25 and 30.40, and should thus enable us to determine what sort of phase the universe might be in at redshifts above one.

Now while full discrimination between these various options would require obtaining as many data points above (ideally well above)  $z = 1$  as there are below  $z = 1$ , it is of interest to note that a few  $z > 1$  data points have recently been obtained [40, 41] including one (SN 1997ff) as high as  $z = 1.75$ . Of these new data points nine have been identified as being particularly reliable and so we have extended our fits to include them. In order to perform the fits, we note that in [2] the 54  $z < 1$  data points are presented as apparent magnitudes  $m$ , while in [41] the nine new  $z > 1$  data points are presented as distance moduli  $m - M$ . With these two methods of characterizing the data being related via  $m = M + 25 + 5\log_{10}d_L$  where  $d_L$  is the luminosity distance as measured in megaparsecs, in the fitting of the 54  $z < 1$  data points of [2]  $M$  is treated as a free parameter, with the notion that the supernovae are indeed standard candles then being confirmed by the fact that in the very good fits of Fig. 2 a common value for  $M$  actually is obtained. On now fitting the combined 63 (=54+9) data points in fits which allow  $M$

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<sup>49</sup>For the moment our interest in these particular cosmologies is only in having cosmologies which are not decelerating above  $z = 1$  so as to provide a contrast with the standard cosmology in that region. However, shortly we shall see that such non-decelerating cosmologies are of interest in their own right, as they are relevant to the alternate conformal gravity solution to the cosmological constant problem which we present below.

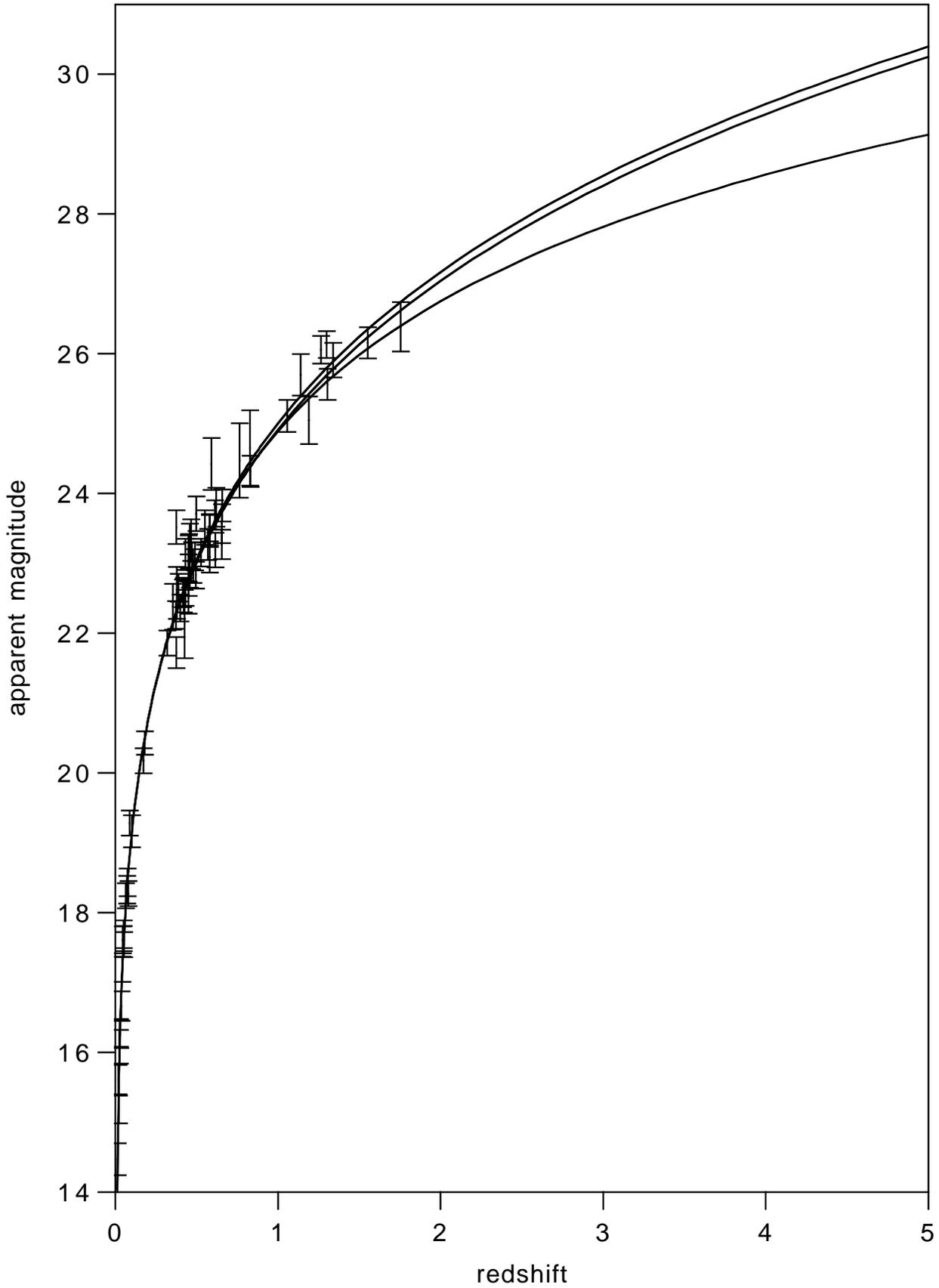


Figure 3: Hubble plot expectations for  $q_0 = -0.37$  (highest curve) and  $q_0 = 0$  (middle curve) conformal gravity and for  $\Omega_M(t_0) = 0.3$ ,  $\Omega_\Lambda(t_0) = 0.7$  standard gravity (lowest curve).

to again vary for the 54 points in the sample, the standard  $(\Omega_M(t_0), \Omega_\Lambda(t_0)) = (0.3, 0.7)$  universe is found to yield  $\chi^2 = 74.5$  and  $M = -19.15$ , with the best  $\Omega_k(t_0) = 0$  universe fit being found to be the slightly lower mass (and thus more cosmically repulsive)  $(\Omega_M(t_0), \Omega_\Lambda(t_0)) = (0.21, 0.79)$  universe with  $\chi^2 = 69.9$  and  $M = -19.19$ . For comparison, the pure  $\Lambda$   $(\Omega_M(t_0), \Omega_\Lambda(t_0)) = (0, 0.37)$  universe with  $q_0 = -0.37$  yields  $\chi^2 = 70.7$  and  $M = -19.14$ , while the pure curvature  $(\Omega_M(t_0), \Omega_\Lambda(t_0)) = (0, 0)$  universe with  $q_0 = 0$  yields  $\chi^2 = 78.4$  and  $M = -19.06$ . As we see then, while a standard  $\Omega_M(t_0) > 0$ ,  $\Omega_\Lambda(t_0) = 1 - \Omega_M(t_0)$  cosmology with  $q(z > 1) > 0$  and an  $\Omega_M(t_0) = 0$ ,  $\Omega_\Lambda(t_0) = 0.37$  cosmology with  $q(z > 1) < 0$  both outperform the  $\Omega_M(t_0) = 0$ ,  $\Omega_\Lambda(t_0) = 0$  cosmology with  $q(z > 1) = 0$  in the  $z > 1$  region, statistically the  $q(z > 1) > 0$  and  $q(z > 1) < 0$  cosmologies are equivalent to each other, with the current  $z > 1$  region data being just as well fitted by a model which is accelerating above  $z = 1$  as by one which is decelerating above it. Consequently, for the moment it is not possible to determine whether the universe is in an accelerating or a decelerating phase above  $z = 1$ , with both options being compatible with currently available data. To be able to display the nine  $z > 1$  data points of [41] on the apparent magnitude plot given in Fig. 3, we have taken the fitted value of  $M$  ( $= -19.14$ ) and used it to extract values for  $m$  for the nine points from their reported  $m - M$  distance moduli, with it being these inferred  $m$  values which are then displayed in the figure. While the  $\chi^2$  fitting itself is not at all sensitive to the value used for  $H(t_0)$  which is required for the luminosity distance  $d_L$ , the extraction of an actual value for  $M$  itself from the fitting is. For instance with  $H(t_0) = 72$  km/sec/Mpc, in megaparsec units the quantity  $5\log_{10}(c/H(t_0))$  evaluates to 18.099, while for  $H(t_0) = 65$  km/sec/Mpc it evaluates to 18.321, a quite substantial difference of 0.222. With an increasing value for the Hubble parameter leading to a decreasing  $5\log_{10}(c/H(t_0))$ , an increasing value for the Hubble parameter thus also leads to an increasing derived  $M$ . In its turn an increased  $M$  leads to an increased (dimmer) value for  $m$  as extracted from a given distance modulus. On making this extraction, as we see from Fig. 3, the extracted  $z > 1$   $m$  values then precisely straddle the three indicated theoretical curves, to thereby prevent current data from being able to differentiate between acceleration and deceleration in the  $z > 1$  region.<sup>50</sup> With the SN 1997ff supernova at  $z = 1.75$  being known to be lensed by two foreground galaxies which lie at  $z = 0.5$  [42], this particular supernova may actually be dimmer (larger  $m$ ) than it appears. Filling in the  $z > 1$  Hubble plot with a large enough sample of data points and allowing for systematic effects such as lensing should thus be very informative.

Apart from the use of type Ia supernovae to explore the  $z > 1$  region, a few other techniques have also been developed, with Daly and coworkers having used the very powerful extended FR II radio sources as standard yardsticks [43, 44] and Schaefer having explored the use of gamma ray bursters as standard candles [45]. For the FR II approach some 20 radio sources which go out to  $z = 2$  have already been fully analyzed (11 being above  $z = 1$ ), with there being some more which go out to  $z = 3$  or so, while the gamma ray burster technique will be able to go out to  $z = 4.5$ . Global fits to the 54  $z < 1$  supernovae studied above and the 20 FR II radio galaxies have been made, with the  $(\Omega_M(t_0), \Omega_\Lambda(t_0)) = (0.3, 0.7)$  standard cosmology giving  $\chi^2 = 74.4$  for the 74 total data points and an  $(\Omega_M(t_0), \Omega_\Lambda(t_0)) = (0, 0.38)$  cosmology giving  $\chi^2 = 74.1$  [46]. The available radio source data are thus seen to be just as compatible with both decelerating and accelerating universes in the  $z > 1$  region as the supernovae data are. Further exploration of the  $z > 1$  Hubble plot using all of the available techniques will thus be necessary in order to determine how the universe behaves in the  $z > 1$  region and ascertain whether in fact  $\Omega_\Lambda(t_0)$  really is 0.7 in the real world.

Having now described and diagnosed the dark matter and dark energy problems, we turn next to an analysis of some alternatives to the standard model in which these problems are resolved. And since the cosmological dark energy problem is an intrinsically general-relativistic one, to seek alternatives to dark matter and dark energy will oblige us to have to consider alternatives to Einstein gravity itself.

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<sup>50</sup>The  $z > 1$  data are however able to exclude the possibility that dust or cosmic evolution may be causing a dimming of the  $z < 1$  supernovae, since both of those effects would dim the  $z > 1$  supernovae even more, an effect which as noted in [41] the  $z > 1$  data now exclude, with the behavior of the Hubble plot thus indeed being of cosmological origin.

## 8 Alternatives to Einstein gravity

### 8.1 Pure metric theories with additional curvature dependent terms

As we had noted earlier, while it is sufficient to use standard Einstein gravity to obtain the standard Schwarzschild metric phenomenology for the solar system, is not in fact necessary. Thus in seeking alternatives to Einstein gravity, as well as look for generalizations of it which contain the Einstein equations in an appropriate limit, one can also look for generalizations which are able to recover the Schwarzschild phenomenology while never reducing to the Einstein theory in any limit at all, with the more general rule thus being to generalize the Schwarzschild solution rather than the Einstein equations themselves. We shall thus look at both kinds of generalization (either reducing to the Einstein equations or bypassing them altogether) and after detailing possible options, shall then look for a fundamental principle which would enable us to unambiguously select just one option from amongst them.

Within the framework of general covariant pure metric theories of gravity, the most straightforward generalization of the standard Einstein theory is to augment the Einstein-Hilbert action of Eq. (103) with additional general coordinate invariant pure metric terms, additional terms the smallness of whose coefficients or the specific structure of which (vanishing in Ricci flat geometries) would cause them to make negligible contributions in the solar system. The choice of possible additional terms is actually unlimited, since it not only encompasses appropriate contractions of arbitrarily high powers of the Riemann tensor (such as the fourth order actions of Eqs. (105) and (106)),<sup>51</sup> provided candidate additions are able to avoid violating locality, in principle such additions need not even be expressible as a power series expansion of Riemann tensor terms at all.<sup>52</sup> Moreover, within the class of additions to the Einstein-Hilbert action, one must also include the addition that Einstein himself made, viz. a cosmological constant term, as well. Without a fundamental guiding principle almost nothing can be said either in favor or against any candidate such addition, with this freedom being due to the fact that it was not an appeal to a fundamental principle which led to the specific choice of the Einstein-Hilbert action in the first place. While it is sometimes stated that the choice of the Einstein-Hilbert action can be justified on grounds of simplicity, simplicity itself is not a law of nature, and the simplicity of the Einstein approach to gravity is not so much in the simplicity of its equations (or the ease with which they might be solved) but in the simplicity of its concepts. For Einstein, the simplicity of the concepts and the intrinsic elegance and beauty of the theory were paramount, and one could argue that in adding in dark matter and dark energy in the ad hoc way that is currently being done, one gives up much of the inner beauty of the theory.

### 8.2 Additional fields

Within the framework of general covariant theories of gravity, the next most straightforward generalization of Einstein gravity is to introduce additional macroscopic gravitational fields to accompany the metric tensor itself, with the most common choice being additional scalar fields. In and of itself, the addition of covariantly coupled scalar fields does not imperil the validity of the equivalence principle, since instead of coupling through the  $I_T$  action of Eq. (6), a test particle will now couple through the  $I_T$  action as augmented by the  $I_T^{(4)}$  action of Eq. (11), to lead to a particle equation of motion

$$(mc + \hat{\kappa}S) \frac{D^2 x^\lambda}{D\tau^2} = (mc + \hat{\kappa}S) \left( \frac{d^2 x^\lambda}{d\tau^2} + \Gamma_{\mu\nu}^\lambda \frac{dx^\mu}{d\tau} \frac{dx^\nu}{d\tau} \right) = -\hat{\kappa}S_{;\beta} \left( g^{\lambda\beta} + \frac{dx^\lambda}{d\tau} \frac{dx^\beta}{d\tau} \right) \quad (145)$$

<sup>51</sup>In a quantization of the standard Einstein theory an infinite set of such higher order terms is generated via radiative correction counter terms – though these terms then come with Planck scale coefficients which are indeed negligible compared to the  $\hbar = 0$  Einstein term itself.

<sup>52</sup>The addition to the Einstein-Hilbert action of terms which behave as  $1/R^\alpha$  has been considered by [47] in an attempt to explain the acceleration of the universe as a gravitational effect rather than one which is generated by dark energy.

which still involves the covariant 4-acceleration  $D^2x^\lambda/D\tau^2$ . While Eq. (145) and its Eq. (12) antecedent can lead to departures from geodesic motion (an issue we return to below when we consider covariant generalizations of MOND), it still respects the equivalence principle requirement that the Christoffel symbol term and the ordinary acceleration term  $d^2x^\lambda/d\tau^2$  always appear in exactly the combination indicated. Within the class of scalar tensor theories the most venerable is the Brans-Dicke theory with scalar tensor sector action [48]

$$I_{BD} = - \int d^4x (-g)^{1/2} \left( SR^\alpha{}_\alpha - w \frac{S_{;\mu} S^{;\mu}}{S} \right) \quad (146)$$

where  $S$  is the scalar field and  $w$  is a constant.<sup>53</sup> In this theory there is no fundamental  $G$ , with  $G$  instead being induced as an effective gravitational constant in solutions to the theory in which  $S$  is non-zero. With  $S$  being able to take different values in different epochs, the Brans-Dicke theory therefore allows for a varying effective  $G$ , and with it then being this epoch-dependent effective  $G$  which will then normalize  $\Omega_M$  in each epoch, while not solving the Friedmann equation fine-tuning problem the Brans-Dicke theory does at least allow one to reformulate it as a need to find a dynamics in which  $S$  could naturally change by 60 orders of magnitude from the early universe until today. While such a mechanism is a general one which actually holds for any theory in which an effective  $G$  is induced by a scalar field, it turns out that solar system constraints on the Brans Dicke theory itself are so tight, that there is no room to choose its parameter  $w$  in a way which could lead to any potential departure from the standard theory. While this is thus a shortcoming of the Brans-Dicke theory, it is not actually a generic shortcoming of all scalar tensor theories of gravity, only a shortcoming of the scalar tensor theory which is based on the particular action of Eq. (146), with every individual candidate scalar tensor theory of gravity needing to be evaluated on a case by case basis.

### 8.3 Additional spacetime dimensions

The third kind of generalization of the standard Einstein theory is one which has now become extremely popular and widespread, namely increasing the number of spacetime dimensions. Such ideas originated with Kaluza and Klein in the nineteen twenties who considered the possible existence of a fifth dimension, with the concept of extra dimensions reemerging much later following the development of superstring theories. Characteristic of these higher dimensional approaches is not just the existence of extra dimensions, but also of extra fields. The Kaluza-Klein approach is based on a 5-dimensional metric, a 15-component quantity whose 4-dimensional decomposition consists of a 10-component rank two tensor graviton, a 4-component vector field, and a 1-component scalar field, while the string theory approach typically involves scalar dilaton fields and not just rank two tensor gravitons.<sup>54</sup> To account for the apparent absence of any dimensions beyond four, it had generally been thought that all but four of the higher dimensions would be compactified down to Planck length size. More recently though it was recognized by Randall and Sundrum [8, 9] that if we lived on a brane (viz. membrane) which was embedded in a higher dimensional anti-de Sitter bulk, the anti de Sitter bulk could (for at least some choices of brane matter fields) inhibit the propagation of gravitational signals into it no matter how large in size it might be.

While these higher dimensional theories all contain a massless graviton, it does not follow that a 4-dimensional observer confined to the brane will then see standard 4-dimensional Einstein gravity on it, with the gravity on the brane being modified by the very embedding itself, a modification which actually opens up a possible new way to address the dark matter and dark energy problems. In order to see how such modifications come about, we note first that before imposing any dynamics whatsoever,

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<sup>53</sup>With its  $1/S$  dependence the scalar field kinetic energy is not expressible as a power series expansion in  $S$ .

<sup>54</sup>String theory typically also generates Planck scale higher order Riemann tensor action terms.

for the Gaussian normal metric

$$ds^2 = g_{MN}dx^M dx^N = dw^2 + q_{\mu\nu}dx^\mu dx^\nu \quad (147)$$

considered in the typical 5-dimensional bulk case,<sup>55</sup> use only of the purely geometric Gauss embedding formula entails that the induced Einstein tensor  ${}^{(4)}G_{\mu\nu}$  on the brane (viz. the one calculated with the induced metric  $q_{\mu\nu}$  seen by an observer at  $w = 0$ ) is given by<sup>56</sup>

$$\begin{aligned} {}^{(4)}G_{\mu\nu} &= \frac{2}{3}G_{\alpha\beta}q^\alpha{}_\mu q^\beta{}_\nu + \frac{2}{3}G_{55}q_{\mu\nu} - \frac{1}{6}G^A{}_A q_{\mu\nu} - E_{\mu\nu} \\ &\quad - K^\alpha{}_\alpha K_{\mu\nu} + K^\alpha{}_\mu K_{\alpha\nu} + \frac{1}{2}(K^\alpha{}_\alpha)^2 q_{\mu\nu} - \frac{1}{2}K_{\alpha\beta}K^{\alpha\beta} q_{\mu\nu} \end{aligned} \quad (148)$$

where  $G_{MN}$  is the 5-dimensional Einstein tensor (viz. the one calculated with the full 5-dimensional metric  $g_{MN}$ ),  $E_{\mu\nu}$  is the particular projection

$$E_{\mu\nu} = q^\beta{}_\mu q^\delta{}_\nu C^5{}_{\beta 5 \delta} \quad (149)$$

of the 5-dimensional Weyl tensor  $C_{ABCD}$ , and  $K_{\mu\nu}$  (the  $(\mu\nu)$  component of the extrinsic curvature  $K_{MN} = (g^A{}_M - n^A n_M)(g^B{}_N - n^B n_N)n_{B;A}$  of the brane) is given by

$$K_{\mu\nu} = q^\alpha{}_\mu q^\beta{}_\nu n_{\beta;\alpha} \quad (150)$$

Similarly, in terms of the acceleration vector  $a_N = n^A n_{N;A}$  which lies in the brane ( $n^N a_N = 0$  when  $n^N n_N = 1$ ), one can derive a second purely geometric relation

$$n^C K_{\mu\nu;C} - R_{BD}q^B{}_\mu q^D{}_\nu = -{}^{(4)}R_{\mu\nu} + \frac{1}{2}(a_{\mu;\nu} + a_{\nu;\mu}) - K^A{}_A K_{\mu\nu} - a_\mu a_\nu \quad (151)$$

Consequently, no matter what form we may assume for the brane world energy-momentum tensor  $T_{MN}$ , we immediately see that even if we do take  $G_{MN}$  to obey the Einstein equations in the higher dimensional bulk space, viz.

$$G_{MN} = -\kappa_5^2 T_{MN} \quad (152)$$

it does not at all follow that  ${}^{(4)}G_{\mu\nu}$  has to then obey the standard Einstein equations in the 4-space.

To see what equation  ${}^{(4)}G_{\mu\nu}$  does in fact obey, we note that if we take the bulk to possess a 5-dimensional cosmological constant  $\Lambda_5$  and the brane to possess some energy-momentum tensor  $\tau_{\mu\nu}$  which is confined to it, viz.

$$T_{MN} = -\Lambda_5 g_{MN} + \tau_{\mu\nu} \delta_M^\mu \delta_N^\nu \delta(w) \quad (153)$$

then since every term on the right-hand side of Eq. (151) is continuous at the brane<sup>57</sup> it follows that the discontinuities on the left-hand side must cancel,<sup>58</sup> so that on crossing the brane we obtain

$$\int_{0^-}^{0^+} dw \left[ \partial_w K_{\mu\nu} + \kappa_5^2 \left( \tau_{\mu\nu} - \frac{1}{3} g_{\mu\nu} \tau^\alpha{}_\alpha \right) \delta(w) \right] = 0 \quad (154)$$

<sup>55</sup>Here we take  $w$  to denote the fifth coordinate,  $M$  and  $N$  to denote 5-dimensional coordinates and  $\mu$  and  $\nu$  to denote the usual four coordinates on the brane, and have chosen the geometry so that the brane is located at  $w = 0$  and the normal to the brane is given by  $n_M = (0, 0, 0, 0, 1)$ . In this coordinate system even though the components of the induced metric  $q_{\mu\nu}$  only range over the 4-space indices,  $q_{\mu\nu}$  is nonetheless a function of all five coordinates.

<sup>56</sup>See e.g. [49] where a more detailed treatment of the embedding issues and formulae discussed here may be found.

<sup>57</sup>Both  ${}^{(4)}R_{\mu\nu}$  and  $a_\mu$  lie in the plane of the brane, and even though the extrinsic curvature has a step function discontinuity at the brane, terms quadratic in it do not.

<sup>58</sup>Since  $K_{\mu\nu}$  is a step function, its covariant derivative is a delta function.

with the fully covariant Israel junction conditions [50]

$$K_{\mu\nu}(w = 0^+) - K_{\mu\nu}(w = 0^-) = -\kappa_5^2 \left[ \tau_{\mu\nu} - \frac{1}{3} q_{\mu\nu}(w = 0) \tau^\alpha_\alpha \right] \quad (155)$$

then emerging.<sup>59</sup> For the brane world set-up a  $Z_2$  symmetry about the brane is imposed so that the extrinsic curvatures on the two sides of the brane are related by  $K_{\mu\nu}(w = 0^+) = -K_{\mu\nu}(w = 0^-)$ , with the extrinsic curvature at the brane then being given by

$$K_{\mu\nu}(w = 0^+) = -K_{\mu\nu}(w = 0^-) = -\frac{\kappa_5^2}{2} \left[ \tau_{\mu\nu} - \frac{1}{3} q_{\mu\nu}(w = 0) \tau^\alpha_\alpha \right] . \quad (156)$$

Given this discontinuity and the original 5-dimensional Einstein equations of Eq. (152) which fix the value of  $G_{MN}$  at the brane, the extraction of the continuous piece of Eq. (148) is then direct and yields [51]

$${}^{(4)}G_{\mu\nu} = \frac{1}{2} \kappa_5^2 \Lambda_5 q_{\mu\nu}(w = 0) - \kappa_5^4 \Pi_{\mu\nu} - \bar{E}_{\mu\nu}(w = 0) \quad (157)$$

where

$$\bar{E}_{\mu\nu}(w = 0) = \frac{1}{2} [E_{\mu\nu}(w = 0^+) + E_{\mu\nu}(w = 0^-)] \quad (158)$$

and

$$\Pi_{\mu\nu} = -\frac{1}{4} \tau_{\mu\alpha} \tau_\nu^\alpha + \frac{1}{12} \tau^\alpha_\alpha \tau_{\mu\nu} + \frac{1}{8} q_{\mu\nu}(w = 0) \tau_{\alpha\beta} \tau^{\alpha\beta} - \frac{1}{24} q_{\mu\nu}(w = 0) (\tau^\alpha_\alpha)^2 . \quad (159)$$

Finally, on introducing a brane cosmological constant  $\lambda$  and decomposing the brane  $\tau_{\mu\nu}$  as

$$\tau_{\mu\nu} = -\lambda q_{\mu\nu}(w = 0) + S_{\mu\nu} \quad (160)$$

where  $S_{\mu\nu}$  represents all other brane matter field sources, we find that Eq. (159) can then be rewritten as

$${}^{(4)}G_{\mu\nu} = \Lambda_4 q_{\mu\nu}(w = 0) - \kappa_4^2 S_{\mu\nu} - \kappa_5^4 \pi_{\mu\nu} - \bar{E}_{\mu\nu}(w = 0) \quad (161)$$

where

$$\begin{aligned} \pi_{\mu\nu} &= -\frac{1}{4} S_{\mu\alpha} S_\nu^\alpha + \frac{1}{12} S^\alpha_\alpha S_{\mu\nu} + \frac{1}{8} q_{\mu\nu}(w = 0) S_{\alpha\beta} S^{\alpha\beta} - \frac{1}{24} q_{\mu\nu}(w = 0) (S^\alpha_\alpha)^2 , \\ \kappa_4^2 &= \frac{\kappa_5^4 \lambda}{6} , \quad \Lambda_4 = \frac{\kappa_5^2}{12} (\kappa_5^2 \lambda^2 + 6\Lambda_5) = \frac{1}{2} (\kappa_4^2 \lambda + \kappa_5^2 \Lambda_5) . \end{aligned} \quad (162)$$

As constructed, Eq. (161) thus represents the generalized Einstein equation seen by an observer on the brane, with the  $\bar{E}_{\mu\nu}(w = 0)$  and  $\pi_{\mu\nu}$  terms heralding a departure from a standard 4-dimensional gravity with an effective gravitational constant  $\kappa_4^2$  and an effective cosmological constant  $\Lambda_4$ .<sup>60</sup> With regard to the effective cosmological constant, we note that it is not  $\lambda$  which is to serve as the cosmological constant for cosmology on the brane but rather  $\Lambda_4$ , and since the anti-de Sitter bulk  $\Lambda_5$  has to be negative, the effective  $\Lambda_4$  has to be smaller than the  $\lambda$  expected from ordinary 4-dimensional early universe phase transition physics. As we had noted earlier, a possible way to quench the contribution of the cosmological constant to cosmology is to have two cosmological constants, one being a dynamical one coming from the energy-momentum tensor (viz.  $\lambda$ ) and the second being an a priori fundamental one. As we now see, what we had thought of as a fundamental cosmological constant from the point of view of 4-dimensional physics might actually be a  $\kappa_5^2 \Lambda_5 / 2$  term coming from the embedding in a

<sup>59</sup>In a  $D$  dimensional bulk and a  $D$ -space brane the factor  $1/3$  is replaced by  $1/(D-2)$ .

<sup>60</sup>From the point of view of 4-dimensional physics, the massless 5-dimensional graviton of Eq. (152) decomposes into a massless 4-dimensional graviton and an infinite tower of massive 4-dimensional spin two modes [8, 9], with these massive modes then contributing to Eq. (161). As such, these massive modes could also serve as potential dark matter candidates.

higher dimensional bulk. While embedding thus does provide a natural rationale for the existence of the  $\kappa_5^2\Lambda_5/2$  term, at the present time no rationale has yet been found for why the  $\kappa_5^2\Lambda_5/2$  term and the  $\kappa_4^2\lambda/2$  terms might be as close to each other as an  $\Omega_\Lambda(t_0) = 0.7$  universe would require. Within this extra dimension picture we note that it has also been suggested [52] that the acceleration of the universe might be explainable not as a cosmological constant effect at all but rather as a consequence of a leakage of gravitational signal into the extra dimensions.

## 8.4 The modified Friedmann equations on the brane

To get a sense of the relevance of the  $\bar{E}_{\mu\nu}(w=0)$  and  $\pi_{\mu\nu}$  terms in Eq. (161), it is convenient to descend to the special case where  $S_{\mu\nu}$  is taken to be the perfect fluid

$$S_{\mu\nu} = (\rho + p)U_\mu U_\nu + pq_{\mu\nu}(w=0) \quad , \quad (163)$$

since in this case  $\pi_{\mu\nu}$  then also takes the form of the perfect fluid

$$\pi_{\mu\nu} = (R + P)U_\mu U_\nu + Pq_{\mu\nu}(w=0) \quad (164)$$

with energy density and pressure

$$R = \frac{\rho^2}{12} \quad , \quad P = \frac{1}{12}(\rho^2 + 2\rho p) \quad . \quad (165)$$

Considering now the cosmological case where  $q_{\mu\nu}(w=0)$  is taken to be a standard Robertson-Walker metric with scale factor  $a(t)$ , the (00) component of Eq. (161) is then found to take the form

$$-\frac{3(\dot{a}^2 + k)}{a^2} = -\Lambda_4 - \kappa_4^2\rho - \frac{\kappa_5^4\rho^2}{12} - \bar{E}_{00}(w=0) \quad , \quad (166)$$

with it thus being Eq. (166) which is to describe cosmic evolution on the brane, to be compared with the standard (unembedded) 4-dimensional Friedmann evolution equation

$$-\frac{3(\dot{a}^2 + k)}{a^2} = -\Lambda_4 - \kappa_4^2\rho \quad (167)$$

where the  $\kappa_5^2\rho^2/12$  and  $\bar{E}_{00}(w=0)$  terms are of course absent. Now while it is geometrically possible to embed a brane with a Robertson-Walker geometry into a 5-dimensional bulk with an anti-de Sitter geometry, it is not obligatory, with the bulk being able to have less than maximal 5-symmetry in general. Thus while the  $\bar{E}_{00}(w=0)$  term will vanish if the bulk is anti-de Sitter (since the bulk Weyl tensor then vanishes), it should not in general be expected to do so. Its non-vanishing would thus entail an unavoidable modification of standard cosmology whose impact on the dark matter and dark energy problems could be of interest. Moreover, even if the bulk geometry is such that the  $\bar{E}_{00}(w=0)$  term does in fact vanish, Eq. (166) will still differ from standard cosmology due to the presence of a term which is quadratic in the energy density  $\rho$ . Such terms will be of great importance when  $\rho$  is big, namely in the early universe, though in permanently expanding cosmologies such terms would redshift away much faster than the term which is linear in  $\rho$  and thus be nicely suppressed today. With the  $\rho^2$  term leading to early universe modifications, it would be of interest to see how they might modify the initial condition fine-tuning flatness problem.

While it might be thought that a late time embedded  $\bar{E}_{00}(w=0) = 0$  cosmology would always have to behave the same way as an unembedded one, this is not necessarily the case. Thus in the specific case where the effective  $\Lambda_4$  is negative, the Robertson-Walker 3-curvature  $k$  is zero and  $\bar{E}_{00}(w=0)$  is zero, Eq. (166) is found to admit of the exact radiation era ( $\rho = A/a^4$ ) solution

$$a^4(t) = -\frac{\kappa_5^2 Ab}{6H^2}\cos(4Ht) + \frac{\kappa_5^2 A(b^2 - H^2)^{1/2}}{6H^2} \quad (168)$$

where  $b^2 = -\kappa_5^2 \Lambda_5 / 6$ ,  $\Lambda_4 = -3H^2$ , and  $\kappa_4^2 = \kappa_5^2 (b^2 - H^2)^{1/2}$ . As we see from Eq. (168), in this particular model  $a(t)$  is not in fact late time expanding. Rather, it only expands to a bounded maximum value given by

$$a_{\max}^4(t) = \frac{\kappa_5^2 A}{6H^2} [b + (b^2 - H^2)^{1/2}] \quad . \quad (169)$$

With the ratio of the terms in Eq. (166) which are quadratic and linear in  $\rho$  being given by

$$\frac{\kappa_5^4 \rho^2}{12\kappa_4^2 \rho} = \frac{H^2}{2[b^2 - H^2 + b(b^2 - H^2)^{1/2}]} \quad (170)$$

at this maximum, we see that when  $H$  is of order  $b$ , the right-hand side of Eq. (170) is of then of order one. Since the  $\rho^2/\rho \sim 1/a^4$  ratio is at its minimum when  $a(t)$  is at its maximum, we thus see that in this particular cosmology at no time is the  $\rho^2$  term ever negligible compared to the  $\rho$  term, with this particular embedded cosmology never being able to approach the associated unembedded one at all. With Eqs. (148) and (151) being completely general purely geometric relations which will hold whenever there is a higher-dimensional embedding of any sort whatsoever (they do not even require the presence of a brane), despite the fact that higher dimensional models do possess a massless graviton, we see that such models could in principle yield modified 4-dimensional physics, and thus each candidate such higher-dimensional model has to be analyzed independently to see what kind of 4-dimensional physics it might lead to.

## 8.5 Relativistic MOND theory

Through the use of additional fields it is also possible to generalize Milgrom's non-relativistic MOND theory to the relativistic regime. The first step towards doing this was taken in [53] where a scalar field  $\psi$  was introduced with action

$$I(\psi) = -\frac{1}{8\pi GL^2} \int d^4x (-g)^{1/2} \tilde{f} \left( L^2 g^{\alpha\beta} \psi_{;\alpha} \psi_{;\beta} \right) \quad (171)$$

( $\tilde{f}$  is a scalar function and  $L$  is a constant), with this action then being added on to the standard Einstein-Hilbert one. Additionally, the scalar field coupling to a test particle was taken to be given by

$$I_m = -m \int e^\psi \left( -g_{\alpha\beta} dx^\alpha dx^\beta \right)^{1/2} \quad , \quad (172)$$

with the test particle thus obeying Eq. (12) as written with  $S = e^\psi$ . With the non-relativistic limit of  $\tilde{f}$  being taken to be the MOND function, the MOND formula of Eq. (96) was then found to ensue with the MOND acceleration parameter being identified as  $a_0 = 1/L$ . While this treatment shows what is in principle needed in order to recover MOND (or in general to produce a covariant theory whose non-relativistic limit is not the standard second order Poisson equation but a modified version of it), as formulated the above specific treatment had two difficulties. One was that  $\psi$  waves would propagate faster than light (the requisite  $\tilde{f}$  not being quadratic), and the other being that  $\psi$  would have no effect on the propagation of light waves (the light cone is invariant under the conformal transformation  $g_{\alpha\beta} \rightarrow e^{2\psi} g_{\alpha\beta}$ ) leaving the theory unable to explain gravitational lensing without dark matter. To resolve the acausality problem a second scalar field  $A$  was introduced [54, 55] with the action of Eq. (171) being replaced by

$$I(\psi, A) = -\frac{1}{2} \int d^4x (-g)^{1/2} \left[ \frac{A^2}{\eta^2} g^{\alpha\beta} \psi_{;\alpha} \psi_{;\beta} + g^{\alpha\beta} A_{;\alpha} A_{;\beta} + \nu(A^2) \right] \quad , \quad (173)$$

where  $\eta$  was a real parameter and  $\nu$  a real valued function. Building on this idea Bekenstein [56] then found a way to avoid the gravitational lensing problem as well in a generalization which involved

both scalar and vector fields. The theory which Bekenstein constructed contains the metric  $g_{\mu\nu}$ , a timelike 4-vector field  $U_\alpha$  which obeys the constraint  $g^{\alpha\beta}U_\alpha U_\beta = -1$ , a propagating scalar field  $\phi$  and a non-propagating scalar field  $\sigma$ . The action for the metric is taken to be the standard Einstein-Hilbert action, to which is added an action

$$I(\phi, \sigma) = -\frac{1}{2} \int d^4x (-g)^{1/2} \left[ \sigma^2 (g^{\alpha\beta} - U^\alpha U^\beta) \phi_{;\alpha} \phi_{;\beta} + \frac{G\sigma^4}{2\ell^2} F(kG\sigma^2) \right] \quad (174)$$

for the scalars and an action

$$I(U_\alpha) = -\frac{K}{32\pi G} \int d^4x (-g)^{1/2} \left[ g^{\alpha\beta} g^{\mu\nu} (U_{\alpha;\mu} - U_{\mu;\alpha}) (U_{\beta;\nu} - U_{\nu;\beta}) - \frac{2\lambda}{K} (g^{\mu\nu} U_\mu U_\nu + 1) \right] \quad (175)$$

for the vector. (Here  $k$ ,  $K$ , and  $\ell$  are constants,  $F$  is a general scalar function, and  $\lambda$  is a Lagrange multiplier which enforces the constraint  $g^{\alpha\beta}U_\alpha U_\beta = -1$ .) With this formulation, MOND can now be regarded as being a fully-fledged, fully consistent relativistic theory which retains all of MOND's non-relativistic features, and it will be of interest to see how it might address the dark energy and cosmological dark matter problems, problems which we noted earlier fall right in the MOND regime. While the above treatment does show how it is possible to embed the MOND formula of Eq. (96) in a fully relativistic setting, a drawback of the approach is that the function  $F$  (to which the MOND function  $\mu(a/a_0)$  of Eq. (96) is related) is not specified by it, as the only constraint on it is that  $F(kG\sigma^2)$  be a general coordinate scalar. It would thus be of interest to find a dynamics which could constrain the function  $F$  in a way that could naturally lead to the behavior which is phenomenologically required of the MOND function in its large and small argument limits.

## 8.6 Modifications of the nature of the spacetime geometry itself

Beyond the above more or less conventional generalizations of Einstein gravity (more Riemann tensor dependent terms, more fields, or more spacetime dimensions),<sup>61</sup> there are also some less orthodox ones which are also worthy of consideration, generalizations which modify the role played by the metric and the nature of the geometry itself. The first of these is to introduce torsion, with the Christoffel symbols then no longer being symmetric in their lower indices. Since such a modification would typically only be expected to be of relevance to microscopic physics, it is not thought to have any effect on the dark matter and dark energy problems. The second modification is to allow the metric to be non-symmetric in its indices, with  $g_{\mu\nu}$  then being a 16-component tensor with ten symmetric and six antisymmetric components. Such a proposal would not affect the line element  $ds^2 = g_{\mu\nu} dx^\mu dx^\nu$  at all since the symmetric  $dx^\mu dx^\nu$  can only couple to the symmetric part of  $g_{\mu\nu}$ . With the antisymmetric part of the metric having precisely the same number of degrees of freedom as the 6-component electromagnetic field strength ( $\vec{E}$  and  $\vec{B}$ ) this generalization also opens up the possibility of unification of electromagnetism with gravity, a program which has been developed in detail by Kursunoglu [57] and generalized by him to encompass unification with the other fundamental forces as well [58]. Models based on a non-symmetric metric have also been explored by Moffat, enabling him [59, 60, 61, 62, 63] to provide a reasonable explanation of galactic rotation curve systematics and the accelerating universe without the need to invoke dark matter or dark energy.

Perhaps the most radical departure from Einstein gravity was actually the first generalization of it ever considered, when only a couple of years after the very introduction of general relativity Weyl [64] actually generalized Riemannian geometry itself. In what was the very first attempt at unification of electromagnetism with gravitation through metrication, Weyl introduced a local transformation which

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<sup>61</sup>Not that any modification of Einstein gravity can really be described as being conventional.

he referred to as a “gauge” transformation under which the metric and the electromagnetic field would jointly transform as

$$g_{\mu\nu}(x) \rightarrow e^{2\alpha(x)} g_{\mu\nu}(x) \quad , \quad (176)$$

$$A_\mu(x) \rightarrow A_\mu(x) - e\partial_\mu\alpha(x) \quad , \quad (177)$$

with gravitation and electromagnetism thus unifying by sharing a common  $\alpha(x)$ . Given such a joint transformation, Weyl then departed from the Riemannian geometry of general relativity and replaced it with a new geometry, “Weyl Geometry”, in which the Riemannian connection was generalized to the  $A_\mu$ -dependent

$$\hat{\Gamma}_{\mu\nu}^\lambda = \frac{1}{2}g^{\lambda\sigma} (\partial_\mu g_{\sigma\nu} + \partial_\nu g_{\sigma\mu} - \partial_\sigma g_{\mu\nu}) + \frac{1}{e}g^{\lambda\sigma} (g_{\sigma\nu}A_\mu + g_{\sigma\mu}A_\nu - g_{\mu\nu}A_\sigma) \quad , \quad (178)$$

so that instead of the covariant derivative of the metric being zero, it would instead be given by

$$g^{\mu\nu}{}_{;\nu} = \partial_\nu g^{\mu\nu} + \hat{\Gamma}_{\nu\sigma}^\mu g^{\sigma\nu} + \hat{\Gamma}_{\nu\sigma}^\nu g^{\sigma\mu} = \frac{2A^\mu}{e} \quad . \quad (179)$$

The utility of this generalized definition of covariant derivative is that Eq. (179) has the remarkable property of being invariant under the joint transformations of Eqs. (176) and (177) (transformations under which  $\hat{\Gamma}_{\mu\nu}^\lambda$  transforms into itself). With the covariant derivative of a tensor transforming as a vector, Weyl was thus able to connect the tensor  $g^{\mu\nu}$  and the vector  $A^\mu$  in an intricate geometrical fashion, though at the price of departing from Riemannian geometry.<sup>62</sup> It was through the development of this theory that the notion of a gauge transformation was introduced into physics, with its usage

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<sup>62</sup>In passing we note that there is no conflict with general coordinate invariance in having a connection which is not simply the standard pure metric based Christoffel symbol  $\Gamma_{\mu\nu}^\lambda(\text{C.S.}) = (1/2)g^{\lambda\sigma} (\partial_\mu g_{\sigma\nu} + \partial_\nu g_{\sigma\mu} - \partial_\sigma g_{\mu\nu})$ . In fact, in general the connection can even be completely independent of the metric altogether and serve as a totally independent dynamical degree of freedom. All that is required of a general connection  $\hat{\Gamma}_{\mu\nu}^\lambda$  is that it transform the same way under a general coordinate transformation as the standard  $\Gamma_{\mu\nu}^\lambda(\text{C.S.})$ , viz. as  $\hat{\Gamma}_{\mu\nu}^\lambda \rightarrow (\partial x'^\lambda / \partial x^\rho)(\partial x^\tau / \partial x'^\mu)(\partial x^\sigma / \partial x'^\nu)\hat{\Gamma}_{\tau\sigma}^\rho + (\partial x'^\lambda / \partial x^\rho)(\partial^2 x^\rho / \partial x'^\mu \partial x'^\nu)$ , since then covariant derivatives constructed via its use will still transform as true general coordinate tensors. (Because the  $(1/e)g^{\lambda\sigma} (g_{\sigma\nu}A_\mu + g_{\sigma\mu}A_\nu - g_{\mu\nu}A_\sigma)$  term in Eq. (178) is a true tensor, then given the way  $\Gamma_{\mu\nu}^\lambda(\text{C.S.})$  transforms, it follows that the quantity  $\hat{\Gamma}_{\mu\nu}^\lambda$  of Eq. (178) does indeed transform as a connection from which true tensor covariant derivatives such as  $g^{\mu\nu}{}_{;\nu}$  of Eq. (179) can be constructed.) However, in geometries with a metric and an independent connection, even though the metric will still give the invariant line element and be responsible for raising and lowering indices, true tensor quantities such as  $g^{\mu\nu}{}_{;\tau} = \partial_\tau g^{\mu\nu} + \hat{\Gamma}_{\tau\sigma}^\mu g^{\sigma\nu} + \hat{\Gamma}_{\tau\sigma}^\nu g^{\sigma\mu}$  will not in general be zero, with it then not being possible to move the metric in and out of covariant derivatives in the usual Riemannian way. Additionally, in terms of a general connection, the Riemann tensor as constructed as  $R_{\mu\nu\kappa}^\lambda = (\partial\hat{\Gamma}_{\mu\nu}^\lambda / \partial x^\kappa) - (\partial\hat{\Gamma}_{\mu\kappa}^\lambda / \partial x^\nu) + \hat{\Gamma}_{\mu\nu}^\eta \hat{\Gamma}_{\kappa\eta}^\lambda - \hat{\Gamma}_{\mu\kappa}^\eta \hat{\Gamma}_{\nu\eta}^\lambda$  will still transform as a true rank four tensor, with  $g^{\mu\kappa} R_{\mu\lambda\kappa}^\lambda$  still being a general covariant scalar. Using this prescription for the Ricci scalar and forming covariant derivatives with the general  $\hat{\Gamma}_{\mu\nu}^\lambda$  connection will thus still enable us to build an action which is a true general coordinate scalar even when the connection is not taken to be the Christoffel symbol or even not specified at all. And while the variation of such a scalar action with respect to the metric will automatically be a true rank two tensor, its independent (Palatini) variation with respect to an independent connection (in general a variation which is independent of that associated with the metric) will still be a true tensor too, since even though the connection is not itself a tensor, its variation is, as the  $(\partial x'^\lambda / \partial x^\rho)(\partial^2 x^\rho / \partial x'^\mu \partial x'^\nu)$  term drops out in the transformation of the difference between any two connections which both transform as  $\hat{\Gamma}_{\mu\nu}^\lambda \rightarrow (\partial x'^\lambda / \partial x^\rho)(\partial x^\tau / \partial x'^\mu)(\partial x^\sigma / \partial x'^\nu)\hat{\Gamma}_{\tau\sigma}^\rho + (\partial x'^\lambda / \partial x^\rho)(\partial^2 x^\rho / \partial x'^\mu \partial x'^\nu)$ . Consequently, dealing with a connection which is independent of the metric is just as covariant as taking the connection to be the standard Christoffel symbol, and doing so can thus generate departures from standard Riemannian-based gravity which are every bit as covariant. (In such cases it should be noted that if the matter field is taken to have an action which is independent of the connection altogether - such as the free Maxwell action - because of the matter field equations of motion, the covariant derivative of the matter field energy-momentum tensor with respect to the Christoffel-symbol based connection will vanish identically. However since this covariant derivative is linear in  $\Gamma_{\mu\nu}^\lambda(\text{C.S.})$ , and since the covariant derivative of the matter field energy-momentum tensor with respect to  $\hat{\Gamma}_{\mu\nu}^\lambda$  is necessarily a true tensor, it follows that the covariant derivative of the matter field energy-momentum tensor with respect to  $\Gamma_{\mu\nu}^\lambda(\text{C.S.})$  is a true tensor too, with its vanishing in one coordinate system entailing its vanishing in all other coordinate systems too.)

While there is actually no difference between the Palatini approach and the standard approach when the action is just

in Eq. (176) entailing a change in the magnitude of  $g^{\mu\nu}$  and thus of its size (this being the meaning of the term “gauge”). With the Einstein-Hilbert action not being invariant under the transformation of Eq. (176), Weyl’s theory thus possessed neither the Einstein equations or Riemannian geometry, and was not so much a generalization of Einstein gravity as a rather substantial departure from it. Moreover, following the development of quantum mechanics with its complex-valued wave functions, Weyl was able to redefine his gauge transformation so that instead of being accompanied by Eq. (176), the gauge transformation of Eq. (177) would instead be accompanied by a change in the complex phase of the wave function ( $\psi \rightarrow e^{i\alpha}\psi$ ). With Weyl’s original invariance of Eq. (176) entailing that all mass parameters would be zero identically (something which at times prior to the development of spontaneous symmetry breakdown was regarded as totally unacceptable), and with the use of Eq. (177) in accompaniment with  $\psi \rightarrow e^{i\alpha}\psi$  then becoming so fruitful, Weyl’s geometric theory has essentially been set aside<sup>63</sup> – so much so in fact that the term “gauge invariance” is now taken to refer exclusively to complex phase transformations, with the transformations of Eq. (176) now being known as conformal or scale transformations instead.

## 8.7 Conformal gravity

In the course of developing his theory Weyl also discovered [67] a tensor with a remarkable geometric property, the so-called conformal or Weyl tensor

$$C_{\lambda\mu\nu\kappa} = R_{\lambda\mu\nu\kappa} - \frac{1}{2}(g_{\lambda\nu}R_{\mu\kappa} - g_{\lambda\kappa}R_{\mu\nu} - g_{\mu\nu}R_{\lambda\kappa} + g_{\mu\kappa}R_{\lambda\nu}) + \frac{1}{6}R^\alpha{}_\alpha(g_{\lambda\nu}g_{\mu\kappa} - g_{\lambda\kappa}g_{\mu\nu}) \quad . \quad (180)$$

As constructed, this particular combination of the Riemann and Ricci tensors and the Ricci scalar has the property that under the local transformation of Eq. (176) it transforms as

$$C^\lambda{}_{\mu\nu\kappa}(x) \rightarrow C^\lambda{}_{\mu\nu\kappa}(x) \quad (181)$$

with all derivatives of  $\alpha(x)$  dropping out identically. As such the Weyl tensor  $C^\lambda{}_{\mu\nu\kappa}$  thus bears the same relation to conformal transformations as the Maxwell tensor  $F_{\mu\kappa}$  does to gauge transformations, with the kinematic relation  $g^{\mu\kappa}F_{\mu\kappa} = 0$  having as a counterpart the kinematic  $g^{\mu\kappa}C^\lambda{}_{\mu\nu\kappa} = 0$ , with the Weyl tensor being the traceless piece of the Riemann tensor.

Given this particular property of the Weyl tensor, a far less radical version of Weyl’s theory is suggested, one in which the invariance of Eq. (176) is retained but the unification with  $A_\mu$  is dropped, so that the geometry is then a standard strictly Riemannian one in which the covariant derivative of the metric is zero as usual. We thus retain the metric as the gravitational field, have it couple covariantly in the usual way, but endow gravity with an additional symmetry beyond coordinate invariance, viz. conformal symmetry. As such, this theory is known as conformal gravity and has been pursued by various authors, with (as we shall see below) the interest of the present author being in applying it astrophysics and cosmology with a view to solving the dark matter and dark energy problems. The great appeal of this conformal symmetry is that its imposition actually leads to a unique choice of

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the Ricci scalar (since the stationary variation with respect to the  $\hat{\Gamma}^\lambda{}_{\mu\nu}$  connection then yields a condition whose solution requires the stationary  $\hat{\Gamma}^\lambda{}_{\mu\nu}$  to actually be  $\Gamma^\lambda{}_{\mu\nu}$ (C.S.) after all), in cases where the action contains additional fields and cross-terms between these fields and the Ricci scalar, the Palatini approach can lead to equations of motion for the metric and the connection which are totally different from the ones in which the connection is taken to be given by the standard  $\Gamma^\lambda{}_{\mu\nu}$ (C.S.) from the outset. We thus note that taking advantage of this fact has enabled the authors of [65] to develop a Palatini-based scalar-tensor generalization of the approach of [47] which also allows for cosmic acceleration to emerge as a purely gravitational effect.

<sup>63</sup>See however [66] where it is suggested that Weyl’s vector  $A_\mu$  field might be become massive by a Higgs mechanism in which it combines with the component of the complex doublet scalar field whose vacuum expectation value spontaneously breaks the  $SU(2) \times U(1)$  electroweak symmetry.

gravitational action, as there is one and only one action which is invariant under the local conformal transformation of Eq. (176), namely the Weyl action

$$\begin{aligned} I_W &= -\alpha_g \int d^4x (-g)^{1/2} C_{\lambda\mu\nu\kappa} C^{\lambda\mu\nu\kappa} \\ &= -\alpha_g \int d^4x (-g)^{1/2} \left[ R_{\lambda\mu\nu\kappa} R^{\lambda\mu\nu\kappa} - 2R_{\mu\kappa} R^{\mu\kappa} + \frac{1}{3} (R^\alpha{}_\alpha)^2 \right] \end{aligned} \quad (182)$$

where  $\alpha_g$  has to be dimensionless, with conformal gravity thus possessing a dimensionless coupling constant.<sup>64</sup> Moreover, not only does the conformal symmetry uniquely select out the Weyl action, it expressly forbids the presence of any fundamental cosmological term, and is thus a symmetry which is able to control the cosmological constant;<sup>65</sup> and as we shall see below, even after the conformal symmetry is spontaneously broken (as is needed to generate particle masses), the contribution of a then induced cosmological constant to cosmology will still be under control.<sup>66</sup>

However, in addition to forbidding a fundamental cosmological constant, the same conformal symmetry also forbids a fundamental Planck mass, and thereby excludes any fundamental Newton constant or any fundamental Einstein-Hilbert action. Nonetheless, the conformal theory cannot be excluded on these grounds, since, as we had noted earlier, as long as a theory contains the Schwarzschild solution as an exterior solution, that will suffice to recover the standard solar system phenomenology. However, such an eventuality does not immediately appear likely for the action  $I_W$  since it involves the Riemann tensor, a tensor which does not vanish in the Ricci flat Schwarzschild solution. However, it turns out that the Lanczos Lagrangian

$$I_L = (-g)^{1/2} \left[ R_{\lambda\mu\nu\kappa} R^{\lambda\mu\nu\kappa} - 4R_{\mu\kappa} R^{\mu\kappa} + (R^\alpha{}_\alpha)^2 \right] \quad (183)$$

happens to be a total divergence [69], so that  $I_W$  can be rewritten as

$$I_W = -2\alpha_g \int d^4x (-g)^{1/2} \left[ R_{\mu\kappa} R^{\mu\kappa} - \frac{1}{3} (R^\alpha{}_\alpha)^2 \right] , \quad (184)$$

to now only involve the Ricci tensor and scalar. Variation of the action of Eq. (184) with respect to the metric then yields

$$\frac{1}{(-g)^{1/2}} \frac{\delta I_W}{\delta g_{\mu\nu}} = -2\alpha_g W^{\mu\nu} = -2\alpha_g \left[ W_{(2)}^{\mu\nu} - \frac{1}{3} W_{(1)}^{\mu\nu} \right] \quad (185)$$

where  $W_{(1)}^{\mu\nu}$  and  $W_{(2)}^{\mu\nu}$  were respectively introduced in Eqs. (107) and (108).<sup>67</sup> Then with the variation of the matter field action  $I_M$  with respect to the metric yielding a matter energy-momentum tensor  $T^{\mu\nu}$ , the variation of the total action  $I_W + I_M$  with respect to the metric yields as equation of motion

$$4\alpha_g W^{\mu\nu} = 4\alpha_g \left[ W_{(2)}^{\mu\nu} - \frac{1}{3} W_{(1)}^{\mu\nu} \right] = T^{\mu\nu} , \quad (186)$$

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<sup>64</sup>The parallel between  $C_{\lambda\mu\nu\kappa}^\lambda$  and  $F_{\mu\kappa}$  also carries over to the form of the action with the  $-\alpha_g \int d^4x (-g)^{1/2} C_{\lambda\mu\nu\kappa}^\lambda C^{\lambda\mu\nu\kappa}$  action being the conformal analog of the  $-(1/4) \int d^4x (-g)^{1/2} F_{\mu\kappa} F^{\mu\kappa}$  action, an action whose coefficient is also dimensionless.

<sup>65</sup>It was the ability of conformal symmetry to control the cosmological constant which first attracted the present author to conformal gravity [68], with its application to the dark matter and dark energy problems only coming later.

<sup>66</sup>Conformal invariance Ward identities are not modified by a change in vacuum, with the conformal energy-momentum tensor having to remain traceless under a change of vacuum, with the induced cosmological constant term then having to be neither bigger nor smaller than the other terms in  $T_{\mu\nu}$ .

<sup>67</sup>Alternatively, one can first make the variation of the action  $I_{W_3} = \int d^4x (-g)^{1/2} R_{\lambda\mu\nu\kappa} R^{\lambda\mu\nu\kappa}$  to obtain [14]  $(-g)^{-1/2} \delta I_{W_3} / \delta g_{\mu\nu} = W_{(3)}^{\mu\nu} = 4R^{\mu\nu;\sigma}{}_{;\sigma} - 2R^{\mu\sigma;\nu}{}_{;\sigma} - 2R^{\nu\sigma;\mu}{}_{;\sigma} - 2R^\mu{}_{\sigma\tau\rho} R^{\nu\sigma\tau\rho} + (1/2)g^{\mu\nu} R_{\sigma\tau\rho\lambda} R^{\sigma\tau\rho\lambda}$ , and with the identical vanishing of variation  $\delta I_L / \delta g_{\mu\nu}$  of the Lanczos action  $I_L = \int d^4x I_L$  yielding [14] the kinematic identity  $-2R^\mu{}_{\sigma\tau\rho} R^{\nu\sigma\tau\rho} + (1/2)g^{\mu\nu} R_{\sigma\tau\rho\lambda} R^{\sigma\tau\rho\lambda} + 4R^{\mu\sigma\nu\tau} R_{\sigma\tau} + 4R^\mu{}_{\sigma} R^{\nu\sigma} - 2g^{\mu\nu} R_{\sigma\tau} R^{\sigma\tau} - 2R^\alpha{}_\alpha R^{\mu\nu} + (1/2)g^{\mu\nu} (R^\alpha{}_\alpha)^2 = 0$ , on using Eq. (187) one can then conclude that  $W_{(3)}^{\mu\nu}$  is not independent of  $W_{(1)}^{\mu\nu}$  and  $W_{(2)}^{\mu\nu}$ .

with use of the kinematic identity

$$R^{\mu\sigma;\nu}{}_{;\sigma} = \frac{1}{2}(R^\alpha{}_\alpha)^{;\mu;\nu} + R^{\mu\sigma\nu\tau}R_{\sigma\tau} - R^\mu{}_\sigma R^{\nu\sigma} \quad (187)$$

allowing one to rewrite Eq. (186) in the compact form<sup>68</sup>

$$4\alpha_g W^{\mu\nu} = 4\alpha_g \left[ 2C^{\mu\lambda\nu\kappa}{}_{;\lambda;\kappa} - C^{\mu\lambda\nu\kappa}R_{\lambda\kappa} \right] = T^{\mu\nu} \quad (188)$$

Finally, as we see from the form of  $W_{(1)}^{\mu\nu}$  and  $W_{(2)}^{\mu\nu}$  given in Eqs. (107) and (108), the Schwarzschild solution is indeed an exact solution to the conformal theory in the  $T^{\mu\nu} = 0$  region, just as required.

Now we had noted earlier that within the framework of gravitational theories which did not reduce to the Einstein equations of motion in any limit whatsoever, as long as those theories admit the Schwarzschild solution they would suffice for solar system phenomenology. In principle such theories could include gravitational actions based on arbitrarily high powers of the Ricci tensor and Ricci scalar, with there thus being an infinite number of such options. With there also being an infinite number of options for theories which do contain the Einstein equations, the beauty of conformal symmetry is that it serves as a principle which chooses one and only one gravitational theory from amongst all possible theories, regardless of whether they contain the Einstein equations or not. Since the conformal symmetry sets to zero any fundamental cosmological constant, it thus addresses the most severe problem in physics head on, and is thus a reasonable starting point for developing a gravitational theory; and so in the following we shall look at its observational implications for astrophysics and cosmology. However, we note now that the ability of the conformal theory to address the cosmological constant problem so directly is that this symmetry simultaneously excludes a fundamental Einstein-Hilbert term as well; with the difficulties faced by the standard theory in addressing the cosmological constant problem perhaps being due to the fact that in any starting point which permits the Einstein-Hilbert action, there is no reason to exclude a cosmological constant term at all, with it being very hard to get rid of something which has no reason not to be there.

In the conformal theory Eq. (186) is to replace the standard Einstein equations of Eq. (1), and with the rank two gravitational tensor  $W^{\mu\nu}$  being kinematically traceless, covariantly conserved, and having to transform as

$$W^{\mu\nu}(x) \rightarrow e^{-6\alpha(x)}W^{\mu\nu}(x) \quad (189)$$

under the conformal transformation of Eq. (176), it follows that the consistency of the conformal theory requires its energy-momentum tensor to be traceless, covariantly conserved and transform as

$$T^{\mu\nu}(x) \rightarrow e^{-6\alpha(x)}T^{\mu\nu}(x) \quad (190)$$

too. The consistency of the theory thus sharply constrains the energy-momentum tensors that are allowable, with there being no freedom to introduce new ad hoc sources (such as dark matter or dark energy) at will, with a typical allowed energy-momentum tensor being the generic one we gave earlier as Eq. (61). Moreover, our use of a conformal invariant action for gravity dovetails with our earlier discussion of the constraints of an  $SU(3) \times SU(2) \times U(1)$  invariant theory of massless fermions and gauge bosons on the structure of the energy-momentum tensor. Specifically, even though the fermion spin connection was only introduced into physics in order to make the fermion kinetic energy term  $-\int d^4x(-g)^{1/2}i\bar{\psi}\gamma^\mu(x)[\partial_\mu + \Gamma_\mu(x)]\psi$  be general coordinate invariant, nonetheless under  $\psi(x) \rightarrow e^{-3\alpha(x)/2}\psi(x)$ ,  $V_a^\mu(x) \rightarrow e^{-\alpha(x)}V_a^\mu(x)$  ( $V_a^\mu(x)$  is the vierbein from which  $\gamma^\mu(x) = V_a^\mu(x)\hat{\gamma}^a$  and  $\Gamma_\mu(x) = [\gamma^\nu(x), \partial_\mu\gamma_\nu(x)]/8 - [\gamma^\nu(x), \gamma_\sigma(x)]\Gamma_{\mu\nu}^\sigma/8$  are constructed) this term happens to be locally

<sup>68</sup>Even though  $C_{\lambda\mu\nu\kappa}(x)$  itself acquires no derivatives of  $\alpha(x)$  under a conformal transformation, because of the way the Christoffel symbols transform its covariant derivatives do, with it being the so-called Bach tensor combination  $2C^{\mu\lambda\nu\kappa}{}_{;\lambda;\kappa} - C^{\mu\lambda\nu\kappa}R_{\lambda\kappa}$  in which all derivatives of  $\alpha(x)$  drop out.

conformal invariant as well. Similarly, the general coordinate invariant Maxwell action is conformal invariant too. With the full symmetry of the lightcone being the conformal group  $O(4, 2)$ , and with the conformal group being homomorphic to the Dirac algebra  $SU(2, 2)$ , we see that requiring massless gauge bosons and massless fermions to couple covariantly to geometry forces them to couple to it conformally as well. The only exception to this requirement is the scalar field kinetic energy, since the minimally coupled scalar field kinetic energy is not invariant under  $S(x) \rightarrow e^{-\alpha(x)}S(x)$ ,  $g_{\mu\nu}(x) \rightarrow e^{2\alpha(x)}g_{\mu\nu}(x)$ . Rather it is only the combination  $-\int d^4x(-g)^{1/2} [S^{;\mu}S_{;\mu}/2 - S^2R^\mu{}_\mu/12]$  which is, an aspect of the conformal theory which will prove central in the following. With conformal gravity we thus resolve the conflict between standard gravity and the tracelessness of  $T^{\mu\nu}$  which is required of a theory built out of massless fermions and gauge bosons, with a traceless  $T^{\mu\nu}$  being coupled in Eq. (186) to a gravitational rank two tensor  $W^{\mu\nu}$  which is itself traceless too. Interestingly, while conformal symmetry actually forces the  $SU(3) \times SU(2) \times U(1)$  theory to be second order in the first place,<sup>69</sup> this same symmetry forces gravity itself to be fourth order. Thus while conformal invariance requires the gravitational side of Eq. (186) be fourth order, the same symmetry obliges its matter side to be second order, so that with conformal invariance the matter  $T^{\mu\nu}$  is of the same order (viz. second) as it is in the standard gravitational theory. Having thus singled out and motivated the use of conformal gravity, we turn now to see how it fares in addressing the dark matter and dark energy problems, an exercise which is of interest not just in and of itself, but also because it will uncover some generic features which could serve as a guide for other theoretical approaches which attempt to address these same problems.

## 9 Alternatives to dark matter

### 9.1 Modifying Newtonian gravity at large distances or small accelerations

As it started to become clear that there were mass discrepancies in galaxies, a few physicists started to entertain the possibility that Newton's law of gravity might need modifying on large distance scales, with some early ideas along these lines being made by Finzi [70], Tohline [71], Kuhn and Kruglyak [72] and Sanders [73], with some related work being made by Fahr [74] and Gessner [75]. Typical of these approaches was the introduction of an extra potential which was to accompany the Newtonian one, and while such approaches could in principle even be covariantized via the use of Eq. (145) and thereby be made compatible with the relativity principle, by and large approaches such as these in the end tended not to fit data well enough to make them compelling.<sup>70</sup> More compelling was Milgrom's MOND approach because, without any need for dark matter, it could encompass a large amount of data extremely well via the use of a very compact formula which involved a very small number of assumptions and a universal acceleration scale criterion which not only explained why departures from the luminous Newtonian expectation are to occur for galaxies, but also why none are to be expected for the solar system. MOND fits to the data sample of Fig. 1 for instance are spot on [24], as are its fits to a large sample of low surface brightness galaxies [77], with reasonable fitting being found for a quite wide range and variety of non-relativistic systems from molecular clouds to clusters of galaxies [78], systems which in first approximation are found to obey the fundamental  $v \sim M^{1/4}$  relation between mass and velocity which the MOND theory possesses.<sup>71</sup> And with MOND having now been successfully covariantized, it should be taken seriously since, as noted, it has captured a basic truth of nature,

<sup>69</sup>An action such as  $\kappa \int d^4x(-g)^{1/2}[F_{\mu\nu}F^{\mu\nu}]^2$  would still be gauge invariant but would not be conformal invariant, and would have to have a dimensionful coupling constant  $\kappa$ . The absence of terms such as these is simply assumed by fiat in the standard  $SU(3) \times SU(2) \times U(1)$  theory, with it being conformal invariance which actually provides a rationale for their absence, to thereby secure the renormalizability of the standard model of strong, electromagnetic and weak interactions.

<sup>70</sup>For a more detailed discussion of candidate alternate theories of this type and of the challenges they face see e.g. [76].

<sup>71</sup>In passing we note that a similar mass-velocity relation was also found to hold [79] in the alternate conformal gravity theory which is described below.

namely that it is a universal acceleration scale which determines when departures from the luminous Newtonian expectation are to occur.

## 9.2 Modifying Einstein gravity – the conformal gravity approach

While one can develop non-relativistic theories by starting with the data themselves,<sup>72</sup> in doing so it is not initially clear what covariantized relativistic theory might ultimately ensue or whether such an approach would ever be able to address problems such as the cosmological constant problem, a problem which is intrinsically relativistic. On the other hand if one starts with a relativistic theory which is chosen because it does address the cosmological constant problem and then works down, one does not know what non-relativistic limit one might then encounter. Encouraged by the fact that conformal gravity could address the cosmological constant problem, Mannheim and Kazanas [80, 81] set out to determine the non-relativistic limit of the conformal theory, and did so at that time with a quite limited objective, namely to see whether a theory which did not contain the Einstein equations could nonetheless still recover the standard Newton/Schwarzschild phenomenology.

To this end they endeavored to find an exact conformal gravity analog of the Schwarzschild exterior and interior solutions to standard gravity, viz. they tried to solve the equation  $W^{\mu\nu} = T^{\mu\nu}/4\alpha_g$  for a static, spherically symmetric source. While this meant having to handle a fourth order differential equation, it turned out that it was possible to greatly simplify the problem through use of the underlying conformal symmetry that the theory possessed. Specifically, they noted that under the general coordinate transformation

$$\rho = p(r) \quad , \quad B(r) = \frac{r^2 b(r)}{p^2(r)} \quad , \quad A(r) = \frac{r^2 a(r) p'^2(r)}{p^2(r)} \quad (191)$$

with initially arbitrary function  $p(r)$ , the general static, spherically symmetric line element

$$ds^2 = -b(\rho)dt^2 + a(\rho)d\rho^2 + \rho^2 d\Omega_2 \quad (192)$$

would be brought to the form

$$ds^2 = \frac{p^2(r)}{r^2} \left[ -B(r)dt^2 + A(r)dr^2 + r^2 d\Omega_2 \right] \quad . \quad (193)$$

On now choosing  $p(r)$  according to

$$-\frac{1}{p(r)} = \int \frac{dr}{r^2 [a(r)b(r)]^{1/2}} \quad , \quad (194)$$

the function  $A(r)$  would then be set equal to  $1/B(r)$ , with the line element then being brought to the convenient form

$$ds^2 = \frac{p^2(r)}{r^2} \left[ -B(r)dt^2 + \frac{dr^2}{B(r)} + r^2 d\Omega_2 \right] \quad . \quad (195)$$

While Eq. (195) is coordinate equivalent to Eq. (192) with the two general functions  $a(\rho)$  and  $b(\rho)$  having been traded for two equally general functions  $p(r)$  and  $B(r)$ , its utility lies in the fact that one of these two functions appears purely as an overall multiplier in Eq. (195). Consequently, through use of Eqs. (189) and (190) the function  $p(r)$  can be removed from the theory altogether (i.e. gauged away) via a conformal transformation, with the full kinematic content of the conformal theory then being contained in the line element

$$ds^2 = -B(r)dt^2 + \frac{dr^2}{B(r)} + r^2 d\Omega_2 \quad (196)$$

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<sup>72</sup>Such approaches while phenomenological should not be thought of as being any more ad hoc than dark matter itself.

in the static, spherically symmetric case. Comparing with the standard theory, we see that instead of having deal with coupled second order equations for two independent functions, we instead have to deal with a fourth order equation for one alone, an equitable enough trade-off.

Evaluating the form that  $W^{\mu\nu}$  takes in the line element of Eq. (196) is straightforward though lengthy, and leads to

$$\begin{aligned} \frac{W^{rr}}{B(r)} &= \frac{B'B'''}{6} - \frac{B''^2}{12} - \frac{1}{3r}(BB''' - B'B'') \\ &\quad - \frac{1}{3r^2}(BB'' + B'^2) + \frac{2BB'}{3r^3} - \frac{B^2}{3r^4} + \frac{1}{3r^4} , \end{aligned} \quad (197)$$

$$\begin{aligned} W^{00} &= -\frac{B''''}{3} + \frac{B''^2}{12B} - \frac{B'''B'}{6B} - \frac{B'''}{r} - \frac{B''B'}{3rB} \\ &\quad + \frac{B''}{3r^2} + \frac{B'^2}{3r^2B} - \frac{2B'}{3r^3} - \frac{1}{3r^4B} + \frac{B}{3r^4} \end{aligned} \quad (198)$$

for its components of interest. Combining Eqs. (197) and (198) then yields the remarkably simple

$$\frac{3}{B} (W^0_0 - W^r_r) = B'''' + \frac{4B'''}{r} = \frac{1}{r}(rB)'''' = \nabla^4 B , \quad (199)$$

so that in terms of the convenient source function  $f(r)$  defined via

$$f(r) = \frac{3}{4\alpha_g B(r)} (T^0_0 - T^r_r) \quad (200)$$

the equation of motion of Eq. (186) can then be written in the extremely compact form

$$\nabla^4 B(r) = f(r) . \quad (201)$$

The remarkably simple equation of motion of Eq. (201) contains the full dynamical content of the conformal theory in the static, spherically symmetric case, an equation of motion which is completely exact and which has not required the use of any perturbative approximations whatsoever. While we have not advocated a preference for one theory over another on the grounds of calculational simplicity, preferring to make assessments on the basis of conceptual simplicity instead, nonetheless from a purely calculational standpoint, Eq. (201) can claim to be just as simple as (and perhaps even simpler than) its standard gravity counterpart given in Eqs. (44) and (45). The solution to Eq. (201) is readily given as the analog of the solution to Eq. (86) given earlier, and can be written as

$$\begin{aligned} B(r > R) &= -\frac{r}{2} \int_0^R dr' r'^2 f(r') - \frac{1}{6r} \int_0^R dr' r'^4 f(r') + w - kr^2 , \\ B(r < R) &= -\frac{r}{2} \int_0^r dr' r'^2 f(r') - \frac{1}{6r} \int_0^r dr' r'^4 f(r') \\ &\quad - \frac{1}{2} \int_r^R dr' r'^3 f(r') - \frac{r^2}{6} \int_r^R dr' r' f(r') + w - kr^2 , \end{aligned} \quad (202)$$

where the  $w - kr^2$  term is the general solution to the homogeneous  $\nabla^4 B(r) = 0$  equation. Finally, on defining

$$\gamma = -\frac{1}{2} \int_0^R dr' r'^2 f(r') , \quad 2\beta = \frac{1}{6} \int_0^R dr' r'^4 f(r') , \quad (203)$$

on dropping the  $kr^2$  term (as it does not couple to the source) and setting  $w = 1$ , we see that up to conformal equivalence the conformal gravity metric exterior to a static, spherically source is thus given without any approximation at all as

$$B(r > R) = -g_{00} = \frac{1}{g_{rr}} = 1 - \frac{2\beta}{r} + \gamma r \quad . \quad (204)$$

In the region where  $2\beta/r \gg \gamma r$  we thus nicely recover the Schwarzschild solution, while seeing that departures from it occur only at large distances and not at small ones, with the standard solar system Schwarzschild phenomenology thus being preserved.<sup>73</sup>

### 9.3 Impact of the global Hubble flow on galactic rotation curves

With both  $\beta$  and  $\gamma$  being related to moments of the energy-momentum tensor of the source, and with neither moment having any reason to vanish, a general conformal gravity source would then furnish both terms, with a source such as a star then producing a non-relativistic potential of the form

$$V^*(r) = -\frac{\beta^* c^2}{r} + \frac{\gamma^* c^2 r}{2} \quad , \quad (205)$$

to nicely yield none other than the potential of Eq. (95) whose merits were discussed earlier. In an application of the potential  $V^*(r)$  to spiral galaxies, we need to integrate the  $V^*(r)$  potential over a disk of  $N^*$  stars distributed as  $\Sigma(R) = \Sigma_0 e^{-R/R_0}$ , with use of Eq. (A15) then enabling us to generalize Eq. (115) to

$$\begin{aligned} \frac{v_{\text{lum}}^2}{R} = g^{\text{lum}} = g_{\beta}^{\text{lum}} + g_{\gamma}^{\text{lum}} &= \frac{N^* \beta^* c^2 R}{2R_0^3} \left[ I_0 \left( \frac{R}{2R_0} \right) K_0 \left( \frac{R}{2R_0} \right) - I_1 \left( \frac{R}{2R_0} \right) K_1 \left( \frac{R}{2R_0} \right) \right] \\ &+ \frac{N^* \gamma^* c^2 R}{2R_0} I_1 \left( \frac{R}{2R_0} \right) K_1 \left( \frac{R}{2R_0} \right) \quad . \end{aligned} \quad (206)$$

With the large  $R$  limit of  $v_{\text{lum}}^2/R$  then being given by

$$\frac{v_{\text{lum}}^2}{R} \rightarrow \frac{N^* \beta^* c^2}{R^2} + \frac{N^* \gamma^* c^2}{2} \quad , \quad (207)$$

we see that we precisely recover the two  $N^*$ -dependent terms present in the formula of Eq. (125) that was extracted purely phenomenologically from a study of the centripetal accelerations of the last data points in each of the galaxies in the eleven galaxy sample of Fig. 1. The potential of the conformal theory thus automatically captures much of the essence of Eq. (125).

However, absent from Eq. (207) is the  $N^*$ -independent  $\gamma_0 c^2/2$  term which we saw was crucial to the validity of Eq. (125), and finding its origin is quite subtle and not at all easy. Specifically, intrinsic to the use of Newtonian gravity and to the intuition one then acquires through its repeated use is a theorem due to Newton, namely that if one has a spherical distribution of matter, the net

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<sup>73</sup>While the vanishing of the Ricci tensor entails the vanishing of its derivatives, the vanishing of the derivatives can be obtained without the vanishing of the Ricci tensor itself, with solutions other than the Schwarzschild solution then being possible. The constructive approach which leads to Eq. (204) shows that not only is the metric of Eq. (204) with its the linear potential an exact exterior solution to the conformal theory, such a metric categorizes all departures from Schwarzschild, as every solution to the conformal theory is conformally equivalent to it. In solving the exterior fourth order Laplace equation  $\nabla^4 B(r) = 0$  one would obtain [82, 80] the solution of Eq. (204) as expressed in terms of what would then be two free integration constants  $\beta$  and  $\gamma$ , with it being the matching [81] to the interior region via the fourth order Poisson equation  $\nabla^4 B(r) = f(r)$  which allows one to express these integration constants in terms of moments of the source and establish their physical meaning.

Newtonian potential produced at a given point depends only on the matter interior to its location, with the potentials due to all the points exterior to it mutually cancelling. In fact so central is this notion to our thinking, that when a mass discrepancy is found we immediately introduce dark matter in precisely the region where the discrepancy is found, a strictly local approach to gravity. However, this theorem of Newton only holds for a  $1/r$  potential, and does not hold in any other case. Consequently, in any modification to Newton whatsoever, one cannot exclude the matter exterior to the point of interest, viz. one cannot exclude the effect of the potentials due to the rest of matter in the universe. Moreover, in modifications to Newton which grow with distance, these effects will be quite pronounced, with the biggest contributors to the gravitational potential at any given point then being not the matter sources which are the nearest but rather those which are most distant, a thus global gravitational effect, one which is quite reminiscent of Mach's principle. The contributions of such distant objects would then not depend on the mass of the local galaxy of interest, and since those distant objects form the Hubble flow, one would even expect their net contribution to come with a cosmologically relevant scale. From the rest of the universe then we can thus naturally expect to obtain some form of  $N^*$ -independent  $\gamma_0 c^2/2$  type term with strength of order the phenomenologically obtained value  $\gamma_0 = 3.06 \times 10^{-30} \text{cm}^{-1}$  cited earlier.<sup>74</sup> However, what does not follow from this discussion is why the net effect of the entire rest of the universe would be to explicitly yield a linear potential term of the form  $V = \gamma_0 c^2 R/2$ , as it is precisely such a contribution and no other which leads to a centripetal acceleration of the form  $v^2/R = \gamma_0 c^2/2$ .<sup>75</sup>

To explicitly obtain this global  $\gamma_0 c^2 R/2$  potential, we need to ask how the comoving Hubble flow would look to an observer who uses a coordinate system in which the center of any chosen galaxy is taken to be at rest. With a Robertson-Walker geometry being homogeneous and isotropic, every point in the geometry can serve as the origin of coordinates, and so we can take the center of any given comoving galaxy of interest to serve as that center, with it being the presence of the galaxy of interest itself which provides this choice for the origin of coordinates. Having fixed the origin of coordinates this way we now need to rewrite the comoving Robertson-Walker metric in static Schwarzschild geometry coordinates as referred to this same origin.<sup>76</sup> To this end we note that the general coordinate transformation

$$r = \frac{\rho}{(1 - \gamma_0 \rho/4)^2} \quad , \quad t = \int \frac{d\tau}{R(\tau)} \quad (208)$$

affects the metric transformation

$$\begin{aligned} & -(1 + \gamma_0 r)c^2 dt^2 + \frac{dr^2}{(1 + \gamma_0 r)} + r^2 d\Omega_2 \\ & \rightarrow \frac{(1 + \gamma_0 \rho/4)^2}{R^2(\tau)(1 - \gamma_0 \rho/4)^2} \left[ -c^2 d\tau^2 + \frac{R^2(\tau)}{(1 - \rho^2 \gamma_0^2/16)^2} (d\rho^2 + \rho^2 d\Omega_2) \right] \quad , \end{aligned} \quad (209)$$

to yield a metric which we recognize as being conformal to a Robertson-Walker metric with spatial 3-curvature  $k = -\gamma_0^2/4$  (the spatial part of the Robertson-Walker metric being written in isotropic coordinates here). Now since a Robertson-Walker metric happens to be conformal to flat, the Weyl tensor

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<sup>74</sup>For a galaxy containing  $N^*$  stars, the distance scale at which the galactic potential which they produce becomes of order one is given by  $r \sim 1/N^* \gamma^*$ , a scale which for a typical  $N^* = 10^{11}$  galaxy is numerically precisely of order  $10^{-30} \text{cm}^{-1}$ , with a cosmological scale of order  $1/\gamma_0$  thus being precisely the scale on which galactic potentials become strong.

<sup>75</sup>An application of the purely local galactic Eq. (206) alone to the eleven galaxy sample of Fig. (1) was actually found [83, 84] to fit the shapes of the rotation curves extremely well, but with the data wanting the strength of the total galactic  $\gamma^* N^*$  to be universal rather than  $\gamma^*$  itself. It took the present author quite some time to grasp that the apparent need for a universal  $\gamma^* N^*$  in the fits was actually due to the fact that in fitting data, Eq. (206) was trying to simulate the contribution of a universal,  $N^*$ -independent,  $\gamma^0$  contribution coming from the rest of the universe.

<sup>76</sup>The Robertson-Walker geometry is spherically symmetric about every point while the Schwarzschild geometry is only spherically symmetric about a single point, with the Robertson-Walker geometry being spherically symmetric about that one particular point too.

associated with it would not only be zero, but would even remain so under conformal transformations such as

$$g_{\mu\nu} \rightarrow \frac{(1 + \gamma_0 \rho/4)^2}{R^2(\tau)(1 - \gamma_0 \rho/4)^2} g_{\mu\nu} \quad , \quad (210)$$

transformations which we are free to make in the conformal theory; with a comoving Robertson-Walker geometry with  $k = -\gamma_0^2/4$  thus being coordinate and conformal equivalent to a static Schwarzschild coordinate system metric with none other than a linear potential. As seen by an observer at rest then it follows [25] that the comoving Hubble flow acts precisely as a universal linear potential. Moreover, not only do we see that this effect is expressly associated with a negative spatial 3-curvature,<sup>77</sup> we shall show below that  $k$  actually is negative (and necessarily so in fact) in the cosmology associated with the same conformal theory. We thus see that despite our familiarity with the local nature of Newtonian gravity, global cosmology can have a local observable effect in galaxies, an effect which is of a quite general nature, and which should thus be present in every theory of gravity, and not just the conformal theory being considered here.

With the linear potential of Eq. (209) being centered at the center of any given spiral galaxy,<sup>78</sup> in the weak gravity limit we can directly add it on to the luminous contribution of Eq. (206), with the total centripetal acceleration seen by a particle in the galaxy being given by

$$\begin{aligned} \frac{v_{\text{tot}}^2}{R} = g^{\text{tot}} = g_{\beta}^{\text{lum}} + g_{\gamma}^{\text{lum}} + \frac{\gamma_0 c^2}{2} &= \frac{N^* \beta^* c^2 R}{2R_0^3} \left[ I_0 \left( \frac{R}{2R_0} \right) K_0 \left( \frac{R}{2R_0} \right) - I_1 \left( \frac{R}{2R_0} \right) K_1 \left( \frac{R}{2R_0} \right) \right] \\ &+ \frac{N^* \gamma^* c^2 R}{2R_0} I_1 \left( \frac{R}{2R_0} \right) K_1 \left( \frac{R}{2R_0} \right) + \frac{\gamma_0 c^2}{2} \quad , \quad (211) \end{aligned}$$

with the  $g_{\beta}^{\text{lum}}$  and  $g_{\gamma}^{\text{lum}}$  terms being due to the luminous material within the galaxy and the  $\gamma_0 c^2/2$  term being due to the luminous material in the rest of the universe. With the large  $R$  limit of  $v_{\text{tot}}^2/R$  being given by

$$\frac{v_{\text{tot}}^2}{R} \rightarrow \frac{N^* \beta^* c^2}{R^2} + \frac{N^* \gamma^* c^2}{2} + \frac{\gamma_0 c^2}{2} \quad , \quad (212)$$

we see that from the above  $g^{\text{tot}}$  we precisely recover the phenomenologically established Eq. (125), with conformal gravity thus providing a theoretical rationale for its validity. With the  $\gamma^*$  and  $\gamma_0$  parameters already being fixed by the application of Eq. (125) to the furthest data point in each galaxy, and with the mass to light ratios of the luminous matter distributions in each galaxy already being fixed by the maximum luminous disk prescription for the inner regions of the rotation curves, there is no more freedom left in the theory. The application of Eq. (211) to the entire 280 or so data points in the eleven galaxy sample of Fig. 1 is thus fully prescribed, with it leading to the fits of [25] which are displayed in Fig. 1. As we see, the conformal theory fully captures the systematics of the data, and just like the MOND fits, shows that it is possible to fit galactic rotation curve data without needing to use any galactic dark matter at all.<sup>79</sup>

## 9.4 Comparison of the conformal gravity and MOND fits

To understand why fits to those of the rotation curves which are flat (viz. the bright spirals) can be obtained in a theory which is based on rising potentials, we recall that the bright spirals are all Freeman

<sup>77</sup>If  $k$  is zero there is no effect, while if  $k$  is positive, the transformation of Eq. (208) would lead to a pure imaginary value for  $\gamma_0$ .

<sup>78</sup>In a coordinate system in which any given comoving galaxy is at rest, no other comoving galaxy could simultaneously be at rest also. Nonetheless, since every comoving galaxy can serve as the origin of the Robertson-Walker coordinate system, for the purposes of determining the motion of particles with respect to the center of any particular galaxy, we can refer Eq. (209) with its universal  $\gamma_0$  to the center of the galaxy of interest each and every time.

<sup>79</sup>As with the analogous MOND case, the numerical values obtained for the conformal gravity parameters used in Eq. (211) ensure that no modifications to the luminous Newtonian expectation are to occur on solar system distance scales.

limit galaxies with a common central surface brightness  $\Sigma_0^F$ , and thus a common value for  $N^*/2\pi R_0^2$ . In a theory with rising potentials, such potentials make a quite small contribution in the inner regions of the rotation curves, with the peak at  $R = 2.2R_0$  being controlled almost completely by the luminous Newtonian contribution, to thus yield

$$\frac{v_{\text{lum}}^2}{c^2} \sim \frac{0.4N^*\beta^*}{R_0} = 0.8\pi\Sigma_0^F\beta^*R_0 \quad (213)$$

at the peak. With reference to say the typical galaxy NGC 3198, we match the inner region peak value at  $R = 2.2R_0$  to the value at the furthest data point at  $R = 10R_0$  where Eq. (125) holds, to obtain

$$\frac{0.4N^*\beta^*}{R_0} = 0.8\pi\Sigma_0^F\beta^*R_0 = \frac{\beta^*N^*}{10R_0} + 5R_0(\gamma^*N^* + \gamma_0) \quad , \quad (214)$$

and thus

$$\gamma^*N^* + \gamma_0 = \frac{0.06N^*\beta^*}{R_0^2} = 0.12\pi\Sigma_0^F\beta^* \quad . \quad (215)$$

With  $\gamma^*N^*$  and  $\gamma_0$  being of the same order for such galaxies, the universality of  $\gamma_0$  and  $\Sigma_0^F$  are thus correlated, to not only automatically enable all Freeman limit galaxies to match their inner and outer regions (something which we recall is completely contrived in dark matter fits), but also to suggest that the very existence of the universal  $\Sigma_0^F$  itself is of cosmological origin, and that it might therefore emerge naturally in a cosmologically based theory of galaxy formation. Additionally, we also note from Fig. 1 that in the intermediate region of the NGC 3198 rotation curve at around  $R = 6R_0$  or so, the luminous Newtonian contribution is at about half of its inner region peak value, while the linear potential contribution is at about half of its outer region value, with the sum of their contributions thus being equal to the value of  $v_{\text{lum}}^2$  obtained in the inner and outer regions themselves. (With one contribution falling and the other rising, there has to be some intermediate point where they cross.) Matching the inner and outer regions to each other via Eq. (214) then automatically takes care of matching the outer and inner regions to the intermediate region as well, to thus make the rotation curve pretty flat over the entire  $2.2R_0 \leq R \leq 10R_0$  region (again we recall how completely contrived this particular matching is in dark matter fits). While such flatness is thus naturally achieved in the conformal theory, it is important to stress that it is only achievable out to some maximum value of  $R$  before the rise due to the linear potential wins out over the falling Newtonian contribution, an aspect of the theory which is immediately evident in the fits to the lower luminosity galaxies shown in Fig. 1. The view of the conformal theory then is that the currently rising low luminosity rotation curves will continue to rise while the currently flat high luminosity rotation curves will eventually start to rise at large enough  $R$ . Such an expectation stands in sharp contrast to both the dark matter and MOND expectations where the rising low luminosity rotation curves are expected to flatten off and the flat high luminosity rotation curves are expected to stay flat. Extending galactic rotation curves to larger distances from the centers of galaxies could therefore be quite instructive.

The essence of the conformal gravity fits presented above lies in the specific role played by cosmology, with the view of conformal gravity being that as a test particle orbits a galaxy, it not only sees the local field produced by the luminous material within that selfsame galaxy, the test particle also sees the gravitational field produced by the rest of the universe. Particles orbiting a galaxy thus probe the global cosmology and thereby enable us to measure the global spatial curvature of the universe, which from the measured value for  $\gamma_0$  is thus given by  $k = -2.3 \times 10^{-60} \text{cm}^{-2}$ . In the conformal theory the rest of the universe thus replaces galactic dark matter, with the view of the conformal theory being that the presence of dark matter in galaxies is nothing more than an artifact which is engendered by trying to describe a global phenomenon in purely local terms.<sup>80</sup>

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<sup>80</sup>This same could of course equally be said of dark matter in even larger systems such as clusters of galaxies where

With Eq. (211) being writable as

$$a = g^{\text{lum}} + \frac{\gamma_0 c^2}{2} \quad , \quad (216)$$

we additionally see that we can write it in the form of Eq. (99), viz.

$$a = \nu \left( \frac{g^{\text{lum}}}{(\gamma_0 c^2/2)} \right) g^{\text{lum}} \quad (217)$$

where

$$\nu(x) = 1 + \frac{1}{x} \quad . \quad (218)$$

We thus recognize Eq. (216) as being in the form of a MOND type equation with  $\gamma_0 c^2/2$  playing the role of a universal acceleration, but with the local  $g_\beta^{\text{lum}}$  being replaced by the entire local  $g^{\text{lum}} = g_\beta^{\text{lum}} + g_\gamma^{\text{lum}}$  which is due to both the Newtonian and the linear potentials of the luminous matter within the galaxy. Conformal gravity thus provides a rationale for why there should be a universal acceleration in the first place, and also for why it should be associated with a cosmic scale.<sup>81</sup> Moreover, while it had been noted earlier that the universal acceleration of MOND had a magnitude of order  $cH_0$ , this was quite puzzling since  $H_0$  is epoch-dependent. However, cosmology can actually supply not just one but two natural scales, viz. both  $H_0$  and  $k$ , and unlike  $H_0$ , the spatial 3-curvature  $k$  is a true general coordinate scalar which is not epoch-dependent at all. It is thus  $k$  which is the natural quantity with which to associate a universal scale, with it being  $k$  which produces a global imprint on galactic rotation curves. The success of the conformal gravity fits thus suggests a role for cosmology in the interpretation of galactic rotation curves as well as a role for the spatial curvature of the universe. We believe these to be general features which should be sought in any theory which attempts to explain galactic rotation curves, with the absence of a fundamental scale such as  $k$  being a possible shortcoming of the standard hierarchical  $\Omega_k = 0$ ,  $\Lambda$ CDM models of dark matter halos.<sup>82</sup>

As a final comment on the conformal gravity fits to galactic rotation curves which we have presented here, we note that even more remarkable than the fits themselves is that the conformal theory was never constructed for this particular purpose. Rather, the theory was selected for an entirely different purpose (controlling the cosmological constant), and when Mannheim and Kazanas set out to determine the non-relativistic limit of the conformal theory, they had no idea what they would find (as indicated earlier, they had only wanted to see whether the theory could recover the standard Newton/Schwarzschild phenomenology), and had no inkling at all as to how the theory might then impact on galactic rotation curves. The fact that the theory then does impact so well on the rotation curves suggests that one should therefore give the theory serious consideration. With the motivation for the conformal theory being the cosmological constant problem, we turn now to a discussion of how the theory treats that particular problem, and in the fitting to the Hubble plot which we shall describe below, we shall again

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the natural interplay in the conformal theory between local and global effects would be even more pronounced. With a cluster of galaxies being a large local inhomogeneity in a homogeneous global cosmological background, a full treatment of clusters in the conformal theory has to await the development of a treatment of the growth of fluctuations in conformal cosmology.

<sup>81</sup>While the MOND and conformal gravity theories agree on the need for a universal scale in galaxies, they depart on how the theories are to behave in the region where the existence of the scale is of consequence. (In passing we note that the numerical relation between the  $a_0$  and  $\gamma_0 c^2$  parameters may be obtained by looking at the behavior of the typical galaxy NGC 3198 in the intermediate region near  $R = 6R_0$  where the Newtonian and linear contributions cross each other in Fig. 1. With  $\gamma^* N^*$  being of order  $\gamma_0$  for this galaxy, at the crossing point we have (c.f. Eq. (212))  $\beta^* N^*/R^2 \sim \gamma_0$  and thus  $v^2/R \sim 2\gamma_0 c^2$ . Comparing this with the MOND value (c.f. Eq. (102)) of  $v^2/R \sim (a_0)^{1/2} (\beta^* N^* c^2/R^2)^{1/2}$  then yields the relation  $a_0 = 4\gamma_0 c^2$ , a relation which numerically is closely obeyed.)

<sup>82</sup>Whatever one's views on alternate theories such as MOND or conformal gravity, if the answer is to be dark matter, then dark matter theory really has to be able to reproduce the regularities in the galactic rotation curve data which these alternate theories so readily capture, and has to do so with the same easy facility.

find the theory working remarkably well on data it had not at all been designed for, data which actually did not even exist at the time [68, 85] conformal cosmology was first developed.

## 10 Alternatives to dark energy

### 10.1 The conformal gravity alternative to dark energy

In order to identify possible alternatives to dark energy, we first present the conformal gravity theory treatment of the issue, and then extract some generic features from it. In applying conformal gravity to cosmology we note that since the Weyl tensor vanishes in a Robertson-Walker geometry, in conformal cosmology the entire left-hand side of Eq. (188) vanishes identically, with the equation of motion for conformal cosmology thus reducing to the extremely simple

$$T^{\mu\nu} = 0 \quad . \quad (219)$$

With this vanishing of the cosmological  $T^{\mu\nu}$ , we see immediately that in the conformal theory the zero of energy is completely determined, to thus give us control of the cosmological constant.<sup>83</sup> To appreciate the implications of Eq. (219), we consider as cosmological energy-momentum tensor the typical matter field conformal energy-momentum tensor given in Eq. (64). With this energy-momentum tensor having to transform according to Eq. (190) under a conformal transformation, its vanishing will persist after any conformal transformation is made. With the scalar field transforming as  $S(x) \rightarrow e^{-\alpha(x)}S(x)$  under a conformal transformation, in field configurations in which  $S(x)$  is non-zero, the full content of the theory can be obtained by working in the particular gauge in which the scalar field takes the constant value  $S_0$ . In such a gauge, use of the matter field equations of motion allows us to then write the energy-momentum tensor in the convenient form

$$T^{\mu\nu} = i\bar{\psi}\gamma^\mu(x)[\partial^\nu + \Gamma^\nu(x)]\psi - \frac{1}{6}S_0^2 \left( R^{\mu\nu} - \frac{1}{2}g^{\mu\nu}R^\alpha_\alpha \right) - g^{\mu\nu}\lambda S_0^4 \quad . \quad (220)$$

An incoherent averaging of  $i\bar{\psi}\gamma^\mu(x)[\partial^\nu + \Gamma^\nu(x)]\psi$  over all the fermionic modes propagating in a Robertson-Walker background will bring the fermionic contribution to  $T^{\mu\nu}$  to the form of a kinematic perfect fluid

$$T_{\text{kin}}^{\mu\nu} = \frac{1}{c} [(\rho_m + p_m)U^\mu U^\nu + p_m g^{\mu\nu}] \quad , \quad (221)$$

with the conformal cosmology equation of motion then taking the form [85, 86, 38]

$$T^{\mu\nu} = T_{\text{kin}}^{\mu\nu} - \frac{1}{6}S_0^2 \left( R^{\mu\nu} - \frac{1}{2}g^{\mu\nu}R^\alpha_\alpha \right) - g^{\mu\nu}\lambda S_0^4 = 0 \quad . \quad (222)$$

Despite its appearance, the  $T^{\mu\nu} = 0$  condition can actually be satisfied non-trivially, since the vanishing of the full  $T^{\mu\nu}$  can be effected by a cancellation of the perfect fluid contribution against the  $-(1/6)S_0^2 (R^{\mu\nu} - g^{\mu\nu}R^\alpha_\alpha/2)$  term associated with the back reaction of the scalar field on the geometry. Thus in the conformal theory there is energy density not just in the matter fields but in the geometry as well (conformal invariance requires the presence of the  $-S^2R^\alpha_\alpha/12$  term in Eq. (61) with its necessarily negative sign), with it being this additional energy density which actually drives the theory. Finally, we note that despite the fact that it is only the full  $T^{\mu\nu}$  which is covariantly conserved, because the Einstein tensor and metric tensor terms in Eq. (222) are independently covariantly conserved, it follows that

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<sup>83</sup>While the trace of the energy-momentum tensor must in general always vanish in the conformal theory to thereby already give us control of its various components, for cosmology the energy-momentum tensor itself must vanish too, to give us yet more control.

$T_{\text{kin}}^{\mu\nu}$  is covariantly conserved also. In the  $S = S_0$  gauge then the perfect fluid does not exchange energy and momentum with the gravitational field, and thus obeys precisely the same covariant conservation condition as is used in the standard theory, with fermionic particles on average moving geodesically in the Robertson-Walker geometry.<sup>84</sup>

With a slight rewriting of Eq. (222) as

$$\frac{1}{6}S_0^2 \left( R^{\mu\nu} - \frac{1}{2}g^{\mu\nu} R^\alpha{}_\alpha \right) = T_{\text{kin}}^{\mu\nu} - g^{\mu\nu} \lambda S_0^4 \quad , \quad (223)$$

we recognize the conformal cosmology Eq. (223) as being of the form of none other than the cosmological evolution equation of the standard theory, save only for the fact that the standard  $G$  has been replaced by an effective, dynamically induced one given by

$$G_{\text{eff}} = -\frac{3c^3}{4\pi S_0^2} \quad . \quad (224)$$

Thus as noted in [85], conformal cosmology is controlled by an effective gravitational coupling constant which is repulsive rather than attractive, and which becomes smaller the larger  $S_0$  might be.<sup>85</sup> If now we define conformal analogs of the standard  $\Omega_M(t)$  and  $\Omega_\Lambda(t)$  via

$$\bar{\Omega}_M(t) = \frac{8\pi G_{\text{eff}} \rho_m(t)}{3c^2 H^2(t)} \quad , \quad \bar{\Omega}_\Lambda(t) = \frac{8\pi G_{\text{eff}} \Lambda}{3c H^2(t)} \quad (225)$$

where  $\Lambda = \lambda S_0^4$ , then in a Robertson-Walker geometry Eq. (223) yields

$$\begin{aligned} \dot{R}^2(t) + kc^2 &= \dot{R}^2(t) \left( \bar{\Omega}_M(t) + \bar{\Omega}_\Lambda(t) \right) \quad , \quad \bar{\Omega}_M(t) + \bar{\Omega}_\Lambda(t) + \bar{\Omega}_k(t) = 1 \quad , \\ q(t) &= \frac{1}{2} \left( 1 + \frac{3p_m}{\rho_m} \right) \bar{\Omega}_M(t) - \bar{\Omega}_\Lambda(t) \end{aligned} \quad (226)$$

as the evolution equation of conformal cosmology,<sup>86</sup> an evolution equation which only departs from the standard Friedmann evolution equation of Eq. (138) through the replacement of  $G$  by  $G_{\text{eff}}$ .<sup>87</sup> With  $G_{\text{eff}}$  being smaller the larger  $S_0$  itself might be, we see that the larger the cosmological constant of the theory (viz. the larger  $\lambda S_0^4$ ), the less it will contribute to cosmic evolution, with  $\bar{\Omega}_\Lambda(t)$  having a self-quenching property not possessed by the standard  $\Omega_\Lambda(t)$ , as the latter comes with a fixed, pre-assigned

<sup>84</sup>The energy-momentum tensor of Eq. (222) provides us with an explicit example of a phenomenon we discussed earlier, namely that covariant conservation of the kinematic fluid energy-momentum tensor is achievable even if the full energy-momentum tensor differs from the purely kinematic one.

<sup>85</sup>In the conformal theory local non-relativistic solar system gravity is controlled by the parameter  $\beta$  in the metric of Eq. (204). With the conformal coupling constant  $\alpha_g$  not participating in Eq. (188) in homogeneous geometries such as the cosmological one in which the Weyl tensor is zero, while participating in the static source  $f(r)$  of Eq. (200) in geometries such as the inhomogeneous Schwarzschild one in which the Weyl tensor is non-zero,  $G_{\text{eff}}$  is completely decoupled from the local  $G = \beta c^2/M$  associated with a source of mass  $M$ . With the sign of the local  $G$  being fixed by the sign of  $\alpha_g$ , a negative effective global cosmological  $G_{\text{eff}}$  is not in conflict with the existence of a positive local  $G$ . The fact that the dynamically induced  $G_{\text{eff}}$  is negative in the conformal theory had been thought of as being a disadvantage since it seemed to imply that the local  $G$  would be given by the same negative  $G_{\text{eff}}$  and then be repulsive too. However, as we see, a repulsive global cosmological  $G_{\text{eff}}$  and an attractive local  $G$  can coexist in one and the same theory, an aspect of the theory which can now actually be regarded as a plus given the recent discovery of cosmic repulsion. (Recall that the best fit line of Eq. (141) to the accelerating universe data would go through the negative  $\Omega_M(t_0) = -0.34$  if  $\Lambda = 0$ .)

<sup>86</sup>In Eq. (226) the quantity  $\bar{\Omega}_k(t)$  is the same as the previously defined  $\Omega_k(t) = -kc^2/R^2$ .

<sup>87</sup>With conformal invariance forcing the matter field action to be second order, in geometries in which the Weyl tensor vanishes the gravitational equations are only second order, and thus coincide in form with those of the standard second order Friedmann equations. Thus even for cosmology a preference for simplicity of the equations would not favor standard gravity over the conformal alternative.

$G$ . Consequently, in the conformal theory a large unquenched  $\Lambda$  can still yield a small  $\bar{\Omega}_\Lambda(t)$ , with it thus being possible to have a  $\Lambda$  as big as particle physics suggests, and yet, as we shall see in detail below, nonetheless not be in conflict with observation. The essence of the conformal gravity treatment of the cosmological constant then will be not to quench  $\Lambda$  at all but to instead quench the amount (viz.  $G_{\text{eff}}$ ) by which it gravitates.

In applying Eq. (226) to cosmology we have to recognize three distinct epochs. Based on our notion of a universe that goes through grandunification and electroweak symmetry breaking phase transitions as it expands and cools, and of each such phase transition having its own scalar field order parameter  $S_0$  and critical temperature  $T_V$ , the universe will start at the highest temperatures above all phase transitions (where every  $S_0$  will be zero), cool through a grand-unified phase transition temperature  $T_V^{\text{GUT}}$  with order parameter  $S_0^{\text{GUT}}$ , and then subsequently go through the electroweak phase transition temperature  $T_V^{\text{EW}}$  with order parameter  $S_0^{\text{EW}}$  at which fermion and intermediate vector boson masses are generated. For temperatures above  $T_V^{\text{EW}}$  all particles will be massless, and  $T_{\text{kin}}^{\mu\nu}$  will act as a radiation fluid with

$$\rho_m = 3p_m = \frac{A}{R^4(t)} \quad . \quad (227)$$

For temperatures below  $T_V^{\text{EW}}$  there will be a net induced cosmological constant which we can represent by an equivalent black-body with necessarily negative  $\Lambda = -\sigma T_V^4/c$  (the free energy having been lowered by the phase transition, not raised) where  $T_V$  represents some blended value of  $T_V^{\text{GUT}}$  and  $T_V^{\text{EW}}$ . Analogously,  $S_0^{\text{GUT}}$  and  $S_0^{\text{EW}}$  can be represented by some blended value  $S_0$ . Since both of the  $T_V^{\text{GUT}}$  and  $T_V^{\text{EW}}$  temperatures are so huge, it will not matter which particular blended  $T_V$  we might use since compared to current temperatures it will be huge also; and with  $\sigma T_V^4$  then being overwhelmingly larger than the currently measured  $\rho_m(t_0)$ , it will not matter if we ignore particle masses altogether.<sup>88</sup> Thus we can define the model as being one whose evolution at temperatures less than  $T_V$  obeys Eq. (226) with  $\rho_m = 3p_m = \sigma T^4$  and  $\Lambda = -\sigma T_V^4/c$  ( $T_V$  being huge), and which at temperatures above  $T_V$  obeys

$$T_{\text{kin}}^{\mu\nu} = 0 \quad (228)$$

with  $S_0$  being zero.<sup>89</sup>

At first glance the condition  $T_{\text{kin}}^{\mu\nu} = 0$  would appear to possess no non-trivial solution. However, that is not the case, since  $T_{\text{kin}}^{\mu\nu}$  is constructed here as an incoherent averaging of all the matter field modes propagating not in a flat space (where the solution to  $T_{\text{kin}}^{\mu\nu} = 0$  would of course be trivial) but in a curved one, with the equation  $T_{\text{kin}}^{\mu\nu} = 0$  being found [86] to actually support a non-trivial solution if the spatial 3-curvature of the Robertson-Walker geometry is expressly negative, since there is then

<sup>88</sup>For  $T_V$  of order say  $10^{15}$ °K and  $\rho_m(t_0)$  of order the detected luminous matter density in the universe at the current time  $t_0$ , the ratio  $c\Lambda/\rho_m(t_0)$  will be of order  $10^{60}$ , with  $\rho_m(t)$  only being competitive with  $c\Lambda$  at early universe temperatures of order  $T_V$  where particle masses are anyway unimportant.

<sup>89</sup>We should perhaps clarify what we mean by  $S_0 = 0$  here. Our view of the field  $S_0$  is that it is a classical order parameter field which represents a spontaneously broken vacuum expectation value of a fermion composite (typically quadrilinear for grandunification and mass generating bilinear for electroweak) which is only operative in the critical regions below  $T_V^{\text{GUT}}$  and  $T_V^{\text{EW}}$ , and thus absent altogether at temperatures above  $T_V^{\text{GUT}}$ . However, there could still be fundamental scalar fields in the theory as well (as there indeed would for instance be if conformal gravity is given a conformal supergravity extension). While such scalar fields would have vanishing vacuum expectation values above the highest critical temperature (the vacuum then being normal rather than spontaneously broken), nonetheless in such a vacuum the expectation value of  $S^2$  would not vanish, since a quantum scalar field  $\hat{S}$  can still connect a normal vacuum  $|\Omega\rangle$  to the one-particle state  $|n\rangle$  created out of it. In such a case the energy-momentum tensor associated with the kinetic energy of the scalar field would be given (c.f. Eq. (64)) as  $T^{\mu\nu} = (2/3)S^{;\mu}S^{;\nu} - (1/6)g^{\mu\nu}S^{;\alpha}S_{;\alpha} - (1/3)SS^{;\mu;\nu} + (1/3)g^{\mu\nu}SS^{;\alpha}_{;\alpha} - (1/6)S^2(R^{\mu\nu} - (1/2)g^{\mu\nu}R^\alpha_\alpha)$  as evaluated in modes which obey  $S^{;\mu}_{;\mu} + (1/6)SR^\mu_\mu = 0$ , with the classical  $S$  now denoting the quantum scalar field matrix element  $\langle\Omega|\hat{S}|n\rangle$ . Then even though the vacuum is such that  $\langle\Omega|\hat{S}|\Omega\rangle$  is zero, an incoherent sum over all  $|n\rangle$  would still generate a non-zero contribution to the scalar field bilinear  $T^{\mu\nu}$  which would then be included as part of  $T_{\text{kin}}^{\mu\nu}$ .

negative energy density present in the gravitational field to effect the needed cancellation. Conformal cosmology thus specifies the sign of  $k$ ,<sup>90</sup> leading to the same negative value that we previously found would impact on galactic rotation curves. With conformal cosmology early universe dynamics thus fixing the sign of  $k$  (something not the case with the standard Friedmann cosmology), and with  $k$  being a general coordinate scalar, the global topology of the universe would then not change as the universe cools down, with  $k$  then still being negative at temperatures below  $T_V$  where Eq. (226) applies.

## 10.2 Conformal gravity and the accelerating universe

Given the above considerations, the signs of various parameters of the theory are then fully specified, with  $k$ ,  $\Lambda$  and  $G_{\text{eff}}$  all being negative and  $\rho_m(t)$  being positive. Given such a pattern of signs, the temperature evolution of the cosmology is then completely determined. Specifically, with the deceleration parameter  $q(t)$  being given by

$$q(t) = \bar{\Omega}_M(t) - \bar{\Omega}_\Lambda(t) \quad (229)$$

when  $p_m = \rho_m/3$ , we see that  $q(t)$  is automatically negative in every epoch, with the universe thus being a permanently accelerating one no matter what the magnitudes of the various parameters.<sup>91</sup> Given the various signs of the parameters, conformal cosmology is found to admit of the exact solution [38]

$$R^2(t) = -\frac{k(\beta - 1)}{2\alpha} - \frac{k\beta\sinh^2(\alpha^{1/2}ct)}{\alpha} \quad (230)$$

where

$$\alpha c^2 = -2\lambda S_0^2 = \frac{8\pi G_{\text{eff}}\Lambda}{3c} \quad , \quad \beta = \left(1 - \frac{16A\lambda}{k^2c}\right)^{1/2} . \quad (231)$$

(When  $A = 0$  this solution reduces to the  $k < 0$ ,  $\alpha > 0$  solution given in Eq. (133).) With Eq. (230) entailing that the initial time  $\dot{R}(t = 0)$  is finite, conformal cosmology is singularity-free ( $G_{\text{eff}}$  being repulsive rather than attractive). The cosmology thus expands from a finite rather than zero minimum size  $R_{\text{min}}^2 = -k(\beta - 1)/2\alpha$ , and thus from a finite maximum temperature  $T_{\text{max}} \sim 1/R_{\text{min}}$ , with the temperature evolution being given by

$$\begin{aligned} \frac{T_{\text{max}}^2}{T^2} &= 1 + \frac{2\beta\sinh^2(\alpha^{1/2}ct)}{(\beta - 1)} \quad , \quad \beta = \frac{(T_{\text{max}}^4 + T_V^4)}{(T_{\text{max}}^4 - T_V^4)} \quad , \\ \tanh^2(\alpha^{1/2}ct) &= \left(1 - \frac{T^2}{T_{\text{max}}^2}\right) \left(1 + \frac{T^2 T_{\text{max}}^2}{T_V^4}\right)^{-1} . \end{aligned} \quad (232)$$

In the solution of Eq. (230) the temperature evolution of the quantities  $\bar{\Omega}_\Lambda(t)$  and  $\bar{\Omega}_M(t)$  can be written in closed form as

$$\begin{aligned} \bar{\Omega}_\Lambda(t) &= \left(1 - \frac{T^2}{T_{\text{max}}^2}\right)^{-1} \left(1 + \frac{T^2 T_{\text{max}}^2}{T_V^4}\right)^{-1} = \left(1 - \frac{T^2}{T_{\text{max}}^2}\right)^{-2} \tanh^2(\alpha^{1/2}ct) \quad , \\ \bar{\Omega}_M(t) &= -\frac{T^4}{T_V^4} \bar{\Omega}_\Lambda(t) \quad , \end{aligned} \quad (233)$$

with Eq. (233) holding at any  $T(t)$  without any approximation at all. Quite remarkably, from Eq. (233) we see that no matter what the actual value of  $T_{\text{max}}$ , at temperatures which are well below it, the late time  $\bar{\Omega}_\Lambda(t \gg 0)$  will be given by

$$\bar{\Omega}_\Lambda(t \gg 0) = \tanh^2(\alpha^{1/2}ct) \quad , \quad (234)$$

<sup>90</sup>That  $k$  would be negative in the conformal theory was first noted in [85].

<sup>91</sup>With  $G_{\text{eff}}$  and  $\Lambda$  being negative and  $\rho_m(t)$  being positive, the quantity  $\bar{\Omega}_\Lambda(t)$  is then positive while the quantity  $\bar{\Omega}_m(t)$  is negative.

to thus have to exclusively lie between zero and one, ultimately asymptoting to a bound of one from below. Thus no matter how big  $T_V$  might be, i.e. no matter how big  $\Lambda$  might be, in the conformal theory the current era  $\bar{\Omega}_\Lambda(t_0)$  is automatically bounded from above. With  $\bar{\Omega}_M(t)$  being completely negligible in the current era ( $\bar{\Omega}_M(t_0)/\bar{\Omega}_\Lambda(t_0) = -T_0^4/T_V^4 = O(10^{-60})$ ), it follows from Eq. (226) that the current era  $\bar{\Omega}_k(t_0)$  is given by

$$\bar{\Omega}_k(t_0) = \text{sech}^2(\alpha^{1/2}ct_0) \quad (235)$$

with curvature thus contributing to current era cosmic expansion,<sup>92</sup> with the late time deceleration parameter  $q(t \gg 0)$  being given by the accelerating

$$q(t \gg 0) = -\tanh^2(\alpha^{1/2}ct) \quad , \quad (236)$$

a deceleration parameter which is therefore automatically bounded between minus one and zero.

Not only does conformal cosmology automatically quench the current era  $\bar{\Omega}_\Lambda(t_0)$  and thereby show that it actually is possible to live with a huge cosmological constant, since the cosmology has no initial singularity, it also solves the Friedmann early universe fine-tuning flatness problem, with the absence of any initial singularity not obliging the initial  $\bar{\Omega}_M(t=0) + \bar{\Omega}_\Lambda(t=0)$  to have to be one [85]. Further, the conformal cosmology also solves the cosmic coincidence problem, since  $\bar{\Omega}_M(t_0)$  is nowhere near close in value to  $\bar{\Omega}_\Lambda(t_0)$  at all; and with  $\bar{\Omega}_M(t_0)$  not needing to be of order one (rather it is of order  $10^{-60}$ ) the cosmology even has no need for any cosmological dark matter. With the conformal cosmology also having been shown to have no horizon problem or universe age problem [86, 38], it thus directly addresses many of the major challenges that contemporary cosmology has had to face, and does so without needing to appeal to an early universe inflationary era.<sup>93</sup>

With  $R(t)$  being completely determined in Eq. (230), it is then straightforward to calculate the dependence of the luminosity function on redshift, with the late time luminosity function being found [38, 39] to be given by none other than the pure  $\Lambda$  universe formula given earlier as Eq. (140), viz.

$$d_L = -\frac{c}{H(t_0)} \frac{(1+z)^2}{q_0} \left( 1 - \left[ 1 + q_0 - \frac{q_0}{(1+z)^2} \right]^{1/2} \right) \quad , \quad (237)$$

as expressed in terms of the current era  $H(t_0)$  and  $q_0$ . With  $q_0$  being treated as a free parameter which has to lie between minus one and zero in the conformal case, as described earlier the accelerating universe data are then fitted extremely well with  $q_0 = -0.37$  in the fits exhibited in Figs. 2 and 3, with no fine-tuning of parameters being required. The success of these fits shows that having a particle physics sized cosmological constant is not incompatible with the supernovae data at all, it is only having a particle physics sized cosmological constant coupled to standard gravity which is incompatible. Since the conformal theory was not constructed after the fact in order to explain the accelerating universe data, and since it then does fit the accelerating universe data as naturally and as readily as it does, it should be given serious consideration. And as we indicated earlier, a  $q_0 = -0.37$ ,  $\bar{\Omega}_\Lambda(t_0) = 0.37$ ,  $\bar{\Omega}_M(t_0) = O(10^{-60})$ ,  $\bar{\Omega}_k(t_0) = 0.63$  theory is a fully falsifiable one whose continuing acceleration predictions for higher redshift exhibited in Fig. 3 can be directly tested.

Beyond our interest here in conformal gravity as a classical theory which is to be used to fit astrophysical and cosmological data without dark matter or dark energy, the theory is also of interest as a quantum theory. Since it is based on an action  $I_W$  with a dimensionless coupling constant, as a

<sup>92</sup>With  $k$  being negative and  $\bar{\Omega}_M(t \gg 0)$  being negligible, the effective late time sum rule  $\bar{\Omega}_\Lambda(t \gg 0) + \bar{\Omega}_k(t \gg 0) = 1$  requires that  $\bar{\Omega}_\Lambda(t \gg 0)$  has to lie below one, with it being the negative 3-curvature of the universe which was fixed in the very early universe thus forcing the current era  $\bar{\Omega}_\Lambda(t_0)$  to be bounded from above.

<sup>93</sup>Since a de Sitter geometry is conformal to flat, conformal gravity also admits of a de Sitter based cosmology [68], with it not being excluded that there could have been an early universe inflationary de Sitter phase in the conformal theory even if the flatness and horizon problems that it addresses can be solved without it.

quantum gravity theory the conformal theory is power counting renormalizable, to thus have far better behavior in the ultraviolet than the non-renormalizable standard theory, a difficulty which one tries to get round in the standard theory by generalizing it to a 10-dimensional string theory. Because it is based on fourth order equations of motion, the unitarity of quantum conformal gravity has been called into question since it is suspected that the theory might possess on-shell negative norm ghost states. However, while this is indeed the case for the theories which involve both second and fourth order components, it has recently been shown [87] not to be the case for pure fourth order theories themselves, a point we elaborate on in Appendix D. Conformal gravity may thus well be a completely consistent quantum gravitational theory, one which is formulated purely in four spacetime dimensions alone.

### 10.3 Quenching the contribution of the cosmological constant to cosmic evolution

The essence of the conformal gravity approach to the cosmological constant problem is to have a symmetry which forces  $\Lambda$  to be zero in the exact symmetry limit (viz. above all cosmological phase transitions), and to then, as per particle physics, have a huge one induced as the temperature drops below the various critical points, but to have the huge induced one gravitate far less than it would do in the standard theory. Thus instead of trying to quench  $\Lambda$ , one can instead try to quench  $G$ . As such, a quenching of the cosmological  $G$  can be thought of as being a generic way to approach the cosmological constant problem, one which could apply in theories other than just conformal gravity, theories which need not depart so far from the Einstein-Hilbert action as the conformal theory does. Thus in the standard theory it could be the case that rather than be fixed,  $G$  could be a running coupling constant whose magnitude would depend on epoch or on distance, with a difference between the cosmological early universe  $G$  and the cosmological late universe  $G$  possibly allowing one to solve the early universe fine-tuning problem, and a difference between the large distance cosmological  $G$  and the short distance local  $G$  possibly allowing one to solve the cosmological constant problem.<sup>94</sup> Also, it might be possible to generate an epoch-dependent  $G$  through an embedding in a higher dimensional space, since as noted in Eq. (161), the 4-dimensional  $G$  is then not fundamental but is instead induced by the embedding. In having an epoch-dependent  $G$  which could quench the contribution of the cosmological constant to cosmology, such a quenching would ordinarily be expected to quench the contribution of ordinary matter to cosmology as well. However, in order to solve the cosmological constant problem per se, it is actually only necessary to quench the amount by which  $\Lambda$  itself gravitates, with there being no need to modify the amount by which  $\rho_m(t)$  gravitates; with models which can do this (by taking advantage of the fact that unlike  $\rho_m(t)$  the constant  $\Lambda$  carries zero 4-momentum) currently being under consideration [89, 90, 91]. Given the fact that attempts to quench the cosmological constant itself have so far foundered, attempting to instead quench the amount by which it gravitates appears to be an attractive alternative.

## 11 Future prospects and challenges

In this review we have examined standard gravity and some of its alternatives, and have identified the extension of the standard Newton-Einstein gravitational theory beyond its solar system origins to be the root cause of the dark matter and dark energy problems. However, attempting to critique the standard theory is not an easy enterprise, in part because the definition of what constitutes the standard theory is something of a moving target. While it is agreed upon that the standard theory is to be based on the Einstein equations, such a definition leaves a great deal undetermined, as it leaves  $T^{\mu\nu}$  completely unspecified, and does not make clear whether or not the gravitational side of the Einstein equations

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<sup>94</sup>For a possible role for a running  $G$  in solving the dark matter problem see [88].

is to also include a fundamental cosmological constant term (to cancel the one induced in by phase transitions in  $T^{\mu\nu}$ ).<sup>95</sup> During the period between the development of the inflationary universe model in 1981 and the discovery of the accelerating universe in 1998 the standard model was taken to be one in which both a fundamental and an induced cosmological constant were zero and the matter density was at the critical density (viz.  $\Omega_M(t) = 1$ ,  $\Omega_\Lambda(t) = 0$ ). Then, following the discovery of the accelerating universe this definition was revised to mean a universe with  $\Omega_M(t_0) = 0.3$ , and a net  $\Omega_\Lambda(t_0) = 0.7$ . Such a universe has a ready test since it predicts deceleration at redshifts above one (as per Fig. 3). However, if acceleration were to be established from analysis of the redshift greater than one Hubble plot, the response would not necessarily be to declare the standard theory wrong but to instead introduce, say, a quintessence fluid (or possibly even two or more such fluids) with a parameter  $w = p/\rho$  which would then depend on redshift in whatever way was needed. The weakness of the standard theory is not so much that one necessarily would do this, but rather that one could do it, i.e. that even now the theory is still not uniquely specified in a way that could enable it to make fully falsifiable predictions from which the theory could not subsequently back away.<sup>96</sup>

Nonetheless, even if the road which led to the present  $\Omega_M(t_0) = 0.3$ ,  $\Omega_\Lambda(t_0) = 0.7$  standard paradigm was circuitous, it is still the case that these particular numbers work extremely well when the standard theory is applied to the anisotropy structure of the cosmic microwave background. In fact so well do they work that it is doubted that there could possibly be any second theory which could do as well. However, while it is extremely unlikely for lightning to strike in the same place twice, as far as the standard theory is concerned one has to ask if lightning has even struck once – namely do  $\Omega_M(t_0) = 0.3$  and  $\Omega_\Lambda(t_0) = 0.7$  really describe the energy content of the universe, or is the true  $\rho_M(t_0)$  in the real world the one given by luminous matter alone and the true  $\Lambda$  the one given by a particle physics scale, since if they are, the standard model fits would then be disastrous.

However, if in the real world  $\rho_M(t_0)$  and  $\Lambda$  are such that  $\Omega_M(t_0)$  is equal to 0.01 and  $\Omega_\Lambda(t_0)$  is of order  $10^{60}$ , one has to ask why does  $\Omega_M(t_0) = 0.3$ ,  $\Omega_\Lambda(t_0) = 0.7$  then work. As far as the supernovae Hubble plot data are concerned an answer can be provided to this question. Specifically, as we noted earlier, the best ( $\Omega_M(t_0)$ ,  $\Omega_\Lambda(t_0)$ ) plane fits to the supernovae data lie along the  $\Omega_\Lambda(t_0) = 1.1\Omega_M(t_0) + 0.37$  curve given as Eq. (141). The best fit minimum is thus quite shallow, and if the conformal gravity theory is correct and the real world lies close to the  $\bar{\Omega}_M(t_0) = 0$ ,  $\bar{\Omega}_\Lambda(t_0) = 0.37$  solution to Eq. (226) (a solution which as far as purely phenomenological fitting is concerned is equivalent to using the  $\Omega_M(t_0) = 0$ ,  $\Omega_\Lambda(t_0) = 0.37$  solution to Eq. (138)), the  $\Omega_M(t_0) = 0.3$ ,  $\Omega_\Lambda(t_0) = 0.7$  standard model solution would then just happen to be close to the real solution. Of course, the argument cuts both ways, and it could be that it is the conformal theory fit which is the one which is accidentally correct. To resolve this issue will require fitting the cosmic microwave background with conformal gravity to see if it works too,<sup>97</sup> and if it does, to then see whether the success of either of the two theories (or of any other candidate theory) could reveal why the successes of the other theories were then only accidental.

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<sup>95</sup>The freedom standard gravity has in choosing  $T^{\mu\nu}$  is due to its not using the traceless  $SU(3) \times SU(2) \times U(1)$  based  $T^{\mu\nu}$  provided by fundamental physics. And we note that if we were to couple standard cosmology to a traceless  $T^{\mu\nu}$ , for  $k = 0$  we would obtain  $R(t) \sim t^{1/2}$  and  $d_L = cz/H(t_0)$ , with a best fit to the 54 data points of Fig. 2 yielding  $\chi^2 = 138$ .

<sup>96</sup>The contrast with conformal gravity theory for instance in this regard is that the conformal theory is a uniquely defined theory (with gravitational action  $I_W$  and  $SU(3) \times SU(2) \times U(1)$  invariant matter action) which unambiguously predicts continuing acceleration above a redshift of one. Showing that the universe is in fact decelerating above a redshift of one would thus rule out the conformal theory, while showing that it was accelerating above a redshift of one would not necessarily rule out the standard theory. Alternate theories such as conformal gravity and MOND which endeavor to describe the universe without dark matter or dark energy do not enjoy the luxury of still having functions to adjust, and any major failure would rule them out.

<sup>97</sup>To apply the conformal theory to the cosmic microwave background requires the development of a theory for the growth of inhomogeneities in the model, to see what size standard yardstick the fluctuations would impose on the recombination sky, and what size they would then appear to us. Since the size of the fluctuation yardstick of the conformal theory would be different from the one used in the standard theory, the conformal gravity value  $\bar{\Omega}_k(t_0) = 0.63$  need not be in conflict with the standard model's  $\Omega_k(t_0) = 0$ .

The challenge to alternate theories to fit the cosmic microwave background data is a daunting one, not just computationally, but also because alternate programs are extremely manpower limited with few people working on them.<sup>98</sup> However, this has to be weighed against the fact that the cosmological constant problem is an extremely challenging one for the standard theory, and here there are many workers and yet little progress.<sup>99</sup> However, while this is all for the future, for the present three instructive guidelines have been identified – namely that there is good evidence that it is a universal acceleration scale which determines when a luminous Newtonian expectation is to fail to fit data, that a good case can be made that there is a global cosmological effect on local galactic motions which galactic dark matter might be simulating, and that in attempting to solve the cosmological constant problem one does not actually need to quench the cosmological constant itself, but only the amount by which it gravitates.

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<sup>98</sup>For the conformal gravity theory for instance, while it has been shown that the metric produced by the sun yields the standard solar system Schwarzschild phenomenology, PPN-type corrections to this phenomenology have yet to worked out. While there is (necessarily) gravitational radiation reaction in the conformal theory, its effect on the decay of the orbit of a binary pulsar has yet to be worked out. Application of Eq. (211) to galactic rotation curves has so far only been made for the typical eleven galaxy sample of Fig. (1). Studies of clusters of galaxies, gravitational lensing by them and anisotropies in the cosmic microwave background have yet to be made and all of them await the development of a treatment of the growth of inhomogeneities in the conformal theory. Such a treatment is also needed for the nucleosynthesis of primordial deuterium as not enough of it can survive [92, 93, 94, 95] in a strictly homogeneous cosmology which expands as slowly as that of Eq. (230) – though unlike the standard cosmology, this very same slow expansion does allow [94] the conformal cosmology to get past the barrier associated with the absence of stable nuclei with  $A = 8$  and primordially synthesize <sup>9</sup>Be and beyond. For cosmology, apart from the fits of Figs. (2) and (3) to the accelerating universe data, at the present time all the other achievements of the conformal theory are theoretical, with the cosmology having been shown to possess no flatness problem, no horizon problem, no universe age problem, no cosmic coincidence fine-tuning problem, no dark matter problem, no dark energy problem, and no cosmological constant problem.

<sup>99</sup>One would especially like to see string theory come up with an explanation for why  $\Omega_\Lambda(t_0)$  is equal to 0.7, and by that we mean not just to explain why it is not of order  $10^{60}$  to  $10^{120}$ , but why it is not even equal to 1.7 or -1.7.

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## A The potential of a thin disk

In order to determine the weak gravity potential of an extended object such as a disk of stars each with gravitational potential  $V(r) = -\beta c^2/r + \gamma c^2 r/2$ , we follow an approach originally developed by Casertano [96] for both thin and thick Newtonian disks and then generalize it to disks with linear potentials. For the Newtonian potential of an initially non-thin axially symmetric distribution of matter sources with matter volume density function  $\rho(R', z')$  we need to evaluate the quantity

$$V_\beta(R, z) = -\beta c^2 \int_0^\infty dR' \int_0^{2\pi} d\phi' \int_{-\infty}^\infty dz' \frac{R' \rho(R', z')}{(R^2 + R'^2 - 2RR' \cos\phi' + (z - z')^2)^{1/2}} \quad (\text{A1})$$

where  $R'$ ,  $\phi'$ ,  $z'$  are cylindrical coordinates of the source and  $R$  and  $z$  are the only observation point coordinates of relevance. To evaluate Eq. (A1) it is convenient to make use of the cylindrical coordinate Green's function Bessel function expansion

$$\frac{1}{|\vec{r} - \vec{r}'|} = \sum_{m=-\infty}^\infty \int_0^\infty dk J_m(kr) J_m(kr') e^{im(\phi - \phi') - k|z - z'|} \quad , \quad (\text{A2})$$

with its insertion into Eq. (A1) yielding

$$V_\beta(R, z) = -2\pi\beta c^2 \int_0^\infty dk \int_0^\infty dR' \int_{-\infty}^\infty dz' R' \rho(R', z') J_0(kR) J_0(kR') e^{-k|z - z'|} \quad . \quad (\text{A3})$$

For the case of infinitesimally thin disks with  $\rho(R', z') = \Sigma(R')\delta(z')$  Eq. (A3) then simplifies to

$$V_\beta(R) = -2\pi\beta c^2 \int_0^\infty dk \int_0^\infty dR' R' \Sigma(R') J_0(kR) J_0(kR') \quad (\text{A4})$$

for observation points in the  $z = 0$  plane of the disk. Hence, for a disk with an exponential matter distribution  $\Sigma(R') = \Sigma_0 e^{-\alpha R'}$  and total number of stars  $N = 2\pi\Sigma_0/\alpha^2$  ( $R_0 = 1/\alpha$  is the scale length of the disk), use of the standard Bessel function integral formulae

$$\int_0^\infty dR' R' J_0(kR') e^{-\alpha R'} = \frac{\alpha}{(\alpha^2 + k^2)^{3/2}} \quad , \quad (\text{A5})$$

$$\int_0^\infty dk \frac{J_0(kR)}{(\alpha^2 + k^2)^{3/2}} = \frac{R}{2\alpha} \left[ I_0\left(\frac{R\alpha}{2}\right) K_1\left(\frac{R\alpha}{2}\right) - I_1\left(\frac{R\alpha}{2}\right) K_0\left(\frac{R\alpha}{2}\right) \right] \quad , \quad (\text{A6})$$

then leads directly to

$$\begin{aligned} V_\beta(R) &= -2\pi\beta c^2 \Sigma_0 \int_0^\infty dk \frac{\alpha J_0(kR)}{(\alpha^2 + k^2)^{3/2}} \\ &= -\pi\beta c^2 \Sigma_0 R \left[ I_0\left(\frac{R\alpha}{2}\right) K_1\left(\frac{R\alpha}{2}\right) - I_1\left(\frac{R\alpha}{2}\right) K_0\left(\frac{R\alpha}{2}\right) \right] . \end{aligned} \quad (\text{A7})$$

Finally, on differentiating Eq. (A7) and using the modified Bessel function relations

$$I'_0(z) = I_1(z) \quad , \quad I'_1(z) = I_0(z) - \frac{I_1(z)}{z} \quad , \quad K'_0(z) = -K_1(z) \quad , \quad K'_1(z) = -K_0(z) - \frac{K_1(z)}{z} \quad , \quad (\text{A8})$$

we obtain the relation

$$RV'(R) = \frac{N\beta c^2 R^2 \alpha^3}{2} \left[ I_0\left(\frac{R\alpha}{2}\right) K_0\left(\frac{R\alpha}{2}\right) - I_1\left(\frac{R\alpha}{2}\right) K_1\left(\frac{R\alpha}{2}\right) \right] \quad (\text{A9})$$

presented in the text.

The utility of the above formalism is that it immediately generalizes to the linear potential case (and by extension to the cubic potential and so on), since on setting  $|\vec{r} - \vec{r}'| = (\vec{r} - \vec{r}')^2 / |\vec{r} - \vec{r}'|$ , we immediately obtain the potential

$$\begin{aligned} V_\gamma(R, z) &= \frac{\gamma c^2}{2} \int_0^\infty dR' \int_0^{2\pi} d\phi' \int_{-\infty}^\infty dz' R' \rho(R', z') [R^2 + R'^2 - 2RR' \cos\phi' + (z - z')^2]^{1/2} \\ &= \pi\gamma c^2 \int_0^\infty dk \int_0^\infty dR' \int_{-\infty}^\infty dz' R' \rho(R', z') \left[ (R^2 + R'^2 + (z - z')^2) J_0(kR) J_0(kR') \right. \\ &\quad \left. - 2RR' J_1(kR) J_1(kR') \right] e^{-k|z-z'|} . \end{aligned} \quad (\text{A10})$$

For an infinitesimally thin disk Eq. (A10) then reduces at  $z = 0$  to

$$V_\gamma(R) = \pi\gamma c^2 \int_0^\infty dk \int_0^\infty dR' R' \Sigma(R') \left[ (R^2 + R'^2) J_0(kR) J_0(kR') - 2RR' J_1(kR) J_1(kR') \right] . \quad (\text{A11})$$

The use of Eqs. (A5) and (A6) and the additional integral formula

$$\int_0^\infty dR' R'^2 J_1(kR') e^{-\alpha R'} = \frac{3\alpha k}{(\alpha^2 + k^2)^{5/2}} \quad (\text{A12})$$

then yields

$$\begin{aligned} V_\gamma(R) &= \pi\gamma c^2 \Sigma_0 \int_0^\infty dk \left( \frac{\alpha R^2 J_0(kR)}{(\alpha^2 + k^2)^{3/2}} - \frac{9\alpha J_0(kR)}{(\alpha^2 + k^2)^{5/2}} + \frac{15\alpha^3 J_0(kR)}{(\alpha^2 + k^2)^{7/2}} - \frac{6\alpha k R J_1(kR)}{(\alpha^2 + k^2)^{5/2}} \right) \\ &= \pi\gamma c^2 \Sigma_0 \int_0^\infty dk J_0(kR) \left( \frac{\alpha R^2}{(\alpha^2 + k^2)^{3/2}} + \frac{15\alpha}{(\alpha^2 + k^2)^{5/2}} - \frac{15\alpha^3}{(\alpha^2 + k^2)^{7/2}} \right) . \end{aligned} \quad (\text{A13})$$

Through use of the modified Bessel function relations given above taken in conjunction with Eq. (A6) and its derivatives, Eq. (A13) readily evaluates to

$$\begin{aligned} V_\gamma(R) &= \frac{\pi\gamma c^2 \Sigma_0 R}{\alpha^2} \left[ I_0\left(\frac{R\alpha}{2}\right) K_1\left(\frac{R\alpha}{2}\right) - I_1\left(\frac{R\alpha}{2}\right) K_0\left(\frac{R\alpha}{2}\right) \right] \\ &\quad + \frac{\pi\gamma c^2 \Sigma_0 R^2}{2\alpha} \left[ I_0\left(\frac{R\alpha}{2}\right) K_0\left(\frac{R\alpha}{2}\right) + I_1\left(\frac{R\alpha}{2}\right) K_1\left(\frac{R\alpha}{2}\right) \right] . \end{aligned} \quad (\text{A14})$$

Finally, differentiation with respect to  $R$  and repeated use of the recurrence relations of Eq. (A8) then yields the expression

$$RV'(R) = \frac{N\gamma c^2 R^2 \alpha}{2} I_1\left(\frac{R\alpha}{2}\right) K_1\left(\frac{R\alpha}{2}\right) \quad (\text{A15})$$

presented in the text.

## B The potential of a separable thick disk

For non-thin disks simplification of Eqs. (A1) and (A10) can be achieved for disks with a separable matter distribution  $\rho(R', z') = \Sigma(R')f(z')$  where the symmetric thickness function  $f(z') = f(-z')$  is normalized according to

$$\int_{-\infty}^{\infty} dz' f(z') = 2 \int_0^{\infty} dz' f(z') = 1 \quad . \quad (\text{A16})$$

Recalling that

$$e^{-k|z-z'|} = \theta(z-z')e^{-k(z-z')} + \theta(z'-z)e^{+k(z-z')} \quad , \quad (\text{A17})$$

we find that Eqs. (A1) and (A10) then respectively yield for points with  $z = 0$

$$V_{\beta}(R) = -4\pi\beta c^2 \int_0^{\infty} dk \int_0^{\infty} dR' \int_0^{\infty} dz' R' \Sigma(R') f(z') J_0(kR) J_0(kR') e^{-kz'} \quad (\text{A18})$$

and

$$\begin{aligned} V_{\gamma}(R) &= 2\pi\gamma c^2 \int_0^{\infty} dk \int_0^{\infty} dR' \int_0^{\infty} dz' R' \Sigma(R') f(z') e^{-kz'} \\ &\times \left[ (R^2 + R'^2 + z'^2) J_0(kR) J_0(kR') - 2RR' J_1(kR) J_1(kR') \right] \end{aligned} \quad (\text{A19})$$

in the separable case. Further simplification is possible if  $\Sigma(R')$  is again the exponential  $\Sigma_0 e^{-\alpha R'}$ , with use of the recurrence relation  $J_1'(z) = J_0(z) - J_1(z)/z$  then yielding

$$RV'_{\beta}(R) = 2N\beta c^2 \alpha^3 R \int_0^{\infty} dk \int_0^{\infty} dz' \frac{f(z') e^{-kz'} k J_1(kR)}{(\alpha^2 + k^2)^{3/2}} \quad (\text{A20})$$

and

$$\begin{aligned} RV'_{\gamma}(R) &= N\gamma c^2 \alpha^3 R \int_0^{\infty} dk \int_0^{\infty} dz' f(z') e^{-kz'} \\ &\times \left( -\frac{4RJ_0(kR)}{(\alpha^2 + k^2)^{3/2}} + \frac{6\alpha^2 R J_0(kR)}{(\alpha^2 + k^2)^{5/2}} - \frac{(R^2 + z'^2) k J_1(kR)}{(\alpha^2 + k^2)^{3/2}} + \frac{9k J_1(kR)}{(\alpha^2 + k^2)^{5/2}} - \frac{15\alpha^2 k J_1(kR)}{(\alpha^2 + k^2)^{7/2}} \right). \end{aligned} \quad (\text{A21})$$

Further simplification is possible if a specific form for  $f(z')$  is specified, with two commonly considered ones being of the form

$$f(z') = \frac{1}{2z_0} \text{sech}^2 \left( \frac{z'}{z_0} \right) \quad , \quad (\text{A22})$$

and

$$f(z') = \frac{1}{\pi z_0} \text{sech} \left( \frac{z'}{z_0} \right) \quad , \quad (\text{A23})$$

with both of them falling off very rapidly once  $z$  is much greater than the scale height  $z_0$ . The thickness function of Eq. (A22) is found to lead to rotational velocities of the form

$$\begin{aligned} RV'_{\beta}(R) &= \frac{N\beta c^2 \alpha^3 R^2}{2} \left[ I_0 \left( \frac{\alpha R}{2} \right) K_0 \left( \frac{\alpha R}{2} \right) - I_1 \left( \frac{\alpha R}{2} \right) K_1 \left( \frac{\alpha R}{2} \right) \right] \\ &- N\beta c^2 \alpha^3 R \int_0^{\infty} dk \frac{k^2 J_1(kR) z_0}{(\alpha^2 + k^2)^{3/2}} \beta \left( 1 + \frac{kz_0}{2} \right) \end{aligned} \quad (\text{A24})$$

and

$$\begin{aligned}
RV'_\gamma(R) &= N\gamma c^2 \alpha^3 R \int_0^\infty dk \left[ 1 - kz_0 \beta \left( 1 + \frac{kz_0}{2} \right) \right] \\
&\times \left( -\frac{2RJ_0(kR)}{(\alpha^2 + k^2)^{3/2}} + \frac{3\alpha^2 RJ_0(kR)}{(\alpha^2 + k^2)^{5/2}} - \frac{R^2 k J_1(kR)}{2(\alpha^2 + k^2)^{3/2}} + \frac{9kJ_1(kR)}{2(\alpha^2 + k^2)^{5/2}} - \frac{15\alpha^2 k J_1(kR)}{2(\alpha^2 + k^2)^{7/2}} \right) \\
&+ N\gamma c^2 \alpha^3 R \int_0^\infty dk \frac{k J_1(kR)}{2(\alpha^2 + k^2)^{3/2}} \frac{d^2}{dk^2} \left[ kz_0 \beta \left( 1 + \frac{kz_0}{2} \right) \right] , \tag{A25}
\end{aligned}$$

where

$$\beta(x) = \int_0^1 dt \frac{t^{x-1}}{(1+t)} . \tag{A26}$$

Similarly, the thickness function of Eq. (A23) leads to

$$RV'_\beta(R) = \frac{2N\beta c^2 \alpha^3 R}{\pi} \int_0^\infty dk \frac{k J_1(kR)}{(\alpha^2 + k^2)^{3/2}} \beta \left( \frac{1 + kz_0}{2} \right) \tag{A27}$$

and

$$\begin{aligned}
RV'_\gamma(R) &= \frac{N\gamma c^2 \alpha^3 R}{\pi} \int_0^\infty dk \beta \left( \frac{1 + kz_0}{2} \right) \\
&\times \left( -\frac{4RJ_0(kR)}{(\alpha^2 + k^2)^{3/2}} + \frac{6\alpha^2 RJ_0(kR)}{(\alpha^2 + k^2)^{5/2}} - \frac{R^2 k J_1(kR)}{(\alpha^2 + k^2)^{3/2}} + \frac{9kJ_1(kR)}{(\alpha^2 + k^2)^{5/2}} - \frac{15\alpha^2 k J_1(kR)}{(\alpha^2 + k^2)^{7/2}} \right) \\
&- \frac{N\gamma c^2 \alpha^3 R}{\pi} \int_0^\infty dk \frac{k J_1(kR)}{(\alpha^2 + k^2)^{3/2}} \frac{d^2}{dk^2} \left[ \beta \left( \frac{1 + kz_0}{2} \right) \right] . \tag{A28}
\end{aligned}$$

In all of these expressions the needed functions of  $\beta(x)$  and its derivatives all converge very rapidly to their asymptotic values as their arguments increase. Consequently the relevant  $k$  integrations all converge very rapidly. As a practical matter, galactic scale heights  $z_0$  are usually much smaller than galactic disk scale lengths  $R_0$ . Consequently, the thickness corrections are only of significance in the inner galactic region, and thus have essentially no effect on the linear potential contribution. For the Newtonian term the thickness corrections of Eqs. (A24) and (A27) both tend to slightly reduce the overall Newtonian contribution, and serve to help ensure that the inner region rotation curves of Fig. 1 are well described by the luminous Newtonian contribution alone.

## C The potential of a spherical bulge

For a spherically symmetric matter distribution such as the central bulge region of a galaxy with radial matter density  $\sigma(r)$  and  $N = 4\pi \int dr' r'^2 \sigma(r')$  stars, the potential is readily found to take the form

$$rV'_\beta(r) = \frac{4\pi\beta c^2}{r} \int_0^r dr' \sigma(r') r'^2 \tag{A29}$$

for a Newtonian potential, and take the form

$$rV'_\gamma(r) = \frac{2\pi\gamma c^2}{3r} \int_0^r dr' \sigma(r') (3r^2 r'^2 - r'^4) + \frac{4\pi\gamma c^2 r^2}{3} \int_r^\infty dr' \sigma(r') r' \tag{A30}$$

for a linear potential. Both of these expressions can readily be integrated once a specific form for  $\sigma(r)$  is specified, and it is from Eq. (A29) that dark matter halo contributions to galactic rotation curves are

calculated. Despite the simplicity of these expressions, because of projection effects, it is unfortunately not the 3-dimensional  $\sigma(r)$  which is directly measured in spherical astronomical systems. Rather, it is only the two-dimensional surface matter distribution  $I(R)$  which is measured, with  $\sigma(r)$  having to be extracted from it via an Abel transform

$$\sigma(r) = -\frac{1}{\pi} \int_r^\infty dR \frac{I'(R)}{(R^2 - r^2)^{1/2}} \quad , \quad I(R) = 2 \int_R^\infty dr \frac{\sigma(r)r}{(r^2 - R^2)^{1/2}} \quad , \quad (\text{A31})$$

to thus initially lead to double integrals in Eqs. (A29) and (A30).

Reduction of these integrals to one-dimensional ones which only involve the measured  $I(R)$  is however possible since on introducing the strip brightness  $S(x)$  which obeys

$$S(x) = 2 \int_x^\infty dR \frac{RI(R)}{(R^2 - x^2)^{1/2}} \quad , \quad \sigma(x) = -\frac{S'(x)}{2\pi x} \quad , \quad (\text{A32})$$

we can rewrite Eqs. (A29) and (A30) as

$$rV'_\beta(r) = -2\beta c^2 S(r) + \frac{2\beta c^2}{r} \int_0^r dr' S(r') \quad (\text{A33})$$

for the Newtonian potential and

$$rV'_\gamma(r) = \gamma c^2 r \int_0^r dr' S(r') - \frac{\gamma c^2}{r} \int_0^r dr' r'^2 S(r') \quad (\text{A34})$$

for the linear one. Then, use of the relations

$$\begin{aligned} \frac{d}{dr} \left\{ 2 \int_r^\infty dR R I(R) \arcsin\left(\frac{r}{R}\right) \right\} &= -\pi r I(r) + S(r) \\ \frac{d}{dr} \left\{ 2 \int_r^\infty dR R I(R) \left[ R^2 \arcsin\left(\frac{r}{R}\right) - r(R^2 - r^2)^{1/2} \right] \right\} &= -\pi r^3 I(r) + 2r^2 S(r) \end{aligned} \quad (\text{A35})$$

enables us to conveniently reexpress Eqs. (A33) and (A34) entirely in terms of  $I(R)$ , to yield

$$rV'_\beta(r) = \frac{4\beta c^2}{r} \int_r^\infty dR R I(R) \left[ \arcsin\left(\frac{r}{R}\right) - \frac{r}{(R^2 - r^2)^{1/2}} \right] \quad (\text{A36})$$

for the Newtonian potential [97] and

$$\begin{aligned} rV'_\gamma(r) &= \frac{2\pi\beta c^2}{r} \int_0^r dR R I(R) + \frac{\gamma c^2 \pi}{2r} \int_0^r dR R I(R) (2r^2 - R^2) \\ &\quad + \frac{\gamma c^2}{r} \int_r^\infty dR R I(R) \left[ (2r^2 - R^2) \arcsin\left(\frac{r}{R}\right) + r(R^2 - r^2)^{1/2} \right] \end{aligned} \quad (\text{A37})$$

for the linear one.

## D The ghost problem in fourth order theories

With the fourth order propagator

$$D(k^2, M^2) = \frac{1}{(k_0^2 - \vec{k}^2)(k_0^2 - \vec{k}^2 - M^2)} \quad (\text{A38})$$

associated with the prototype action

$$I = -\frac{1}{2} \int d^4x \left( M^2 \partial_\mu S \partial^\mu S + \partial_\mu \partial_\nu S \partial^\mu \partial^\nu S \right) \quad (\text{A39})$$

and equation of motion

$$(-\partial_0^2 + \nabla^2)(-\partial_0^2 + \nabla^2 - M^2)S = 0 \quad (\text{A40})$$

being writable as a sum of two opposite signated second order propagators, viz.

$$D(k^2, M^2) = \frac{1}{M^2(k_0^2 - \vec{k}^2 - M^2)} - \frac{1}{M^2(k_0^2 - \vec{k}^2)} \quad (\text{A41})$$

it is quite widely thought that when quantized, fourth order theories such as conformal gravity would possess ghost states. However, the equation of motion which is given in Eq. (A40) is not actually the equation of motion of a pure fourth order theory per se, rather it is that of a second plus fourth order theory. With the pure fourth order action being the one obtained by setting  $M^2$  to zero in Eq. (A39), and with the separate positive and negative signated second order propagators in Eq. (A41) becoming singular in this limit, no conclusion about the particle content of the pure fourth order theory can be immediately drawn.

To investigate what the particle content of the fourth order theory does look like, it is sufficient to specialize to field configurations of the form  $S(\bar{x}, t) = q(t)e^{i\vec{k}\cdot\bar{x}}$ , configurations in which Eq. (A40) then reduces to

$$\frac{d^4q}{dt^4} + (\omega_1^2 + \omega_2^2) \frac{d^2q}{dt^2} + \omega_1^2 \omega_2^2 q = 0 \quad (\text{A42})$$

where

$$\omega_1^2 + \omega_2^2 = 2\vec{k}^2 + M^2 \quad , \quad \omega_1^2 \omega_2^2 = \vec{k}^4 + \vec{k}^2 M^2 \quad (\text{A43})$$

with Eq. (A42) itself reducing to the equal frequency  $\omega_1 = \omega_2$  when  $M^2 = 0$ . The equation of motion given in Eq. (A42) can be derived by variation of the acceleration-dependent action first introduced by Pais and Uhlenbeck [98]

$$I_{PU} = \frac{\gamma}{2} \int dt \left[ \dot{q}^2 - (\omega_1^2 + \omega_2^2) q^2 + \omega_1^2 \omega_2^2 q^2 \right] \quad (\text{A44})$$

where  $\gamma$  is a constant, and we thus seek to quantize the theory based on  $I_{PU}$  to see what happens when we take the  $\omega_1 \rightarrow \omega_2$  limit.

For the theory associated with  $I_{PU}$  considered first as a classical theory, we cannot treat  $q$  and  $\dot{q}$  as independent coordinates, since if  $\dot{q}$  is to be an independent coordinate with canonical conjugate  $\partial L / \partial \dot{q}$ , we could not then use  $\partial L / \partial \dot{q}$  as the canonical conjugate of  $q$ . In order to reexpress this theory in terms of an unconstrained set of variables, we introduce a new variable  $x$  to replace  $\dot{q}$ , with the method of Dirac constraints then converting the classical theory associated with  $I_{PU}$  into a theory of two coupled classical oscillators  $(q, p_q)$  and  $(x, p_x)$ , with the unconstrained classical Hamiltonian

$$H = \frac{p_x^2}{2\gamma} + p_q x + \frac{\gamma}{2} (\omega_1^2 + \omega_2^2) x^2 - \frac{\gamma}{2} \omega_1^2 \omega_2^2 q^2 \quad (\text{A45})$$

then resulting [87]. With this Hamiltonian generating the closed set of Poisson bracket relations

$$\begin{aligned} \{x, p_x\} &= 1 \quad , \quad \{q, p_q\} = 1 \quad , \\ \{x, H\} &= \frac{p_x}{\gamma} \quad , \quad \{q, H\} = x \quad , \quad \{p_x, H\} = -p_q - \gamma(\omega_1^2 + \omega_2^2)x \quad , \quad \{p_q, H\} = \gamma\omega_1^2\omega_2^2 q \end{aligned} \quad (\text{A46})$$

for Poisson brackets defined by

$$\{A, B\} = \frac{\partial A}{\partial x} \frac{\partial B}{\partial p_x} - \frac{\partial A}{\partial p_x} \frac{\partial B}{\partial x} + \frac{\partial A}{\partial q} \frac{\partial B}{\partial p_q} - \frac{\partial A}{\partial p_q} \frac{\partial B}{\partial q} \quad (\text{A47})$$

this Hamiltonian is indeed the correct classical one, and thus the appropriate one for quantization.

At the classical level, Eq. (A46) yields equations of motion of the form

$$\dot{x} = \frac{p_x}{\gamma} \quad , \quad \dot{q} = x \quad , \quad \dot{p}_x = -p_q - \gamma(\omega_1^2 + \omega_2^2)x \quad , \quad \dot{p}_q = \gamma\omega_1^2\omega_2^2q \quad , \quad (\text{A48})$$

to nicely recover Eq. (A42), with the substitution of this solution into the Hamiltonian of Eq. (A45) yielding a stationary Hamiltonian

$$H_{\text{STAT}} = \frac{\gamma}{2}\dot{q}^2 - \frac{\gamma}{2}(\omega_1^2 + \omega_2^2)q^2 - \frac{\gamma}{2}\omega_1^2\omega_2^2q^2 - \gamma q \frac{d^3q}{dt^3} \quad (\text{A49})$$

whose time independence in solutions to Eq. (A42) can readily be confirmed. The unequal frequency theory has explicit two-oscillator solution

$$q(t) = a_1 e^{-i\omega_1 t} + a_2 e^{-i\omega_2 t} + \text{c.c.} \quad (\text{A50})$$

with energy

$$H_{\text{STAT}}(\omega_1 \neq \omega_2) = 2\gamma(\omega_1^2 - \omega_2^2)(a_1^* a_1 \omega_1^2 - a_2^* a_2 \omega_2^2) \quad , \quad (\text{A51})$$

while the  $\omega_1 = \omega_2 = \omega$  equal frequency theory has explicit solution

$$q(t) = c_1 e^{-i\omega t} + c_2 t e^{-i\omega t} + \text{c.c.} \quad (\text{A52})$$

with energy

$$H_{\text{STAT}}(\omega_1 = \omega_2) = 4\gamma\omega^2 (2c_2^* c_2 + i\omega c_1^* c_2 - i\omega c_2^* c_1) \quad . \quad (\text{A53})$$

As we see, despite the fact that the equal frequency solution involves a temporal runaway, the appropriately defined energy is still time independent, with temporal runaways not being the problem for fourth order theories they are often thought to be.

Before proceeding to quantize the fourth order theory, we note that a comparison of Eqs. (A51) and (A53) reveals a difference between the unequal and equal frequency theories already at the classical level, one that will prove to be of great significance in the following. Specifically, while the unequal frequency theory is diagonal in the  $(a_1, a_2)$  basis, with one of the oscillators having negative energy (the disease which translates into negative norm ghost states in the quantum theory), the equal frequency theory is not diagonal in the  $(c_1, c_2)$  basis. Moreover, if we set  $c_2 = 0$  we obtain

$$H_{\text{STAT}}(\omega_1 = \omega_2) = 0 \quad , \quad (\text{A54})$$

while if we set  $c_1 = 0$  we obtain

$$H_{\text{STAT}}(\omega_1 = \omega_2) = 8\gamma\omega^2 c_2^* c_2 \quad . \quad (\text{A55})$$

Of these two modes, we see that only the  $c_2$  one carries energy. The zero energy result of Eq. (A54) is reminiscent of the conformal gravity zero energy theorem, with Boulware, Horowitz and Strominger [99] having shown that at the classical level the fourth order gravity theory energy would vanish identically if the gravitational field solutions were required to be asymptotically flat. With the restriction to asymptotic flatness in the gravitational case being equivalent in the Pais-Uhlenbeck case to requiring no temporal runaway, the equivalence of Eq. (A54) to the zero energy theorem of [99] is manifest. However, in our analysis of the conformal gravity analog of the Schwarzschild solution, we found that conformal gravity also possessed the non-asymptotically flat linear potential term of Eq. (204). We now see from Eq. (A55) that when the asymptotic flatness requirement is dropped, the temporal runaway solution would then be allowed, and would (for an appropriate choice of the sign of  $\gamma$ ) actually have a perfectly acceptable positive energy, with the constraints of the zero energy theorem (a theorem long considered to be a shortcoming of conformal gravity) then being evaded. With this in mind, we shall

now show that the quantization of the equal frequency Pais-Uhlenbeck theory will lead to one oscillator which propagates with positive energy and to a second oscillator which does not propagate at all, with the would be ghost mode not being an eigenstate of the Hamiltonian.

To quantize the theory based on the Hamiltonian of Eq. (A45), the introduction of the four Fock space operators  $a_1$ ,  $a_1^\dagger$ ,  $a_2$  and  $a_2^\dagger$  defined via

$$\begin{aligned} q(t) &= a_1 e^{-i\omega_1 t} + a_2 e^{-i\omega_2 t} + \text{H.c.} \quad , \quad p_q(t) = i\gamma\omega_1\omega_2^2 a_1 e^{-i\omega_1 t} + i\gamma\omega_1^2\omega_2 a_2 e^{-i\omega_2 t} + \text{H.c.} \quad , \\ x(t) &= -i\omega_1 a_1 e^{-i\omega_1 t} - i\omega_2 a_2 e^{-i\omega_2 t} + \text{H.c.} \quad , \quad p_x(t) = -\gamma\omega_1^2 a_1 e^{-i\omega_1 t} - \gamma\omega_2^2 a_2 e^{-i\omega_2 t} + \text{H.c.} \end{aligned} \quad (\text{A56})$$

then furnishes us with a Fock space representation of the quantum-mechanical commutation relations

$$[x, p_x] = [q, p_q] = i \quad , \quad [x, q] = [x, p_q] = [q, p_x] = [p_x, p_q] = 0 \quad (\text{A57})$$

provided the Fock space operators obey

$$[a_1, a_1^\dagger] = \frac{1}{2\gamma\omega_1(\omega_1^2 - \omega_2^2)} \quad , \quad [a_2, a_2^\dagger] = \frac{1}{2\gamma\omega_2(\omega_2^2 - \omega_1^2)} \quad , \quad [a_1, a_2^\dagger] = 0 \quad , \quad [a_1, a_2] = 0 \quad . \quad (\text{A58})$$

In terms of the Fock space operators the quantum-mechanical Hamiltonian takes the form

$$H = 2\gamma(\omega_1^2 - \omega_2^2)(\omega_1^2 a_1^\dagger a_1 - \omega_2^2 a_2^\dagger a_2) + \frac{1}{2}(\omega_1 + \omega_2) \quad (\text{A59})$$

with its associated commutators as inferred from Eq. (A46) then automatically being satisfied. With  $\omega_1$ ,  $\omega_2$  and  $\gamma(\omega_1^2 - \omega_2^2)$  all being taken to be positive for definitiveness, we see that the  $[a_1, a_1^\dagger]$  commutator is positive definite while the  $[a_2, a_2^\dagger]$  commutator is negative definite. And with the Hamiltonian being diagonal in the  $a_1^\dagger a_1$ ,  $a_2^\dagger a_2$  occupation number operator basis, we see that the state defined by

$$a_1 |\Omega\rangle = 0 \quad , \quad a_2 |\Omega\rangle = 0 \quad (\text{A60})$$

is its ground state, that the one-particle states

$$|+1\rangle = [2\gamma\omega_1(\omega_1^2 - \omega_2^2)]^{1/2} a_1^\dagger |\Omega\rangle \quad , \quad |-1\rangle = [2\gamma\omega_2(\omega_1^2 - \omega_2^2)]^{1/2} a_2^\dagger |\Omega\rangle \quad (\text{A61})$$

are both positive energy eigenstates with respective energies  $\omega_1$  and  $\omega_2$  above the ground state, that the state  $|+1\rangle$  has a norm equal to plus one, but that the state  $|-1\rangle$  has norm minus one, a ghost state. Thus, as anticipated from Eq. (A41), we find that the Hamiltonian of the unequal frequency theory can indeed be diagonalized in a basis of positive and negative norm states.<sup>100</sup> In the presence of interactions that would be added on to the Pais-Uhlenbeck action  $I_{PU}$ , the asymptotic states for the S-matrix that such interactions would then generate would be the eigenstates of the free Hamiltonian of Eq. (A59). While there are negative norm asymptotic states in the eigenspectrum of the free Hamiltonian, in and of itself that would only be a disaster if interactions would connect in and out states of opposite signature, and it is thus of interest to note that Hawking and Hertog [100] have argued via a path integral treatment of the constraints of  $I_{PU}$  that this does not in fact occur, with the unequal frequency theory then being viable. Hence even when  $M^2$  is non-zero, assessing fourth order theories on the basis of the structure of Eq. (A41) is too hasty, and a careful quantization of the theory which takes constraints into account first needs to be made.

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<sup>100</sup>Despite the presence of ghost states, the propagator of the theory is still causal, with the relative minus sign in Eq. (A41) not affecting the fact that each of the two second order propagators which appear in  $D(k^2, M^2)$  is a standard causal second order propagator.

With the commutation relations of Eq. (A58) being singular in the  $\omega_1 \rightarrow \omega_2$  limit, we see that we cannot infer anything about the structure of the equal frequency Fock space from a study of the unequal frequency one. To take the limit we need to find some new set of Fock operators whose commutation relations would instead be non-singular. To this end we thus introduce the basis

$$\begin{aligned} a_1 &= \frac{1}{2} \left( a - b + \frac{2b\omega}{\epsilon} \right) , & a_2 &= \frac{1}{2} \left( a - b - \frac{2b\omega}{\epsilon} \right) , \\ a &= a_1 \left( 1 + \frac{\epsilon}{2\omega} \right) + a_2 \left( 1 - \frac{\epsilon}{2\omega} \right) , & b &= \frac{\epsilon}{2\omega} (a_1 - a_2) . \end{aligned} \quad (\text{A62})$$

for the unequal frequency theory where we have set

$$\omega = \frac{(\omega_1 + \omega_2)}{2} , \quad \epsilon = \frac{(\omega_1 - \omega_2)}{2} . \quad (\text{A63})$$

These new variables are found to obey commutation relations of the form

$$[a, a^\dagger] = \lambda , \quad [a, b^\dagger] = \mu , \quad [b, a^\dagger] = \mu , \quad [b, b^\dagger] = \nu , \quad [a, b] = 0 , \quad (\text{A64})$$

where

$$\lambda = \nu = -\frac{\epsilon^2}{16\gamma(\omega^2 - \epsilon^2)\omega^3} , \quad \mu = \frac{(2\omega^2 - \epsilon^2)}{16\gamma(\omega^2 - \epsilon^2)\omega^3} . \quad (\text{A65})$$

In terms of these new variables the coordinate  $q(t)$  of Eq. (A56) gets rewritten as

$$q(t) = e^{-i\omega t} \left[ (a - b) \cos \epsilon t - \frac{2ib\omega}{\epsilon} \sin \epsilon t \right] + \text{H.c.} , \quad (\text{A66})$$

and thus has a well defined  $\epsilon \rightarrow 0$  limit, viz.

$$q(t, \epsilon = 0) = e^{-i\omega t} (a - b - 2ib\omega t) + \text{H.c.} . \quad (\text{A67})$$

Similarly, in the new variables the Hamiltonian of Eq. (A59) gets rewritten as

$$H = 8\gamma\omega^2\epsilon^2(a^\dagger a - b^\dagger b) + 8\gamma\omega^4(2b^\dagger b + a^\dagger b + b^\dagger a) + \omega , \quad (\text{A68})$$

and it too has a well-defined limit, viz.

$$H(\epsilon = 0) = 8\gamma\omega^4(2b^\dagger b + a^\dagger b + b^\dagger a) + \omega , \quad (\text{A69})$$

with the following commutators of interest having limiting form

$$\begin{aligned} [H(\epsilon = 0), a^\dagger] &= \omega(a^\dagger + 2b^\dagger) , & [H(\epsilon = 0), a] &= -\omega(a + 2b) , \\ [H(\epsilon = 0), b^\dagger] &= \omega b^\dagger , & [H(\epsilon = 0), b] &= -\omega b , \\ [a + b, a^\dagger + b^\dagger] &= 2\hat{\mu} , & [a - b, a^\dagger - b^\dagger] &= -2\hat{\mu} , & [a + b, a^\dagger - b^\dagger] &= 0 , \end{aligned} \quad (\text{A70})$$

where  $\hat{\mu} = \mu(\epsilon = 0) = 1/(8\gamma\omega^3)$ .

With the equal frequency theory now being well-defined, we find that for the Fock vacuum  $|\Omega\rangle$  defined by  $a|\Omega\rangle = b|\Omega\rangle = 0$ ,  $H(\epsilon = 0)|\Omega\rangle = \omega|\Omega\rangle$ , the  $H(\epsilon = 0)$  Hamiltonian possesses only one and not two one-particle states. The state  $b^\dagger|\Omega\rangle$  is an eigenstate with expressly positive energy  $2\omega$ , while the state  $a^\dagger|\Omega\rangle$  is not an eigenstate at all, a quantum-mechanical one-particle spectrum which thus reflects the features of the classical spectrum exhibited in Eqs. (A54) and (A55). With this same pattern repeating for the multi-particle states, the equal frequency theory is thus a very unusual one in which the full Fock space has the same number of basis states as a two-dimensional harmonic oscillator, while

the Hamiltonian itself only has the number of basis states associated with a one-dimensional one. The  $a^\dagger|\Omega\rangle$  states thus only couple off shell (where they can serve to regulate the ultraviolet behavior of theory) but do not materialize as on-shell asymptotic states. While the disappearance of these modes from the eigenspectrum of the Hamiltonian solves the ghost problem in the theory, this disappearance is still quite perplexing and requires further explanation.

The reason for this highly unusual outcome derives from the fact that while the normal situation for square matrices is that the number of independent eigenvectors of a square matrix is the same as the dimensionality of the matrix, there are certain matrices, known as defective matrices, for which this is not in fact the case. A typical example of such a defective matrix is the non-Hermitian two-dimensional matrix

$$M = \begin{pmatrix} 1 & c \\ 0 & 1 \end{pmatrix} \quad (\text{A71})$$

with  $c$  non-zero but otherwise arbitrary, since even though this matrix has two eigenvalues both of which are real (despite the lack of hermiticity) and equal to 1 (no matter what the value of  $c$ ), solving the equation

$$\begin{pmatrix} 1 & c \\ 0 & 1 \end{pmatrix} \begin{pmatrix} p \\ q \end{pmatrix} = \begin{pmatrix} p + cq \\ q \end{pmatrix} = \begin{pmatrix} p \\ q \end{pmatrix} \quad (\text{A72})$$

leads only to  $q = 0$  when  $c \neq 0$ , and thus to only one eigenvector despite the two-fold degeneracy of the eigenvalue, with the space on which the matrix  $M$  acts not being complete. The defective matrix  $M$  given above is in the form of a Jordan block matrix, with Jordan having shown that under a similarity transform an arbitrary square matrix can be brought to either a diagonal form, or to a triangular form such as that of Eq. (A71) in which every matrix element on one side of the diagonal is zero. Jordan block form matrices have the property that no matter what values the matrix elements take on the other side of the diagonal, the secular equation for the eigenvalues only involves the elements on the diagonal itself. Then if these diagonal elements are all real, the matrix will have real eigenvalues despite not being Hermitian. (While Hermitian matrices must have real eigenvalues, there is no converse theorem which would oblige the eigenvalues of non-Hermitian matrices to necessarily be complex, with non-Hermitian matrices still being able to possess real eigenvalues in certain cases.) As such, a Jordan block form matrix can be thought of as being a diagonal matrix  $M_1$  to which has been added a second matrix  $M_2$  with non-zero elements on only one side of the diagonal. With this second matrix being a divisor of zero, while its addition to the diagonal matrix does not affect eigenvalues, its being a divisor of zero does affect eigenvectors, to thereby cause the full Jordan block matrix  $M_1 + M_2$  to have a smaller number of eigenvectors than eigenvalues.

To see how these considerations apply in the case of interest to us here, we need to track the  $\epsilon \rightarrow 0$  limit carefully. As regards first the unequal frequency Hamiltonian of Eq. (A68), we note that its action on the one-particle states  $a^\dagger|\Omega\rangle$ ,  $b^\dagger|\Omega\rangle$  yields

$$\begin{aligned} Ha^\dagger|\Omega\rangle &= \frac{1}{2\omega} \left[ (4\omega^2 + \epsilon^2)a^\dagger|\Omega\rangle + (4\omega^2 - \epsilon^2)b^\dagger|\Omega\rangle \right] \\ Hb^\dagger|\Omega\rangle &= \frac{1}{2\omega} \left[ \epsilon^2 a^\dagger|\Omega\rangle + (4\omega^2 - \epsilon^2)b^\dagger|\Omega\rangle \right] . \end{aligned} \quad (\text{A73})$$

In this sector we can define a matrix

$$M(\epsilon) = \frac{1}{2\omega} \begin{pmatrix} 4\omega^2 + \epsilon^2 & 4\omega^2 - \epsilon^2 \\ \epsilon^2 & 4\omega^2 - \epsilon^2 \end{pmatrix} \quad (\text{A74})$$

whose eigenvalues are given as  $2\omega + \epsilon$  and  $2\omega - \epsilon$ . For such eigenvalues, energy eigenvectors which obey

$$H|2\omega \pm \epsilon\rangle = (2\omega \pm \epsilon)|2\omega \pm \epsilon\rangle \quad (\text{A75})$$

are then readily constructed as

$$|2\omega \pm \epsilon\rangle = [\pm\epsilon a^\dagger + (2\omega \mp \epsilon)b^\dagger] |\Omega\rangle . \quad (\text{A76})$$

As we see, as long as  $\epsilon \neq 0$ , the two one-particle sector eigenvectors of the Hamiltonian  $H$  are distinct, and  $H$  has two eigenvectors to go with its two eigenvalues. However when we now let  $\epsilon$  go to zero, the two eigenvectors in Eq. (A76) collapse onto a single eigenvector, viz. the vector  $b^\dagger|\Omega\rangle$ , the two eigenvalues collapse onto a common eigenvalue, viz.  $2\omega$ , and the matrix  $M(\epsilon = 0)$  of Eq. (A74) becomes defective, with the  $H(\epsilon = 0)$  Hamiltonian then being free of negative norm eigenstates. The equal frequency theory is thus fully acceptable, to thereby resolve what had been thought of as being one of the biggest difficulties for fourth order theories.

As regards the conformal gravity theory itself, we note that in a  $g_{\mu\nu} = \eta_{\mu\nu} + h_{\mu\nu}$  linearization of the theory around flat spacetime, the conformal gravity rank two tensor of Eq. (185) reduces to

$$W^{\mu\nu} = \frac{1}{2}\Pi^{\mu\rho}\Pi^{\nu\sigma}K_{\rho\sigma} - \frac{1}{6}\Pi^{\mu\nu}\Pi^{\rho\sigma}K_{\rho\sigma} \quad (\text{A77})$$

where

$$K^{\mu\nu} = h^{\mu\nu} - \frac{1}{4}\eta^{\mu\nu}h^\alpha{}_\alpha , \quad \Pi^{\mu\nu} = \eta^{\mu\nu}\partial^\alpha\partial_\alpha - \partial^\mu\partial^\nu . \quad (\text{A78})$$

In the conformal gauge

$$\partial_\nu g^{\mu\nu} - \frac{1}{4}g^{\mu\sigma}g_{\nu\rho}\partial_\sigma g^{\nu\rho} = 0 \quad (\text{A79})$$

(viz. the gauge condition which is left invariant under  $g_{\mu\nu}(x) \rightarrow e^{2\alpha(x)}g_{\mu\nu}(x)$ ) the source-free region gravitational fluctuation equation  $W^{\mu\nu} = 0$  then reduces to

$$(-\partial_0^2 + \nabla^2)^2 K^{\mu\nu} = 0 . \quad (\text{A80})$$

With this equation of motion being decoupled in its tensor indices, we see that each tensor component precisely obeys none other than the  $M^2 = 0$  limit of Eq. (A40). Consequently, it would be of interest to see if the structure we have found for the prototype equal frequency Pais-Uhlenbeck theory can carry over to the conformal gravity theory once its gauge and tensor structure is taken into consideration, since that would then permit the construction of a fully renormalizable, fully unitary gravitational theory in four spacetime dimensions, one which despite its fourth order equation of motion, would nonetheless only possess one on-shell graviton and not two.<sup>101</sup>

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<sup>101</sup>It would also be of interest to see whether the typical linear in time growth given in Eq. (A52) might enable conformal gravity early universe tensor fluctuations to grow to macroscopic size by recombination.